Brightest Stars, expanded version

by Toomas Karmo

PREFACE: Scope and purpose of this essay

This online essay is an extended version of the essay in the printed-edition Handbook, containing all the material of its printed-edition accompaniment, but adding material of its own. The accompanying online table is likewise an extended version of the printed-edition table, *(a)* with extra stars (after providing for multiplicity, as we explain below, the brightest MK-classified 323, allowing for variability, where the printed edition has almost 30 fewer, allowing for variability: our cutoff is mag. ~3.55), and *(b)* with additional remarks for most of the duplicated stars. We use a dagger superscript (†) to mark data cells for which the online table supplies some additional information, some context, or a caveat.

The online essay and table try to address the needs of three kinds of serious amateur: amateurs who are also astrophysics students (whether or not enrolled formally at some campus); amateurs who, like many in the RASC, assist in public outreach, through some form of lecturing; and amateurs who are planning their own private citizen-science observing runs, in the spirit of such "pro-am" organizations as AAVSO. Additionally, we would hope that the online project will help serve a constituency of sky-lovers, whether professional or amateur, who work with the heavens in an unambitious and contemplative spirit, seeking to understand at the eyepiece, or even with the naked eye, the realities behind the little that their limited circumstances may allow them to see. (This is the same contemplative exercise as is proposed for the Cyg X-1 black hole, with its gas-dumping supergiant companion HD226868, in the Handbook "Expired Stars" essay: with a small telescope, or even with binoculars, we first find HD226868, and then take a moment to ponder in awe the accompanying unobserved realities of gas-fed hot accretion disk, event horizon, and spacetime singularity.)

Our online project, started as a supplement to the 2017 Handbook, must be considered still in its rather early stages. We cannot claim to have fully satisfied the needs of our various constituencies. Above all, we cannot claim to have covered all the appropriate points from stellar-astronomy news in our "Remarks" column, important though news is to amateurs of all three types. We would hope in coming years to remedy our deficiencies in several ways, most notably by relying more in our writing on recent primary-literature journal articles, with appropriate explicit citations.

In our citations, we favour the now-preferred astrophysics "bibcode" formalism. The formalism is documented in simbad.u-strasbg.fr/guide/refcode/refcode-paper.html, and again in section 1.2.3 (headed "Bibliographic Identifiers") in adsabs.harvard.edu/abs_doc/help_pages/data.html.

A bibcode can be transformed into the display of a more human-readable bibliography entry, often with clickable hyperlink to an underlying online full-text, all-illustrations PDF publication, in various ways. We illustrate some possibilities by taking an extreme case, namely our bibcode reference to the classic 1910 Joel Stibbins *Astrophysical Journal* paper that reports the electric-photometry discovery of a secondary minimum in the Algol light curve. Old though the paper is, it is nevertheless available online. The bibcode (as we state again in our "Remarks" for the Algol entry in our table) is 1910ApJ....32..185S. A browser display with hyperlink to the desired full-text, all-illustrations PDF is available from the Centre de Données Stellaires (CDS) server (probably in Strasbourg) as simbad.u-strasbg.fr/simbad/sim-ref?bibcode=1910ApJ....32..185S. If something has gone wrong—and experience suggests that things can go wrong, even when a bibcode appears to casual inspection to be correctly typed, at any rate in some such autonomous-agent computing environment as Microsoft Office—then one can recover through CDS as simbad.u-strasbg.fr/simbad/sim-fid if the star of interest and year of publication are known. In this particular case, recovery involves giving simbad.u-strasbg.fr/simbad/sim-fid some convenient identifier, for instance the IAU-promulgated name "Algol" or the Bayer identifier "beta Per." In the Algol-specific input form generated, one next asks, in the "References" section of the form, for all references from 1910 to 1910. The duly displayed bibcode, 1910ApJ....32..185S, for the sole 1910-through-1910 reference, is shown as a clickable hyperlink. Upon further clicking, the hyperlink eventually yields the PDF. A similar browser display is available from a (probably North American) ADS-NASA server as ui.adsabs.harvard.edu/abs/1910ApJ....32..185S/abstract. As a fourth possibility, the PDF is retrievable through a self-evident set of steps that starts by copying and pasting the bibcode into the "Bibliographic Code Query" box at the paper-workflow, as distinct from the more obviously accessible paperless-workflow, online form ui.adsabs.harvard.edu/paper-form. This fourth method has the advantage that multiple bibcodes can be entered within a single query. As a fifth possibility, which in our view cannot be guaranteed to work (but there seem to be intermittent problems with the fourth possibility as well; and in general, servers should not be presumed fully reliable, in any discipline) is simply to put 1910ApJ....32..185S naively into a general Google search, and to explore the ensuing chain of hyperlinks: in the case of at least a heavily cited paper, one is likely soon enough to reach an abstract at ADS-NASA or some similar authority, with accompanying PDF.

The bibliographic support of simbad.u-strasbg.fr/simbad and ui.adsabs.harvard.edu, as the principal tools for our primary-literature searching, is herewith gratefully acknowledged, as are *Wikipedia* (in exact-science topics, generally careful and up to date); *Sky & Telescope*, the web materials of Prof. James Kaler, and at a more technical level, two key sources of data, the Washington Double Star Catalog (WDS) and AAVSO. Helpful at AAVSO are not only the general graphing facility and the general AAVSO record of observations, but also a more recent offering, the VSX online database.

SECTION 1: Selection bases for our 315 nominal "bright stars," strictly 323 MK-classified bright stars

Of our selected 315 nominal stars, three call for extra comment pertinent to this mag. ~3.55 naked-eye selection criterion. (1) κ (kappa) CMa (at RA ~6h50) brightened in the 1960s or 1970s, just managing to meet the cutoff, and has remained bright. This change was unfortunately not noted in the RASC Handbook until 2019. Should κ CMa now once again fade, we propose to keep listing it for at least a few years, since it is a variable of the γ Cas type (and may therefore be liable to yet further episodes of brightening during the 21st century; in general, the γ Cas variables, whether temporarily bright or temporarily faint, are desirable targets for ongoing, regular, citizen-science spectroscopy, and even naked-eye monitoring, being closely associated with the amateur-relevant "Be phenomenon," which we discuss in the final subsection of this essay). (2) We discontinued listing L_2 Pup (at RA ~07h14) in the 2017, 2018, and 2019 Handbooks. Now, however, we revert to our pre-2017 policy, since L_2 Pup is a semi-regular pulsator, occasionally bright. (3) T CrB (at RA ~16h00) has shown nova behaviour, brightening from its current very faint state (mag. ~10) to mag. 2.0 in 1866 and to mag. 3.0 in 1946. We have for years listed this star in our table and propose to continue listing it (since the history suggests the possibility of a 21st-century outburst).

An omission from our selection of 315 nominal stars also calls for comment: while mindful of the fact that η Car brightened greatly, attaining even mag. 0 for a few years from 1837 onward, we omit it from our table since there is no firm prognosis of a 21st-century repetition of that outburst. In February 2022, η Car was being reported as at mag. 4.1 in the V band, and in January 2021 was being variously reported visually as at mag. 4.3 and 4.5, in all cases being decidedly fainter than our naked-eye cutoff. The star was not reported by AAVSO visual observers as any brighter than mag. 3.9 in the second half of 2020.

We may now explain in what sense the set of 315 is nominal. In a strict accounting, the selection is a set of 315 objects that are in naked-eye terms "bright stars," i.e. are bright, unresolved point sources of starlight. Three kinds of situation need to be distinguished here, as we move from naked-eye impressions to underlying physical realities:

(1) In some cases (to cite an example at random, β Tau (Elnath)) what is to the naked eye a point source actually is, so far is as known, a solitary star.

(2) Very common is a situation in which a bright star is a component in a multi-star system, with the other member(s) making either a very small or a negligible contribution to the naked-eye retinal signal. An instance of the former type of pairing is γ And A, from which at a distance of just over 9″ lie two fainter stars, γ And B and γ And C, themselves separated by a mere 0.2″, and so faint that the BC pairing shines at around mag. 5. This has the consequence that BC makes just a modest contribution to the overall γ And ABC naked-eye neuron response. An instance of the latter type of pairing is α CMa A (Sirius), with α CMa B a white dwarf shining at mag. 8.5, in other words shining so feebly as to play essentially no role in the signal generated by the naked-eye retina. This binary constitutes a not trivial, and yet also at the present favourable time a not hopeless, project for the small telescope. (At www.rasc.ca/sirius-observing-challenge, RASC notes that with apastron due in 2025, "an extremely difficult feat has become merely a very demanding one.") Since our table is officially a table of bright stars, we take care, at any rate in our various table revisions from early 2021 onward, to write in our first table column "γ And A" (not "γ And AB" or "γ And"), and "α CMa A" (not "α CMa AB" or "α CMa"). Helpfully, the naming rules promulgated since around 2016 at IAU, and reflected in our concluding "Remarks" column, stipulate, in parallel with our first-column decision, that a name such as "Sirius" applies to a star such as α CMa A, rather than to the binary system α CMa AB.

(3) In eleven other cases, the naked-eye point, shining at mag. ~3.55 or brighter, is the combined light of two binary-system components, each individually bright enough to count as a "bright star"—perhaps with each component exceeding our mag. ~3.55 cutoff, but also perhaps with one or both components just a little fainter than our mag. ~3.55 cutoff, and yielding a "star" brighter than mag. ~3.55 upon combining the light.

These eleven, so-to-speak awkward, cases (awkwardly forcing us to write the binary designations "AB" or "Aa,Ab" in the first column) are the following:

- β Phe AB (with each of A, B individually around mag. 4, yielding an aggregated naked-eye impression of mag. 3.2)
- γ Per Aa,Ab (with each of Aa, Ab a little brighter than mag. 4, yielding an aggregated naked-eye impression of mag. 2.91)
- α Aur Aa (Capella), Ab (with each of Aa, Ab very close to mag. 0)
- β Aur Aa (Menkalinam), Ab (with magnitudes nearly equal, yielding an aggregated naked-eye impression a little brighter than mag. 2)
- ζ CMa Aa,Ab (magnitudes nearly equal, at 3.6 and 3.8; this very tight binary is not as yet well observed, with as of at any rate 2022 March 2 just 4 WDS-documented measurements, from 2019 and 2020)
- γ Vir A (Porrima), B (magnitudes nearly equal, and with each individually very close to our mag. ~3.55 cutoff, yielding an aggregated naked-eye impression a little brighter than mag. 3)
- β Cen Aa (Hadar), Ab (magnitudes nearly equal, with each individual star much brighter than our mag. ~3.55 cutoff)
- η Oph A (Sabik), B (with B at mag. 3.5)
- λ Sco Aa (Shaula), Ab (with even Ab well above our cutoff, at mag. ~2.8)
- ζ Sgr A (Ascella), B (with B at mag. 3.5)
- π Sgr A (Albaldah), B (a poorly documented pairing, with the faint outlier C also poorly documented: WDS implies that B is of nearly the same magnitude as A, with each of these two stars very close to our mag. ~3.55 cutoff)

It is tempting to consider the η Peg system to be a twelfth case, requiring entry as "η Peg Aa (Matar), Ab". But since Ab is decidedly fainter than our mag. ~3.55 cutoff, and Aa only slightly fainter than our mag. ~3.55 cutoff, we are obliged instead to enter this case simply as "η Peg Aa (Matar)", drawing attention in the table "Remarks" column to the fact that our stated magnitude of 2.93 is for the combined light. Somewhat like the η

Peg system is the o (omicron) Leo system, where o Leo Aa (Subra) is very close to mag. 3.5, and where o Leo Ab is, while fainter than mag. 3.5, nevertheless bright enough to make a non-trivial contribution to the overall visual impression. Before 2021, our table unfortunately had the erroneous information that o Leo Aa,Ab is a binary system in which the components are of equal magnitude. Also somewhat like the η Peg system is θ Tau Aa, shining a bit below our magnitude cutoff at 3.74, and rising to just above the cutoff in the combined light of θ Tau Aa,Ab (where Ab is well below our cutoff, at mag. 4.86).

We thus have a table of nominally 315 stars, comprising a more refined, i.e less nominal, analysis 315 + 11 = 326 bright stars. With 3 exceptions, each of the 326 has a known (at worst, an uncertainty-flagged) MK temperature type and MK luminosity class (with the Sun, of course, better observed than any of the others). (The case of ζ CMa Aa,Ab is admittedly rather indeterminate. Here we have a longstanding MK type, from decades before the 2019 discovery of binarity. Should it be assumed now that the type applies only to ζ CMa Aa, or should it be understood as approximately correct also for ζ CMa Ab? Since the magnitude difference is small, around 0.2 mag., and since the stars are likely to be of the same age, as products of co-genesis from an ISM molecular cloud, we take the latter option.) The final result is accordingly a set of 326 –3 = 323 bright stars of known MK classification.

SECTION 2: General characteristics of our 323 MK-classified bright stars

Our 323-element sample is found to lie in a region, around 3000 ly in radius, essentially confined to the sandwich-filler, or "thin disk," part of the overall galactic disk, within the Orion Arm. Of the few Sample-S interlopers born outside the sandwich filling, and now temporarily passing through it on orbits oblique to the thin disk, the best known is α Boo (Arcturus). It is convenient here to use the term "Population P" for the ensemble of non-brown-dwarf, non-white-dwarf stars in the much larger, 3000-ly radius, subdisk-of-the-thin-disk from which our (tiny) Sample S is drawn. This P-region is itself only a (tiny) fraction of the overall galactic thin-disk region of stars, ~50,000 ly in radius. The various pages at atlasoftheuniverse.com are a useful resource for visualization of the Orion Arm, furnishing both a zooming-out to the wider galactic context and a zooming-in to detailed features within the Arm.

Sample S, being formally defined by an apparent-magnitude cutoff as opposed to a distance cutoff, is itself far from statistically representative of Population P. *(a)* In P, the O stars are vanishingly rare. A tabulation by Glenn Ledrew, in *JRASC* **95** (2001), pp. 32ff (bibcode 2001JRASC..95...32L) suggests an O-star frequency within P of just 0.00003%. By contrast, O stars comprise a hefty ~2% of S. A similar overrepresentation occurs for the B, A, F, G, and K stars, with Ledrew's tabulation suggesting that these MK temperature types might have a respective frequency within P of 0.1%, 0.6%, 3.2%, 8.0%, and 12.9%. By contrast, the first three of these five rare types comprise ~30%, ~20%, and ~10%, respectively, of S, and the last ~20% of S. *(b)* In P, something on the order of 76% or 78%—different authorities are perhaps mildly discrepant—must be M stars. (Ledrew's tabulation, in particular, suggests an M-star frequency of 78.2%.) Only a few of these (the Ledrew tabulation suggests 0.04%) have evolved to beyond the main-sequence stage of stable-core hydrogen fusion. By contrast, the M stars comprise just ~5% or ~10% of S. All of them have evolved beyond the main sequence, having started their lives as types hotter than M or K.

The statistically anomalous character of S is further illustrated by the fact that in S, in each of the Big Six MK temperature types hotter than M, the numerical majority comprises the stars that have ended stable-core hydrogen fusion (and so have, as a generally reliable rule—we return below to a necessary caveat regarding reliability—evolved out of MK luminosity class V into one of the brighter MK luminosity classes IV, III, II, or I). In Ledrew's tabulation, the percentages of evolved stars in F, G, and K, as a percentage of the overall respective F, G, and K populations, are just 2.0%, 2.5%, and 3.8%. Consistently with this, the 1991 Gliese-Jahreiss catalogue of the nearest 1000 stars (containing, admittedly, not only the local OBAFGKM VI, V, IV, III, II, and I stars, but also at least many of the local white dwarfs) assigns less than 1% of its population to MK luminosity classes IV, III, II, or I.

Sample S—so rich in varieties of star statistically infrequent within Population P—harbours physical extremes. Although the extremes are for the most part not written into our table, they can be studied easily, from such sources as Prof. James Kaler's stars.astro.illinois.edu/sow/sowlist.html.

At least 58 of our 323-star set each radiate, across the full spectrum from X-ray through UV and optical to IR and radio, at least as much power as is radiated by 10,000 Suns. Possibly the most dramatic is ζ Ori, with a bolometric luminosity of 375,000 Suns—making ζ Ori notable not within S alone, but even within the overall galaxy. Several others are not far behind, among them ζ Pup (360,000 Suns, suggests Kaler, as of July 2008 revising his earlier, circa-1999, suggestion of ~750,000 Suns). We believe that just two stars in Sample S, nearby τ Cet and nearby α Cen B, radiate more feebly than our Sun, each at about half of the Sun's bolometric luminosity.

The principal determinant of stellar luminosity, for any given phase in stellar evolution, is mass, with even small variations in mass translating into large variations in energy output. The exceptional luminosities of ζ Ori and ζ Pup, in particular, are a consequence of their exceptionally high respective masses, 20 $M_\odot$ and 40 $M_\odot$. (Kaler now suggests 40 $M_\odot$ for ζ Pup, while having previously suggested 60 $M_\odot$. He additionally notes from the literature the lower suggested value of 22.5 $M_\odot$.)

Theory does predict, although our small Sample S does not succeed in illustrating, the possibility of masses up to the Eddington stellar-mass limit, somewhere above 100 $M_\odot$, and even of some "super-Eddington" stars. (Eddington's limit is by definition attained when luminosity rises so high as to make the outward radiation push, tending to tear a star apart, exceed the inward gravitational pull.)

Rotation periods in Sample S vary from far in excess of our Sun's to far short of our Sun's (which we may here take as a nominal 27 d; refined treatments of solar rotation provide for rotation-period variations both with solar latitude and with solar depth). Spectroscopy yields for γ Cep a period of 781 d, i.e. of 2.14 y. Kaler suggests that the respective rotation periods of α Hya and ε Crv could be as long as 2.4 y and 3.9 y. Perhaps our slowest rotator, however, is α Ori, now (cf 2009A&A...504..115K) assigned the period of 8.4 y. At the other extreme, Kaler suggests for ζ Aql A, α Aql, and ζ Lep, respectively, 16 h, at most 10 h, and around 6 h.

Radii (as distance from centre to outermost opaque layer, perpendicular to the axis of stellar rotation) are typically greater than the solar radius. Two notable instances of stellar expansion—in other words, of notably tenuous stellar atmosphere—are α Sco (with a radius of 3.4 AU, not far short of the Sun-Jupiter distance) and α Ori (with a radius of 4.1 AU or 4.6 AU from interferometry, or alternatively 3.1 AU or 3.4 AU from luminosity-temperature deductions). Results in these extreme cases depend strongly on the wavelength selected for evaluating opacity. Observations within Population P do indicate, although our sample S does not succeed in illustrating, the possibility of still more-extreme stellar radii, to values approaching ~10 AU. (Among these extreme-radius cases is a vividly red star well known to binocular-equipped observers, though a bit too faint for our table, μ Cep.)

The broad range of temperatures (a topic whose MK conceptual subtleties we examine in subsection 5.1, below) is reflected in the fact that all of the Big Seven temperature-type bins in the traditional MK temperature sequence are well occupied, however statistically skewed (as we have argued above) is the distribution in the MK Big Five luminosity-class bins. At the MK temperature extremes are the hot ζ Pup (O5; 42,000 K) and the cool o (omicron) Cet (M5–M10; a typical temperature for this variable is variously suggested as ~2000 K or ~3000 K).

Interesting spectral anomalies in Sample S include the "Be phenomenon" and "shell spectrum" stars, as discussed at length in our final subsection.

SECTION 3: Initial user guide to the columns in our 315-entry table

In our first column, we use the flags "+nP" (n = 1, 2, ...) for companions of sub-stellar mass, such as have been found outside our Solar System, in an accelerating sequence of discoveries, from the 1990s onward, that has now reached even the tiny Sample S. Such companions are typically planets but could in principle also be brown dwarfs. We do not attempt here to define formally the difference between a planet and a brown-dwarf companion.

In this same column, we apply the WDS naming scheme for multiplicity, both in the case of true binarity and in the case of mere optical doubles (in all but eleven awkward cases, as noted in section 1 above, putting into the first column just the name of the brightest WDS-catalogued component; but we additionally try to supply particulars, at any rate in the online table, for binary and mere-optical companions brighter than mag. 10, in the "Remarks" column, while for the most part regarding binary and mere-optical companions fainter than this limit), using underlining, as discussed in Section 4 below, to flag instances in which a binary system possesses a published orbital solution.

Apparent Visual Magnitude ($\mathbf{m}_V$ = V): Apparent magnitudes, with "v" appended for large-amplitude variables, are from *HIPPARCOS*. In the case of variable, we take as authoritative the ranges (where possible, in V), and also the periods, published in the online AAVSO(VSX) database. Our reasoning here is that AAVSO has critically appraised and filtered data originally presented in more upstream sources, such as the primary (journal-article) literature. Our "V" is the usual "V" of UBV photometry, as introduced by H.L. Johnson and W.W. Morgan in 1953ApJ...117..313J. The (yellow) V filter corresponds roughly to the response of the eye. We retain, without having attempted our own independent error analysis, the assertion of our Handbook predecessor R.F. Garrison (working essentially before *HIPPARCOS*) that the "probable error" of each of our cited V values is at most 0.03 mag. (in other words, that of the actually and potentially available V measurements from the world's duly competent photometry facilities, at least half will lie within 0.03 mag. of our own cited V values). Some small inaccuracies in magnitudes may be present in cases of combined light: readers needing confirmation may check our values against WDS, or where possible against the magnitude-specifying atlas pages of AAVSO. (By the nature of its mission, AAVSO is constrained to supply in its cartography not only details of variables, but also magnitudes of stars that are constant, and which can be used by amateur photometrists as comparison stars and check stars.) We hope to rectify these possible small inaccuracies in the next major revision of this document, in the "8.x.x" series, perhaps around 2022 Dec. 31.

Spectral Classification (**MK Type**): The "MK temperature type" (O, B, A, F, G, K, M) is given first, followed by a finer subtype (0-9) and an "MK luminosity class" (Roman numerals I–V, with "a" or "b" added occasionally to indicate slightly brighter or fainter stars within the class). As we discuss in detail in subsection 5.1 below, O stars are the hottest, M stars coolest; Ia stars are termed the most luminous "supergiants"; III stars are termed "giants"; and V stars are termed "dwarfs." V stars form the largest class in the cosmos, comprising the observational Main Sequence (MS) (as a region in two-dimensional MK-luminosity-class-versus-MK-temperature-type classification space). Other MK symbols include "e" for hydrogen emission; "f" for broad, non-hydrogen emission in hot stars; "m" for strong metallic absorption; "n" or "nn" for unusually broad absorption; "p" for peculiarities; "s" for a mixture of broad and sharp lines; and ":" for a minor uncertainty. (The flags "n" and "nn" are a signature of rotation. It seems that historically "n" and "nn" signified "nebulous", as references to the photographic-plate appearance of a rotationally broadened absorption line.) Where a single star (e.g. α CMa A) is given two types, with the second flagged "m", the first is the type that best characterizes the hydrogen lines, the second the type that best characterizes the metal lines.

MK classifications are in some cases controverted. We have inherited our own types for the most part from the judgements of our predecessor R.F. Garrison, who, as a principal historical authority in MK classification, drew both on what he judged to be the best of the literature and on some of his own unpublished classifications. As of 2021 Jan. 13, we have made a modest beginning at flagging the cases of controverted MK phenomenology (in our online, but not in our printed-edition, "Remarks" column), in two ways: *(a)* Where the literature suggests a real difficulty in MK classification, we draw attention to the difficulty, discussing it in a few words; *(b)* Where we have not found reason in the literature to suspect an MK-classification uncertainty, but nevertheless find our assigned MK type diverging (even in a small way) from the type assigned as of epoch 2021.5 in the official United States Naval Observatory and HM Nautical Almanac Office publication

Astronomical Almanac, Section H (bright stars), we document the divergence, without further discussion.

Parallax (**π**), *Proper Motion* (**μ**), and *Position Angle* (**PA**): Parallaxes, in milliarcseconds (mas), proper-motion vector norms (″/y), and vector position-angles (degrees, from N through E) are derived from the *HIPPARCOS* 2007 data reduction, with a few exceptions. It may be hoped that in future years more precise parallaxes will be forthcoming from the *Gaia* mission, which has now found an engineering solution significantly easing its initial restriction to the fainter stars. (Detector overload had been feared.) Like *HIPPARCOS*, *Gaia* has to cope with the special challenges posed in measuring to high precision (*i*) the parallax of a (orbitally wobbling) star possessing a gravitationally bound, and not necessarily well documented, companion, and (*ii*) the parallax of a star with perturbed photosphere, and consequently with displaced photocentre (as when a tight binary system contains a bright mass-transfer stream).

Absolute Visual Magnitude (**M_V**) and *Distance in Light-Years* (**D**): Absolute magnitudes and distances are determined from parallaxes, except where a colon follows the absolute magnitude; in these cases, both quantities are determined from a calibration of the spectral classification. The absolute magnitude is left uncorrected for interstellar absorption. The appropriate correction is typically ~+0.06 mag. per 100 ly outside the (admittedly very-far-from-spherical) Local Bubble, i.e. beyond ~100 ly. A special difficulty, not fully grasped by us, arises in the case of the controverted ε Aur system distance (for which we now use *Gaia* DR2, additionally supplying references to the recent literature).

We take account of uncertainties in parallaxes by stating the derived distances, in ly, to no more than the appropriate number of significant figures (rounding where necessary). In cases where rounding would itself be misleading, we use a tilde as an indicator of imprecision.

Radial Velocity (**V_{rad}**): Radial velocities are from BSC5. “SB” indicates a spectroscopic binary, an unresolved system whose duplicity is revealed by periodic Doppler oscillations in its spectrum and for which an orbit is possibly known. If the lines of both stars are detectable, “SB2” is used; “+” and “–” indicate, respectively, motion away from and toward the observer. “V” indicates a variable velocity in a star not observable as a spectroscopic binary. (In most “V” cases, the orbit is unknown.)

Remarks: Remarks include data on variability and spectra, particulars of any companions, and (for the most part, only in our online table) prominent bits of observational-astronomy news. We are often a little casual with rounding, stating physical quantities for a given star (as, to take a random example, the angle between the α Cen AB orbital plane and the plane of the sky) to a lower precision than is now available from the primary literature. In a departure from our practice prior to 2017, we now give star names in all and only those cases in which star names are formally promulgated in the International Astronomical Union (IAU) star-naming project, as launched in 2016 at www.iau.org/public/themes/naming_stars. Readers requiring further information on names could start with the individual star descriptions in stars.astro.illinois.edu/sow/sowlist.html. Richard Hinckley Allen’s 1899 book *Star Names: Their Lore and Meaning* has been much cited over the decades. More recent scholarship, with due professional attention to Arabic philology, is, however, presented in Paul Kunitzsch and Tim Smart, *Short Guide to Modern Star Names and their Derivations* (Wiesbaden, 1986), and (by the same pair of authors) *Dictionary of Modern Star Names: A Short Guide to 254 Star Names and their Derivations* (Cambridge, MA, circa 2006). In the **Remarks** column, a **boldface** star name indicates a navigation star.

SECTION 4: Supplementary user guide, concerning our treatment of double-star astrometry

SUBSECTION 4.1: General background remarks on double-star astrometry

4.1.1: Introductory remarks; prevalence of binarity in Sample S: The observer at the eyepiece seeks physical understanding. Here is a speck of starlight, and here at medium power, around 10″ (around a quarter-Jupiter) away is another, with perhaps an intriguing tint difference between pure white and yellowish white, over and above a notable magnitude difference: how far away is each of these two stars, and what stages have they attained in their respective lives, and how do their masses and photosphere temperatures differ?

When the specks are paired, as in this imagined example, a further question arises, however, no less

important than the ones already mentioned. Are these two physically unrelated stars, neighbours on the two-dimensional celestial sphere through coincidence (with one star perhaps twice or five times as far from Earth as the other, but perhaps with the stars even at rather similar distances from Earth)? Or is it, on the contrary, the case that each star experiences the gravitational attraction of the other so strongly as to keep the pair in a mutual orbit, constraining them to move through galactic space as a binary system?
Most stars in our galaxy, and in particular most stars in that portion of our galaxy that is our nearby Population P, not only are red dwarfs in the stable core-hydrogen-fusion phase of their lives but are solitary. (Our own galaxy is a barred spiral. In elliptical galaxies, red dwarfs are still more common, and solitary stars therefore presumably likewise still more common.) However, binarity becomes more prevalent as stellar mass increases. In our Sample S, high-mass stars predominate, and binarity is correspondingly more evident within Sample S than in the overall Population P.

4.1.2: Hierarchically organized systems; contrast with clusters: A situation by no means rare in binarity-rich Sample S is the many-star system hierarchically organized, with binarity at each hierarchy level. One of the good instances is the six-star system whose combined light becomes, in a compact amateur atlas, the starlight point α Gem, and whose most prominent star is Castor (with the ancient name "Castor" since 2016, under IAU rules, designating just a single star, but before that ruling often used loosely, as a name marking merely the overall naked-eye point of light).

We may have, for instance, stars *w* and *x* in a mutual orbit (with *w* describing an ellipse about the *wx* centre of mass C_{wx}, and *x* for its part also describing an ellipse about C_{wx}, but with some remote star *y* experiencing the entire *wx* system as essentially a point mass). In an appropriately chosen rest frame, *y* describes some rather wide, rather slow orbit in tandem with the *wx* pair. In this situation, C_{wx} moves in an ellipse about the *wxy* centre of mass C_{wxy}, and *y* for its part also moves in an ellipse about C_{wxy}. There can now be a still more remote star *z*, experiencing the *wxy* trio as essentially a point mass. In such a case, C_{wxy} will, in an appropriately chosen rest frame, describe an ellipse around C_{wxyz}, and *z* will for its part also describe an ellipse about C_{wxyz}.

Here "appropriateness of choice" consists in taking some inertial frame of mass in which the centre of mass for the given pairing is held at rest. For present purposes, it is enough first to confine our discussion to classical Newtonian mechanics (neglecting the general relativity analysis of gravitation as a geometrical feature of spacetime), and then within this confined framework to further bypass a conceivable deep conceptual problem, saying—simply and loosely—that an inertial frame anchored on some point in space is one in which that point is held at rest (for instance, by being made the origin of Cartesian coordinates for that frame), and in which additionally the frame is "stipulated not to rotate." (A deep question, extraneous to our own limited purposes in this Handbook article, does admittedly arise even in the confined Newtonian-mechanics framework: with respect to what physical standard is it asserted that a given frame "does not rotate"? Without venturing into details here, we remark that in our Handbook view, the deep question—historically prominent in the Leibniz-Clarke correspondence—can be given a conceptually satisfactory answer, by defining "does not rotate" in terms of "absence of centrifugal pseudo-forces violating Newton's force-as-product-of-mass-and-acceleration principle".)

Still further geometries are possible in the case of four stars. In particular, as an alternative to the situation just described, *wxyz* could instead be a double double, with *wx* a tight pair, and *yz* a tight pair, and with these two tight pairs sufficiently distant from each other to make each experience the other as essentially a point mass (with, in this case, C_{wx} and C_{yz} describing, in an appropriate choice of rest frame, elliptical orbits about C_{wxyx}).

Does an intricate system of, say, eight stars, organized into a four-level hierarchy as "{{{{1,2},{3,4}},{5,6}},{7,8}}" or into a six-level hierarchy as "{{{{{{1,2},3},4},{5,6}},7},8}" differ in kind, or only in degree, from a stellar cluster (be it an open cluster, such as the Hyades, or a globular, such as M13)? We argue that the difference is one of kind.

In developing our argument, we first run through some formal preliminaries.

Consider any pair of point masses *P*, *Q* in three-dimensional space, subject to any arbitrary assemblage of forces, and never occupying the same location in space at the same time, and with their respective masses constant through time. For each member of the pair, in any frame of reference, its velocity at any instant (no matter how simple or intricate may be the assemblage of acting forces) is perpendicular to its angular momentum at that instant. We call this the "Basic Momentum-Geometry Fact."

For any binary point-mass system *P*, *Q* (perhaps experiencing some very intricate assemblage of forces, some of them exerted somehow by bodies other than *P* and *Q*), call a force experienced by *P* or by *Q* "central" if and only if it is either parallel or antiparallel to the vector from *P* to *Q*. Newtonian gravitation in an isolated two-body system is a central force, as also are electrostatic forces of attraction and repulsion for isolated two-body systems in particle physics. (More concretely, *P* experiences in the astronomical or particle-physics isolated-binary-system case a force exerted by *Q*, acting along the *P*–*Q* line, whereas *Q* experiences an equal and opposite force exerted by *P*.) If, as is the case for an isolated two-point-mass gravitational or electrostatic system, *P* experiences at every moment nothing but the one central force exerted on it by *Q*, and *Q* experiences nothing but the equal-and-opposite central force exerted on it by *P*, then (via a one- or two-line proof in differential vector calculus) the angular momentum vector of *P* is constant and the angular momentum vector of *Q* is constant. But then, by the "Basic Momentum-Geometry Fact," each of *P* and *Q* describes, in any inertial frame in which C_{PQ} is at rest, a curve confined to a plane.

In astronomy, gravitating bodies are often close to possessing a spherically symmetric mass distribution. It is provable (with a many-line argument; Newton held up publication of *Principia* by a couple of decades until he at last had the proof) that such bodies behave as point masses.

With these formal preliminaries complete, we now give our argument.

In a hierarchical system, each component (whether a single star, or a pair of stars, or a double double, or a lone outlier in tandem with a tight double, …) experiences as non-negligible only a central force exerted by its nearest companion. It therefore describes, in any inertial centre-of-mass frame for that binary within the hierarchy of nested binaries, a curve that is confined to a plane. In an arbitrary stellar cluster, on the other hand, the described curves are not in general plane curves, but can be so-to-speak warped ("twisted," "skewed").

We add here that warped-curves cases might arise even in systems simpler than open clusters and globular clusters, including groupings that are not stable: one example of such a small, unstable, grouping is the handful of stars at the heart of the Orion Nebula, whose four brightest members constitute the Trapezium (and whose very brightest member shines at mag. 5.1, less than two magnitude steps below some of our own "Brightest Stars").

4.1.3: Two-body systems and the conic sections, and orbiting as a kind of oblique falling: None of this argument makes any assumptions about the mathematical form of the central force. It is, however, provable (again as a non-trivial theorem, requiring a multi-line proof) that in the special situation in which the central force obeys an inverse-square law (as is the case for gravitation, and indeed also for the electrostatics dominant in particle physics), the curve is not only confined to a plane, but assumes the specially simple form of a conic section. The section is a hyperbola if the central force happens to be repulsive and is a hyperbola or parabola or ellipse if the central force happens to be attractive. The hyperbola case is illustrated in the mechanics of our solar system by those comets that are moving too quickly to be captured into closed-curve orbits around the Sun. Ellipses are of course common in celestial mechanics: for an isolated binary in an inertial centre-of-mass frame, there is a plane P such that each of the two components describes an ellipse in P with the centre of mass at one focus, with the two components always on opposite sides of the centre of mass, and with the more massive component moving in the smaller of the tandem ellipses, and the tandem ellipses being of the same shape even if quite different in size (being, that is, figures in P that are similar-even-if-not-congruent). That conic section that is the parabola is, on the other hand, from the point of view of real-life celestial mechanics, a mere mathematical idealization, constituting the so-to-speak infinitely thin boundary between the cases of the hyperbola and the cases of the ellipse. More formally, a parabolic trajectory is realized in an inertial centre-of-mass rest frame of an isolated two-point-mass system when and only when the relative speed of the two masses exactly equals the least upper bound of those various relative speeds, relative to the centre of mass, which are low enough to yield an ellipse.

The soccer ball, as a projectile on the sports ground, constitutes a two-body system with Earth. It is often stated, as an approximation, that the impelled ball describes a parabola. Here, however, the truth is that the trajectory is a segment of an ellipse, minutely divergent from a parabola, and that the trajectory would become a perfect parabola if the gravitational field on the sports ground were, contrary to fact, to be of constant direction. (Take the Earth to be a sphere of perfectly spherical mass distribution: then the gravitational field across the soccer ground, exerted by Earth on the ball, changes everywhere in direction, pointing everywhere in the soccer ground to that single point that is the Earth's centre.)

In the limiting case in which a soccer ball is impelled almost directly upward, and so falls almost directly downward, the trajectory becomes a segment of some almost-degenerate ellipse, with its minor axis of almost negligible length. Since Earth is so overwhelmingly more massive than the soccer ball, the common centre of mass of the Earth-ball system nearly coincides with the centre of the Earth, and the Earth's ellipse around the centre of mass becomes correspondingly of sub-sub-atomic dimensions. (Admittedly, this is the situation in a Newtonian setting. It will be interesting to see what becomes of such entailed sub-subatomic ellipses if, at some future era, general relativity becomes successfully unified with quantum mechanics.) Upon reflecting on this sportsground example as an extreme case, it can rather soon be seen (we omit details) that a binary system in stellar astronomy is a case of falling—in which, however, the two bodies fall toward each other in such a way as to stray a little off the line at any instant connecting them, and so are destined never to meet up. If the falling is at all times only a little off the instantaneous connecting line, the ellipse is severely elongated (has an "eccentricity" just slightly less than 1; in our brightest-stars table "Remarks," we write "e" in our occasional reports of known orbit geometries); if, on the other hand, the falling is at each instant as far off the instantaneous connecting line as geometrically possible, so that at each instant each body is moving perpendicularly to what is at that instant the inter-body connecting line, then the ellipse is a circle (with e=0).

4.1.4: Binary systems and the determination of individual masses: For foundational astrophysical reasons, much effort has historically been, and is still now being, expended on the documenting of binary stars. This work was pioneered with the filar micrometry of William Herschel (1738–1822) and (more systematically) F.G.W. von Struve (1793–1864). The work took on fresh vigour with the late-Victorian advent of radial-velocity spectroscopy, as spectrograms began to be measured under the microscope for Doppler shifts. Since 1980 or so, it has taken on still greater vigour with the rise of interferometry–speckle interferometry and aperture-masking interferometry in the case of observatories possessing lone telescopes of large aperture, but more notably with optical tables and beam combiners, at those few facilities in the astronomical community possessing multi-telescope interferometric arrays (notably, at CHARA, NPOI, and VLTI).

Always, from the pioneering filar micrometry onward, the astrophysical motivation has been the same. Once a full orbital solution for a binary system of known distance is determined, the individual masses of the two components are known—both *(a)* as multiples of Solar System quantities and *(b)* in the absolute Système international d'unités (SI) laboratory unit, which is the kilogram. Here a "full orbital solution" is a set of half a dozen geometrical parameters, or "orbital elements," in essence angles, describing the ellipticity of the orbit and its orientation in three-dimensional space (including its angle of inclination with respect to the plane of the sky). With these, plus a determination of the distance to the binary system, the orbital trajectory of the binary (in any inertial centre-of-mass rest frame) is fully described, in particular with the length of its semimajor axis determined in the laboratory SI unit of metres.

It is, admittedly, an intricate task to proceed to the orbital elements from the little that is available at the telescope. In the case of traditional filar microscopy, orbital elements are in principle obtainable from some years or decades of raw astrometry, with each night supplying just the angular separation of the components and their position angle (as an angle in the half-open interval [0°, 360°[, taken from sky North through East, South, and finally West; in our table "Remarks," we write "PA"). Also in principle obtainable are orbital solutions for binaries of known distance in which there is no filar-micrometer astrometry, and also no other (for instance, interferometer-procured) astrometry, but in which the plane of the orbit happens to be exactly perpendicular to the plane of the sky, and in which spectroscopic Doppler-shift measurements have over many successive observing sessions supplied the changing radial velocities of each of the two components. Although the situation in which the plane of a binary orbit is exactly perpendicular to the plane of the sky is seldom, if ever, realized in observational work, the situation is approximated to usable precision by cases in which the two stars are found to eclipse each other (as with, e.g. β Per Aa1 (Algol) and β Per Aa2). Finally, an orbital solution may be obtained from a combination of (perhaps imperfect) astrometry and (perhaps imperfect) spectroscopy.

Let it, in any case, now be taken that a binary system of known distance has had its orbital elements determined, by some means or other.

The solution to problem *(a)* (individual stellar masses to be determined as multiples of Solar System quantities) rests on the fact that there exists a "universal gravitational constant" G, such that for any two-body system of constant point masses m_1 and m_2, in the gravitationally bound, elliptical-orbit case the sum m_1+m_2 obeys, under Newton's generalization of Kepler's Third Law, the equality "$m_1+m_2 = (4\pi^2\, a_{1,2}{}^3)/(G\, P_{1,2}{}^2)$" (with $a_{1,2}$ the length of the semimajor axis of the orbit (in any convenient inertial rest frame of the m_1,m_2 centre of mass), and $P_{1,2}$ the m_1,m_2 orbital period). The law as stated here is independent of units: masses could be measured in kilograms, or in any other convenient units; the distance, which is $a_{1,2}$, could be measured in metres, in light-seconds, or in any other convenient units; and time could be measured in any convenient units. Let, now, M_{EM} and $M_\odot$ be the respective masses of the Earth-Moon binary and the Sun. It then follows as a special case, and under the (in practice sufficiently good) idealization of the Earth-Moon binary and the Sun as an isolated system of two point masses that $M_{EM}+M_\odot = (4\pi^2\, a_{\mathrm{EM},\odot}{}^3)/(G\, P_{\mathrm{EM},\odot}{}^2)$ (where $P_{\mathrm{EM},\odot}$ is the orbital period, in any convenient inertial rest frame of the centre of mass of those two entities, which are the tight Earth-Moon binary and the Sun, of that wide binary, which is the Earth-Moon centre-of-mass and the Sun). Equating the ratio of the left-hand sides of this pair of equations with the ratio of the right-hand sides of this pair of equations, and additionally equating $a_{\mathrm{EM},\odot}$ to the physical quantity 1 AU as defined since 2012 in the SI unit of metres at IAU (this equating, while not exact, is an excellent approximation), we have $(m_1+m_2)/(M_{EM}+M_\odot) = (a_{1,2}/1\ \mathrm{AU})^3/(P_{1,2}\ /\ 1\ \mathrm{y})^2$. Conveniently, however, M_{EM} is to one significant figure a mere 3-millionths the mass of the Sun. This justifies replacing, for most ordinary astrophysical purposes, $M_{EM}+M_\odot$ with $M_\odot$, yielding as a final, good approximation, the following solution to problem *(a)*: $m_1+m_2 = M_\odot\, (a_{1,2}/1\ \mathrm{AU})^3/(P_{1,2}\ /\ 1\ \mathrm{y})^2$.

It remains to determine not just m_1+m_2 in terms of the quantities 1 AU, 1 y, but the individual stellar masses m_1, m_2 in terms of this pair of quantities. This, however, is a comparatively modest further step. Once the orbit of the binary system, in some convenient centre-of-mass inertial rest frame, is given, the mass ratio m_1/m_2 can be found by comparing the respective dimensions of the two similar-though-in-general-not-congruent ellipses (the smaller ellipse for the larger of the two masses) traced around the common centre of mass in any convenient inertial centre-of-mass rest frame. With m_1+m_2 known and m_1/m_2 known, the individual values of m_1 and m_2 follow.

Problem "*(a)*" has thus been solved without recourse to a laboratory determination of the troublesome constant G. It is for problem "*(b)*" (determination of m_1, m_2 individual values in kilograms) that G itself is needed. Work on the laboratory problem has been proceeding for a little over two centuries. Google or YouTube searches under such terms as "Cavendish experiment apparatus" reveal the possibilities for repeating, under a constrained high-school budget or a frugal lone hobbyist's budget, the result published by Henry Cavendish in 1798, and falling with around 1% of the now-accepted value. As the current state of the art, where expense is surely not spared, www.pnas.org/content/113/36/9949 cites a *Phys. Rev. Lett.* year-2000 determination of G to an uncertainty of 0.0014%. Even this modern level of precision compares unfavourably with the precision attainable for, e.g. the electron charge-to-mass ratio, the speed of light, the electrical permittivity of free space, and the magnetic permeability of free space. Nevertheless, G, while continuing to be something of a laboratory embarrassment, is sufficiently well known to facilitate work in those areas of astronomy (notably in planetary science) where actual kilogram masses are useful. Already in Cavendish's day, for instance, it was determined (we here rephrase Cavendish's result in modern terminology, while conserving its substance) that the mass-per-unit-volume of Earth is higher than the mass-per-unit-volume of ordinary rock (planet Earth 5515 kg/m^3, but basalt and granite merely ~3000 kg/m^3), and that therefore the rocks familiar to geology are not representative of Earth's deeper interior. SI-unit density determinations, resting on the determination of masses in kilograms, are now needed not in geophysics alone but in exoplanet work, for instance for supporting hypotheses regarding a given exoplanet's composition (gas, in the manner of Jupiter? or something more dense, in the manner of Earth?).

4.1.5: Some further reading: Tutorial resources on the Web include a conspicuously thorough source of pages from an author of the *Cambridge Double Star Atlas* (2009, second edition 2015), Bruce MacEvoy (the colleague author for this book is the celestial cartography authority, Wil Tirion), at www.handprint.com/ASTRO/index.html.

SUBSECTION 4.2: Our notational conventions in table "Star Name" column for double-star astrometry

Our treatment of double stars follows the WDS naming rules, but with additionally our own (purely Handbook-local) underline-flagging convention.

Suppose, as a hypothetical case, that a certain bright naked-eye point source has been familiar from Johann Bayer's 1603 atlas onward as "omega FooBaris," or ω FBr. Suppose ω FBr to have been discovered by some 1830s filar micrometrist to be a tight double, with components separated on the celestial sphere by an angular distance of 0.7″. It does not matter whether the pair is a binary or a mere line-of-sight coincidence: in either case, at the 1960's launch of WDS, the pairing is catalogued as ω FBr A and ω FBr B.

Now suppose, as a refinement of this basic scenario, that around 1910, ω FBr A was found by some specroscopist to be a spectroscopic binary (in our penultimate-column notation, to be an "SB"), and that nothing further was known about ω FBr A until 1974. What are the 1973 WDS implications of the 1910 discovery? Under WDS rules, ω FBr A had at that early stage in the development of WDS to be ω FBr A (not ω FBr Aa, ω FBr Ab), since as of 1973 its components had not been measured in the celestial-sphere terms of PA and angular separation.

Stellar interferometry was launched in a modest way in the 1920s. It is perhaps reasonable to say that a "Second Generation" of optical interferometers was ushered in by the team of Robert Hanbury Brown, operating the Narrabri Stellar Intensity Interferometer from 1963 to 1974. Suppose, then, that in 1974 some interferometer, such as Narrabri, succeeded in resolving ω FBr A into two components, say at a measured separation of 0.1″. At this stage, the WDS multiplicity catalogue was at last able (and under its self-imposed rules was required) to refer not to "ω FBr A" but to ω FBr Aa and ω FBr Ab.

Finally, suppose that in the current, arguably "Third," generation of optical interferometry, some such instrument as CHARA or NPOI or VLTI, perhaps working in the year 2020 or 2030, measures ω FBr Ab itself as a (very tight, very rapid) binary, with the separation even at apastron found to be just a few tens of milliarcseconds. At this stage, WDS is able (and under its self-imposed rules is required) to refer not to ω FBr Aa and ω FBr Ab but, rather, to ω FBr Aa, ω FBr Ab1, and ω FBr Ab2.

In the leftmost column of our table, we indicate with underlining that a published orbital solution is asserted to exist in WDS. In our notation, "ω FBr Aa" signifies the existence of a published orbital solution for ω FBr Aa and {ω FBr Ab1, ω FBr Ab2} (where the star ω FBr Aa experiences the outlying pair of stars ω FBr Ab1, ω FBr Ab2 as essentially a point mass), whereas "ω FBr Aa" signifies in our notation the existence of a published orbital solution for the entire {ω FBr Aa, ω FBr Ab1, ω FBr Ab2} three-star system, considered as a point mass, in its wide and slow orbit with the remote ω FBr B. In various cases in which this notation is, whether definitely or at least arguably, unclear in its intent, we explain in the "Remarks" what is and is not available in the published ω FBr orbital-solutions literature.

Although the presence of underlining in our leftmost column is a safe indication that a given double is a binary, the absence of underlining is not a safe indication that a given double is a mere line-of-sight coincidence. In some cases lacking underlining, it is a known fact that the given double is a binary (typically with some very wide, slow, orbit, that will defy mathematical modelling until some further centuries or millennia of astrometry become available); in other such cases, it is a known fact that the given double is a mere line-of-sight coincidence (for instance, because either the parallaxes or the proper motions of the two stars are severely discrepant); and in very many other such cases, the answer to the question "Binary, or not?" is currently unknown. Although we do not here try to flag the first and the second of these three possibilities in our leftmost column, WDS does try to track the current state of knowledge with its own (duly elaborate) flagging system.

SECTION 5: Supplementary user guide, concerning the more detailed interpretation of our MK-classification column

SUBSECTION 5.1: Conceptual underpinnings of the MK classification system

In strict conceptual accuracy, the MK temperature types are a purely phenomenological record of which elements are present *(a)* in which stages of ionization, and *(b)* at what densities (in other words, under what local strength of the local downward-directed gravitational field) in the photosphere of the given star.

Decades before the 1943 Morgan-Keenan-Kellman publication of the full two-dimensional MK scheme, it had already been found possible to set up the phenomenological spectral types under our heading "*(a)*" in a single orderly OBAFGKM sequence, in which individual types gave way smoothly to their neighbouring types. (This process was itself not quite straightforward. First came a simple Harvard "A, B, C, D, ..." scheme. This was followed by the realization that "A," for example, linked smoothly in its phenomenology with "B" and "F," with some of the old alphabet having to be altogether dropped or repurposed. In working out this ordering, it was found necessary by the Harvard pioneers to subdivide the OBAFGKM categories, for instance in the sense of "G rather similar to F" and "G rather similar to K" and "G about equidistant between F and K." Hearnshaw's *Analysis of Starlight*, now in its second edition as 2014anst.book.....H, is the definitive history both of the MK scheme and of its predecessors.)

It was then not a matter of definition, but of astrophysical discovery (cf, e.g. the already-cited 2014anst.book.....H, or again 1994AJ....107..742G, or again the detailed MK reference-work exposition 2009ssc..book.....G), that the OBAFGKM sequence corresponded to a temperature-ordered sequence of stellar groupings, running from the hottest photospheres to the coolest, with each of the various subdivisions within each of the O, B, A, F, G, K, and M types corresponding to a particular temperature range.

With the 1943 introduction of the two-dimensional MK scheme, the luminosity classes I, II, III, IV, V likewise had strictly a phenomenological, not an astrophysical, definition (proceeding now from our heading "*(b)*", as opposed to the "*(a)*" that yielded O, B, A, F, G, K, and M). It was then once again conceptually speaking not a matter of definition, but of astrophysical discovery, that the I-through-V sequence corresponded to a luminosity-ordered sequence of stellar groupings, running from the most luminous to the least luminous.

Admittedly, this conceptual picture, for the history of work under our heading "*(b)*", is idealized. It was evident on the theoretical front even some decades before 1943 that the "*(b)*"-heading phenomenological features highlighted in 1943 by the developers of the MK taxonomic system, and signalling differences in photospheric gas densities (i.e. to differences in the strength of the local downward-pointing gravitational field) in fact correspond to differences in stellar luminosities. The developers of the MK taxonomy thus had a theoretical motivation for their definitions of classes I, II, III, IV, and V, resolutely phenomenological though their definitions were required to be, under observational-astrophysics methodology. — The MK system now serves as a paradigm of successful taxonomy, even for fields outside astronomy. A classification system is defined in terms of mere phenomenological fieldwork, and yet in the expectation (successfully realized in the case of MK) that the phenomenological classification bins will in due course be discovered by the theoreticians to correspond to relevant, important, physical differences in the materials observed. (Parallels might be suggested with, e.g. 18th- or 19th-century medicine: whereas *(i)* the old clinical-phenomenology definition of "tertian fever" and "quartan fever," in terms of the observed duration of body-temperature anomalies, have been found in physiology theory not to correspond to useful fundamental realities at the level of microbiology, *(ii)* the gross empirical observation, as with the pre-Victorian stethoscope, of heartbeat anomalies has been found to correspond to useful fundamental realities at the level of cardiac neuroanatomy.)

When the MK system was introduced, it was already evident that if the classes I through V signalled a progressive decrease in stellar luminosities, then they had to signal a corresponding progressive decrease in stellar radii. The temperature of a given photosphere determines the amount of energy that photosphere radiates per unit time per unit of photosphere area. Consequently, if two stars in the same temperature type are found to differ in luminosity class, the one in the brighter luminosity class must have a larger total photosphere area, and so must be of greater radius.

It was therefore natural to adopt theory-informed, but nevertheless in official terms purely mnemonic, labels for the phenomenologically conceived luminosity classes, with I called for convenience the "supergiants," II the "bright giants," III the "giants," and IV the "subgiants." V had to be given some mnemonic label opposed to "giant," with "dwarf" consequently pressed into service, and "subdwarf" used for the underluminous class VI (important in studies of congenital metallicity, but irrelevant to our own Sample S). (It is admittedly troublesome that the terms "white dwarf"—and nowadays also "brown dwarf"—prove necessary in other contexts, with the "white dwarfs" and the now-celebrated "brown dwarfs" radiating at luminosities far below even classes V and VI.)

SUBSECTION 5.2: MK classification and stellar evolution: preliminary remarks

In 1943, when the MK system was introduced, stellar-evolution theory was not yet on a sound footing. Only the broad outline, that a star may be expected to increase in photospheric radius after completing the fusion of hydrogen in its innermost portion, was at that point known. With the theoretical nuclear-physics advances of the 1950s and 1960s, and with the advent of increasingly detailed computer modelling from the 1960s onward, it became possible to map the elaborate excursions (we outline these in subsections 5.7 and 5.8 below) that evolving stars perform in the two-dimensional luminosity-class-versus-temperature-type phenomenologically defined MK plane (the "observational HR diagram"). In particular, it is now known that every star in the phenomenological class V in our 323-star set from our 315-entry table is still performing stable fusion of hydrogen in its innermost portion. (We repeat that this class V is best termed, with correct deference to the MK classification conceptual underpinning, not simply the "Main Sequence" (MS), but the "observational MS"—as at p. 342 of the authoritative 2006ima..book.....C.) Further, membership in the phenomenological class IV is a good (though even in our small 323-star set not an infallible) indicator that stable hydrogen fusion in the innermost portion is over, with the subject star now having performed at least some part of its (in general, elaborate) later-life excursions over the MK phenomenological plane.

The distribution of the set of 323 stars across MK luminosity classes I through V accordingly turns out to be a reasonable indication of the evolutionary spread of the set.

It follows that the naked-eye bright-star night sky is a different place from the daytime sky, with its lone proximate class-V star. Something on the order of a mere fifth of our 323 MK-classified bright nighttime stars (for the most part stars in luminosity class V) resemble the Sun (the sole daytime object in our set of 323 MK-classified bright stars) in stably burning hydrogen at their centre. Even most of these are far hotter than the Sun and are consequently destined to spend less time than the Sun in in this process of stable burning. All the rest have in one way or another moved beyond that stage, as shown by their luminosity classes—with the nocturnal 322 falling overwhelmingly into classes III and IV, but with classes I and II also rather well populated.

SUBSECTION 5.3: MK classification and stellar evolution: starbirth and MS

A star has at birth (i.e. has upon condensing sufficiently from its local ISM cloud to begin hydrogen fusion) four key characteristics. If the star happens not to be in the disturbing environment of some proximate star (most notably, in the disturbing environment of a binary companion so close as to transfer matter) then these four characteristics jointly entail its various other characteristics, for each point in its entire subsequent career. Prominent among those other characteristics are the duration of overall life, and at each point in the overall life additionally those time-varying key characteristics, which are radius, luminosity, and its photosphere effective temperature. Here, then, are the "Governing Four": *(a)* birth-epoch mass (the more massive stars are also the hotter, the more luminous, and the shorter-lived); *(b)* birth-epoch elemental composition (the most important aspect of composition is simply the birth-epoch "metallicity"—i.e. the extent to which, thanks to the specific properties of the local gestating ISM cloud, the subject star contains at the time of birth any elements, in whatever detailed proportions, heavier than hydrogen and helium); *(c)* absence or (possible) presence of inherited fossil magnetism, from (possible) magnetism in the gestating ISM cloud; and *(d)* birth-epoch speed of rotation.

Of the four listed properties, the first is the most important, accounting, along with the accidental circumstances of distance-from-Earth and time-elapsed-since-gestation, for essentially all the stellar variety that the naked eye can discern.

Regarding the accidental circumstance of time-elapsed-since-gestation, a parenthetical caveat, relevant even to interpreting the casual naked-eye experience, is needed: stars condensed from the same ISM cloud are of the same age. This is the case not only with binaries but also, more dramatically, with associations (such as the dramatic naked-eye association in the northern sky whose most familiar members comprise β UMa (Merak), γ UMa A (Phecda), δ UMa A (Megrez), ε UMa A (Alioth), and ζ UMa Aa (Mizar), in other words comprise all but the first and last of the seven Big Dipper stars).

In contrast with mass and present age, congenital elemental composition does not vary greatly across our set of 323 MK-classified bright stars. The pronounced chemical differences across the set of 323 (evident from the notations for chemical peculiarities in many of the 323 bright-star MK types in our 315-entry table) are due, rather, to processes of stellar aging, notably *(i)* gravitational settling and radiational lofting of selected elemental species in cases in which the outer layers are quiet (in particular, not rotationally disturbed), and *(ii)* processes known as "Dredge-Up" (discussed again in subsection 5.8, below), when convection in an evolving star raises such elements as carbon or nitrogen into the photosphere from the buried thermonuclear furnaces.

We will not attempt to discuss congenital magnetism. But we do remark that, like chemical peculiarities, magnetism can develop and change as a star ages (with, for instance, convection in outer layers, under rotation, producing a dynamo, and with the dynamo in turn generating the kind of looping-field locally dipolar magnetic structures present in the Sun, and hinted at in the small telescope by the Sun's appearance through a hydrogen Balmer-α filter).

The fourth property in our list, congenital rotation, is a consequence of the vagaries of possible motions in the gestating ISM. The local part of the condensing gas was likely to have some kind of coordinated spin, and this spin tended to increase, under conservation of angular momentum, as the gas became more and more condensed—even though some angular momentum also was possibly shed via gas outflows, as the condensation proceeded toward starbirth.

We will not discuss congenital rotation further. We do, however, remark that the rotation speed of a solitary, undisturbed star is once again a property that can evolve as the given star ages, under the combined influence of its evolving mass distribution (although the mass of all but the hottest stars remains rather constant until late in life, after cessation of core hydrogen fusion the mass gets distributed over larger radii, forcing (under conservation of angular momentum) an increase in rotation period and its (possibly, as already noted, evolving) magnetism.

The process of change has two aspects. On the one hand, as an aging star evolves out of luminosity class V into IV, III, and in the case of congenitally massive stars even into II or I, increases in its radius cause (because angular momentum is conserved) a slowing of rotation.

On the other hand, and quite apart from this general slowing-through-bloating, a spin-braking mechanism exists within class V for those stars that succeed in generating the right kind of local, looping, dipole magnetic-field structures. The mass shed by such a star in winds, although modest, is nevertheless constrained by magnetic fields not to orbit the star freely, but to rotate at the about the same angular velocity as the star itself. Under conservation of angular momentum, this so-called "magnetic braking" then slows the rotation. In the overall galactic population of V stars, those cooler than MK temperature type F5 are capable of achieving magnetic braking, and those hotter than F5 are not. The F5 type thus constitutes a so-called "rotation break" within class V.

In our set of 323 MK-classified bright stars, all but six of the class V stars lie on the hot side of the break. The brightest V-class stars in Earth's sky have to be either the most luminous, and therefore the hottest, or those nearest to Earth. The scarcity of V-class bright stars on the slow side of the rotation break therefore indicates that it is the first of these two brightness-promoting characteristics that predominates, in our overall set of 323.

Although we here largely neglect stars in the disturbing environment of other proximate stars, we do have to remark parenthetically that in the case of a close binary, rotation (like also chemical composition) can be affected by processes of mass transfer. This is very notably the case with one of the more heavily studied stars in the 323-member set, α Leo A (Regulus). Here the rapid rotation is the result of a now-completed spinning-up process, involving a copious mass transfer, from the now diminutive, and therefore now observationally elusive, pre-white dwarf. In the "Remarks" for α Leo A in the table, we point out that this elusive companion, having for decades escaped observation, is at last reported in 2020ApJ...902...25G as detected spectroscopically.

The F5 "rotation break" within MK luminosity class V is ultimately due to, and is nearly coincident with, a transition (as one proceeds along the observational MS from the hottest stars to the coolest, i.e. as one advances in the sense OBAFGKM) from stars in which the hydrogen fusion is predominantly the work of the carbon-nitrogen-oxygen (CNO) cycle to stars in which the hydrogen fusion is predominantly the work of the proton-proton chain. The point at which the two processes deliver, per unit of fusion-depth mass, roughly equal energy-per-unit time is at or near a total stellar mass of 1.2 $M_{\odot}$.

To what extent are the four key properties reflected in the MK type of a young star (in observational terms, a star found to lie in MK luminosity class V)?

(a) Mass is well correlated with MK temperature type, in the sense that the OBAFGKM progression within class V proves to be a progression from the most massive stars to the least massive. This fact is itself far from obvious. It was, however, established in the early decades of the 20th century by spectrally classifying the elements of binary systems, of known distance, in which the orbit is not so tight as to allow the disturbing feature of mass transfer, and yet in which the orbit is tight enough, and consequently fast enough, to permit determination of orbital geometry and orbital period. For such binaries, individual masses can be determined from Newtonian mechanics.

(b) Birth-epoch elemental composition is not really reflected in the observationally assigned MK class. We have already remarked that the elemental-composition flags present in many of the 323 bright MK types are due, if not to "Dredge-Up" in the case of an aging star, then to segregation of elements through gravitational settling and radiative lofting (processes that can occur even for a young star, provided its atmosphere is quiet, as in cases where rapid rotation is absent).

(c) The MK scheme does not attempt to flag magnetism, even though magnetism is observed spectroscopically, through the Zeeman splitting of emission and absorption lines when a magnetic field is strong.

(d) Rotation can be inferred in favourable cases, but not in all cases, from the presence of the MK-type flags "n" and "nn." In a favourable case, a rapidly rotating star is seen more or less equator-on, causing its emission and absorption lines to be Doppler-broadened (since half of the photosphere is rapidly receding from the spectrograph, and the other half rapidly approaching it). In, however, the unfavourable case in which the star is seen more or less pole-on, there is no rotational broadening. A particularly well-known example of a rapid pole-on rotator (with "n" and "nn" therefore absent from the observed MK type) is α Lyr A (Vega).

We might add by way of background that it is only in recent decades that the detection of pole-on rotators has become feasible at all. If the star is close and bright enough, interferometry, while powerless to detect the shape deformation of a pole-on rapid rotator, may nevertheless succeed in picking up the equatorial darkening that accompanies rotational flattening (in the pole-on case, as an anomalous darkening, over and above the normal "limb darkening," toward the edges of the interferometrically discerned stellar disk, at whose centre is the Earth-facing stellar pole).

SUBSECTION 5.4: MK classification and stellar evolution: rotation largely neglected here

It is now helpful to outline the various possibilities for stellar evolution, as experienced by that majority of stars in the 323 MK-classified set that are already in MK luminosity classes IV, III, II, or I, as opposed to the "observational MS" which is class V. But an initial caveat is needed: we here largely neglect the disturbing influence of stellar rotation, important though that influence is.

Regarding rotation, we do remark at this point that rotation can produce flows of matter along lines of stellar longitude ("meridional flows"), and that where such flows extend some significant distance into the stellar interior, they help replenish the supply of hydrogen, as a thermonuclear fuel, in the stellar depths. The effect of rotation is in general to somewhat shift the evolutionary track of a star on the phenomenological MK plane (by promoting mixing of stellar layers that would otherwise be more sharply separated) without radically changing the shape of the track.

Difficulties in constructing an evolutionary model for the interior of a rapid rotator are among the themes of Section 1 in 2011ApJ...732...68C. This same paper discusses difficulties involved in deducing the mass and age of a rapid rotator, and the problem of deviations from the von Zeipel 1925 gravity-darkening law for oblate-spheroid stars. The law would give the correct result for gravity darkening if the flattened star had a purely radiative envelope. With rotation, however, gravity darkening can lower the photosphere effective temperature at the equator, causing convection to set in there even when the envelope is radiative at the poles. In our 323-star set, this pathology is present in at least α Aql A (Altair) and α Cep A (Alderamin).

Even where the convective regime is uniform, it is not possible to assign a single photospheric effective temperature to a rapid rotator: its observed MK temperature type is now a mongrel, the result of light entering the spectrograph from the differing temperature regimes of (hot) poles and (cool) equator.

SUBSECTION 5.5: MK classification and stellar evolution: structure, energy flows

As a further preface to details of evolution, it is now necessary to introduce discussion-guiding concepts of stellar structure and stellar energy flows.

A star still stably fusing hydrogen in its innermost portion (whether predominantly via the CNO cycle or predominantly via the proton-proton chain) is said to have a hydrogen-fusing "core." The layers outside the energy-producing "core" of such a star are said to comprise its "envelope." Under this definition of "envelope," the envelope is not a place of energy generation, but merely a place of energy transport. This transport involves a cascade, in which a single core-produced photon is absorbed by some envelope atom, causing the envelope atom to re-radiate multiple photons, each individually less energetic, and with the same aggregate energy as the now-vanished input photon. Each of these less energetic photons is in turn absorbed by some envelope atom in a still higher layer, which in its turn re-radiates a plurality of correspondingly less energetic photons. Eventually, as that outer-skin part of the envelope that is the photosphere is reached, photons begin travelling freely, without processes of absorption and re-radiation.

Those young stars with cores hot enough to have the CNO cycle as their principal mode of hydrogen fusion have convective cores. In the case of the very hottest O stars (perhaps hotter than any of the 35 or 40 or so O stars in our set of 323 MK-classified bright stars), not only the core but even the envelope is convective. The more usual case, however, for a CNO-dominated star, and perhaps the only case appearing for the CNO-dominated subset of our 323-star set, involves a convective core overlain by a radiative envelope.

Where the temperatures at the core are low enough for the proton-proton chain to predominate, the core of a young star is radiative. High envelope opacities in this low-temperature case make radiation an inefficient mode of energy transport, causing envelopes to be convective. As one advances along the temperature sequence in the sense OBAFGKM, stars at first present just a thin convective layer (settling in at a photosphere effective temperature of ~8300 K), with convection then running deeper and deeper (and in particular, in the case of our own Sun, as a G2V star, pervading the entire envelope).

Here (once again) a caveat is necessary regarding rotation. A rapid rotator can straddle the ~8300 K boundary, with convection absent at its (hot) poles, and at least a thin convective layer present at its (cooler) equator.

As an irony of nature, an extreme case exists at the cool end of the OBAFGKM progression, just as for its already-discussed hot end. In the coolest young M stars, convection extends all the way down to the core. As for the extreme O stars, so also, however, the extreme-M case is irrelevant for us: our set of 323 MK-classified bright stars contains no young M stars at all.

SUBSECTION 5.6: MK phenomenology of early evolution within the theoretically defined MS

Having so far mentioned just the "observational MS," we may now proceed to the theoretical definition of the MS, or more strictly of departure-from-MS (and soon we shall also be relating this bit of theory to the already-presented observational MS concept). The theoretical MS will turn out (subsection 5.8, below) to be defined in such a way that departure perhaps can occur already within class V, but can also be delayed until an aging star has brightened enough to take it into class IV.

It is a sufficient, although not a necessary, condition for a star lying within the theoretical MS that it be still fusing hydrogen within its core.

Even within this early, seemingly placid, stage of a star's life, large changes can occur. While our own Sun has another four or five gigayears of life before its core-hydrogen fusion is over, the placid process of early MS evolution will, after just a single gigayear, already drive its luminosity high enough to destroy Earth's biosphere.

At the heart of this early MS process is a gradual change in core composition, as helium ash accumulates. With the core becoming progressively helium richer even while core hydrogen nuclei continue to fuse, the number of particles constituting the aggregate of gas that is the core progressively falls. Given this rise in the mean mass of the core-gas particles (the free electrons, and a diminishing number of hydrogen nuclei, and a rising number of helium nuclei: but the increased helium comes at the expense of the hydrogen, with two hydrogens yielding one helium) the core, while maintaining the pressure needed to support the overlying envelope, is under the Ideal Gas Law forced to contract. Under a physical principle known as the Virial Theorem, half of the gravitational potential energy liberated by the contraction is translated into thermal energy, i.e. into a rise in the temperature of the core. With this rise in temperature, core hydrogen fusion (a process already decidedly dependent on temperature in the case of the proton-proton chain, and very strongly dependent on temperature in the case of the CNO cycle) becomes more vigorous. As a result, the star overall becomes more luminous, and also experiences a modest increase in radius.

It is now convenient to distinguish in our set of 323 MK-classified bright stars between *(A)* the very massive ones (possessing at birth a mass greater than around 8 $M_\odot$ or 10 $M_\odot$) and *(B)* all the others. The very massive stars are destined to die as supernovae (leaving behind perhaps a black hole, perhaps a neutron "star"). The others are destined to die as white dwarfs.

SUBSECTION 5.7: MK phenomenology of evolving high-mass stars (eventual supernovae)

In observational terms, the very massive MS stars are of MK temperature class O, or else of the hot B subdivisions B0, B1, or B2. In our set of 323 MK-classified bright stars, at least the following ten (in order of increasing RA) can be said with confidence to meet this condition: η Ori Aa (B0.5 V), θ Car (B0.5 V), α Cru B (B1 V), β Mus Aa (B2 V), π Sco A (Fang; B1 V), β Sco Aa (Acreb; B0.5 V), τ Sco (Paikauhale; B0 V), ζ Oph (O9.5 V), and α Ara A (B2 V). Additionally, 31 are observed to be on the borderline for meeting this condition (being in IV, or being classified "IV–V," or being of MK temperature class B2.5).

In the process leading up to the supernova climax, these massive stars will eventually rise in observational terms into the MK "supergiant" luminosity class I. In the set of 323, 35 are clearly now at that late stage in their development.

We will not discuss at any length the details of massive-star evolution once core hydrogen is exhausted, instead contenting ourselves with just five brief points:

(i) The very concept of MS is a little misleading for the most extreme of the massive stars, since in the most extreme cases scarcely has starbirth (the commencing of core hydrogen fusion) been achieved before gross observable evolutionary changes have set in. We will not here attempt to chart this territory (and in particular will not attempt to define for this group of stars the tricky theoretical concept of "departure from MS"). We remark only that a safe early life theoretical concept for the most massive stars is the concept of a mere instant, as opposed to an interval—namely arrival on the "Zero Age [Theoretical] MS," as the instant at which core hydrogen fusion starts.

(ii) In their so-short lives, these very massive stars fuse progressively heavier elements, in a central aggregation and in shells overlying the aggregation. The fusion after helium is finished is fuelled first by carbon, then by oxygen, and neon, and magnesium, and finally by sulphur, and silicon, yielding the eventual dumping of iron ash, from sulphur-silicon burning in a shell, onto a growing inert central aggregate of iron.

(iii) A "core-collapse" supernova eventuates after the iron aggregate exceeds the "Chandrasekhar limit" of ~1.4 $M_{\odot}$

(iv) The complexities of core and shell burning, with burning at various levels switching itself on and off in the process leading up to the supernova, translates in observational terms into movements across the MK luminosity-class-vs-temperature-type surface, with luminosity not changing much, but with temperature type changing dramatically (and with changes possible both in the redward, or OBAFGKM, sense and in the blueward, or MKGFABO, sense). Each of the MK types OBAFGKM is represented in our group of 35 supergiants, with at the hot (blue) extreme ζ Pup (Naos; O5 Ia) and ζ Ori Aa (Alnitak; O9.5 Ib), and at the cool (red) extreme α Sco A (Antares; M1.5) and α Ori Aa (Betelgeuse; M2 Iab).

(v) In its redward or blueward progressions, an evolving supergiant can pass, possibly more than once, through the "Instability Strip" (IS) in the luminosity class-vs-temperature type MK plane, thereby temporarily becoming a Cepheid variable. This possibility is presently actualized in our set of 35 class-I stars by (in order of increasing RA) α Umi Aa (Polaris), β Dor, l (ell) Car, η Aql A, and δ Cep A.

SUBSECTION 5.8: MK phenomenology of evolving lower-mass stars (eventual white dwarfs)

(B) We may now proceed to explain the sense in which, extreme cases of lower-mass cases of rotation aside (where rotation yields gas flows so violent as to leave no gas unmixed), all stars in the 323-star set with masses below ~8 $M_{\odot}$ or ~10 $M_{\odot}$, and not disturbed by mass transfer from some companion star, proceed from a readily definable theoretical-MS interval of life to the theoretical Sub-Giant Branch (SGB), then to the theoretical Red Giant Branch (RGB), then to either the theoretical Horizontal Branch (HB) or the theoretical Red Clump, then to the theoretical Asymptotic Giant Branch (AGB), and finally (as almost-corpses or corpses) to a post-theoretical-AGB phase, which, in the fullness of time, yields a white dwarf.

It might seem natural to set up a definition of "theoretical MS" for our eventual-white-dwarf stars on which such a star is deemed to leave the theoretical MS upon finishing core hydrogen fusion. The definition actually employed is, however, different (Carroll-and-Ostlie 2006ima..book.....C, pp. 452, 453). That definition has (surely?) been motivated, over the past few decades of theory construction, by a desire to make the theoretical-astrophysics demarcations correspond as closely as possible to the actual spectrograph-observable changes of direction (i.e. to the actual observed bends) as a star traces its path, over a span of megayears or tens or hundreds or thousands of megayears, on the phenomenological I-through-V vs O-through-M surface. Under the standardly employed definition, a star is said to remain on the theoretical MS not only through the process of luminosity increase attributed in subsection 5.6 to the Ideal Gas Law, but somewhat later, even a little after the depletion of core hydrogen has brought core fusion to a halt.

The matchup of theory and phenomenology is, despite efforts at fine-tuning the theoretical definitions, imperfect. Awkwardly enough, not only can a star be on the theoretical MS even after finishing core-hydrogen fusion: conversely, a star can even have left the observational MS, in other words can have left the MK luminosity class V, while residing so far within the theoretical MS as to be still burning its core hydrogen. In terms of our table, this awkward converse possibility is illustrated by at least the following (in order of increasing RA): χ Car (B3 IV (p?)), λ UMa (Tania Borealis; A1 IV), β Cru A (Mimosa; B0.5 III), ν Cen (B2 IV), ζ Cen (B2.5 IV), ι Lup (B2.5 IVn), α Tel (B3 IV), and the celebrated variable β Cep Aa (Alfirk; B1 III). Additionally, α Lyr A (Vega) is still far within the theoretical MS, and yet might erroneously be thought to have evolved to the edge of the observational MS, since its MK class is A0 Va. Here the cause of the "Va," as distinct from "V," is rotation (with Vega presenting itself to the spectrograph pole-on while rotationally flattened, in other words presenting a misleadingly increased radius).

At the moment at which the depletion of core hydrogen has brought core fusion to a halt, the luminosity of the star derives from fusion in a core-surrounding hydrogen shell, now raised to a fusion-capable temperature by the increased temperature of the inactive helium-ash core. For some modest time after core-hydrogen fusion has ceased, nothing dramatic happens from an observational MK standpoint. Departure from the MS is defined as occurring when the central deposit of non-fusing helium ash becomes so massive as to trigger a rapid internal reorganization of the star, with one or the other of two possible types of rapid contraction, to be distinguished below as "*(B.a)*" and *"(B.b)*." This is the point at which something MK-noteworthy, i.e. something that registers strongly in the spectrograph, finally happens.

(B.a) For stars in the 323-star set of mass below ~1.25 $M_{\odot}$, the growing central deposit of still-inert helium ash becomes so massive as to trigger a further, this time rapid, contraction of the core. Some of the gravitational potential energy present before the abrupt contraction, and now liberated by infall, is under the Virial Theorem translated into an increase in the thermal energy of the shell (in which fusion of hydrogen is therefore in turn speeded up). Paradoxically, although the core has decreased in radius, the rise in temperature of the shell causes the shell to expand, increasing the radius of the star overall.

Two contending factors are now at work. On the one hand, the star has become more luminous. On the other hand, it is now larger. The latter factor outweighs the former, entailing a fall in the photosphere effective temperature. (Total luminous output from the photosphere is determined both by the attained photosphere effective temperature and by the attained photosphere radius, i.e. by the extent of stellar bloat. If the overall radius increase is large, then a reasonable modest increase in total luminous output has to be accompanied by a temperature decrease.)

In MK observational terms, the star, now defined to have departed the theoretical MS and simultaneously arrived on the theoretical SGB, has on the one hand moved some modest distance upward out of luminosity class V, and has on the other hand advanced redward, i.e. has evolved in the sense OBAFGKM.

(B.b) For stars of mass above ~1.25 $M_{\odot}$ (and nevertheless not, we repeat, attaining the ~8 $M_{\odot}$ or ~10 $M_{\odot}$ threshold that makes an eventual supernova possible), the star is found under computer modelling to undergo a more radical internal reorganization. On this more radical scenario, not just the inactive helium-rich core, but the entire star, suffers a rapid contraction. It is this spectrograph-detectable event that is in the "*(B.b)*" case taken to define the end of the theoretical MS phase.

As in the less radical *"(B.a)*" scenario, the star increases in luminosity, with some of the liberated pre-contraction gravitational potential energy once again translated into an increase in temperature (with, once again, a consequent speeding up of hydrogen fusion in the shell). In contrast with the "*(B.a)*" scenario, however, the star is of a reduced radius overall. Under the unavoidable correlation of overall luminous output with both attained photosphere effective temperature and attained photosphere radius, the now shrunken, and yet now brightened, photosphere must now be of a higher temperature. In MK observational terms, the star therefore now quite abruptly not only advances upward in the V–IV–III–II–I sense, but also advances blueward, i.e. evolves in the sense MKGFABO.

Whereas in scenario "*(B.a),*" the star is said to arrive on the theoretical SGB simultaneously with its departing the theoretical MS, in the "*(B.b)*" scenario now under consideration arrival on the theoretical SGB is defined as occurring just a little later than departure from the theoretical MS, with a further episode of core contraction following the overall contraction that under "*(B.b)*" defines departure from the theoretical MS. This further episode of core contraction yields a cooling of the photosphere, and consequently a spectrograph-observable change in the sense OBAFGKM.

In scenario "*(B.a),*" i.e. for stars exceeding ~1.25 $M_{\odot}$, movement through the SGB is rapid, making the detection of such stars statistically improbable, and generating the so-called "Hertzsprung Gap" in HR-diagram plots of same-age stars when the subject population is so selected as to be duly rich in masses exceeding ~1.25 $M_{\odot}$, and duly rich both in observational-MS stars and in observational-RGB stars. (Many open clusters meet this sampling requirement.) The statistical improbability notwithstanding, our 323-star set does succeed in capturing several fleeting residents of the Hertzsprung Gap, at any rate (in order of increasing RA) α Aur Ab (the close Capella companion), ε Leo, ζ Leo A (Adhafera), ο UMa A (Muscida), and ζ Her A.

From this point onward, it is no longer necessary to distinguish scenarios "*(B.a)*" and "*(B.b)*." Under both scenarios, residency on the SGB (admittedly started, as we have just said, in one way in the "*(B.a)*" scenario, in a different way in "*(B.b)*," with residency in the former case brief) in due course yields a cooling of the photosphere. With this cooling, the photosphere opacity rises, causing not only the photosphere-proximate layers but even much of the deeper interior to convect. Since, however, convection is a markedly efficient mode of energy transport, the star becomes progressively more luminous and larger, while keeping its photosphere effective temperature roughly constant. As this observationally dramatic increase in luminosity starts, the star is defined as leaving the theoretical SGB and (simultaneously) arriving on the theoretical RGB.

As in the late phases of theoretical MS life, and as in the theoretical SGB, so also here on the theoretical RGB, the star is fusing hydrogen in a shell overlying an increasingly massive, although still inactive, central ball of helium. Now, however, luminosity is much higher than in the MS and SGB phases. As the still-inactive central helium ball increases in mass, it gradually contracts under its own weight. Some of the gravitational potential energy thus liberated once again becomes thermal energy in the ball, as dictated by the Virial Theorem. With the helium ball now getting gradually hotter, the overlying hydrogen-fusing shell becomes gradually hotter also, producing in turn a gradual speeding-up of its hydrogen fusion, and therefore a gradual increase in the star's (already high) luminosity.

RGB life comes to an end with one of two possible kinds of transition to core helium burning, both entailing a decrease in overall luminosity and yet without much change in photosphere temperature. The transition is violent in the case of the less-massive stars in our set, less violent in the case of the more-massive stars in our set: we again omit details. The core-helium-fusion phase is analogous to, and yet is briefer than, the core-hydrogen fusion that characterizes the earlier part of the theoretical MS. The exact destination of this transition depends on whether the star was at the time of its birth (its arrival on the theoretical MS) metal-poor or metal-rich.

For a star born as metal-poor, exit from the RGB takes it rapidly to the "theoretical HB." This region of the theoretical luminosity-vs-photosphere-effective-temperature plot corresponds to a long, roughly horizontal, roughly straight locus of points, which we might term the "observational HB," on the MK surface. Since globular clusters are metal-poor, the observational HB becomes prominent when a globular is (at least partly) resolved into its constituent stars, for which spectroscopy then yields individual MK types. Different metal-poor stars switching on their core-helium fusion are found to arrive at different points on the observational HB, i.e. to attain different photosphere effective temperatures. The particular attained photosphere effective temperature is found in computer modelling to depend chiefly not on the mass of the newly ignited helium core (this proves on modelling to be rather constant across the metal-poor population), but on the mass of the outer, non-helium, layers.

However, with just two or three or so known exceptions—the most celebrated of these being α Boo (Arcturus)—our 323 MK-classified bright stars are metal-rich. Moreover, the exceptions in our set of 323 are perhaps all at phases of evolution either preceding or following residency on the theoretical HB and observational HB. We will therefore not discuss the HB further.

For a star born as metal-rich, exit from the RGB, i.e. the switching on of core-helium fusion, involves a rapid transition to the theoretical and observational "Red Clump" (in effect the redmost rump of the grander theoretical and observational HB), as further discussed at, e.g. https://en.wikipedia.org/wiki/Red_clump. Since the Red Clump is the helium-fusion analogue of theoretical-MS core-hydrogen fusion, it is unsurprising that it is followed by an evolutionary phase analogous to the observational RGB and theoretical RGB, namely the observational AGB and theoretical AGB.

On the AGB, helium-core fusion has come to an end, with the star at this late stage in its life harbouring an inert core rich in carbon and oxygen. Fusion now proceeds, simultaneously or alternately, in an inner shell of helium and an overlying (and in terms of overall luminous output, for most of the AGB lifetime dominant) shell of hydrogen. With more than one shell in play, evolution becomes rather elaborate. In particular, it is possible for the helium shell to be temporarily inactive, simply accreting mass from the helium ash being dumped on it by the overlying hydrogen shell. Once the helium shell becomes sufficiently massive, it turns on helium fusion, causing the overlying hydrogen shell to expand and briefly switch off. The net result of this is a temporary drop in the luminosity of the star, until the helium burning in turn subsides and the hydrogen burning resumes. In its overall evolution along the AGB, and in its post-AGB transition to the quiet, dead state of a white dwarf, a star can undergo even many tens of such “helium shell flash” episodes. Additionally characteristic of evolution on the AGB are pulsation and mass loss. The possibility is dramatically illustrated in our 323-star set by α Her Aa (Rasalgethi), and with a still higher mass loss by o (omicron) Cet Aa (Mira).

We will skip over the further details of stellar evolution toward the white-dwarf corpse phase, remarking here only that in the case of a star nearly, but not quite, massive enough to die as a supernova, even carbon may be fused before all thermonuclear activity finally ceases.

Two concluding remarks are now in order.

(1) Mention has already been made of “Dredge-Up” as a process affecting the elemental composition of the spectroscopically observed photosphere. In terms of the concepts now laid out, it can be added that “Dredge-Up” may occur in the violent and deep convection of the RGB, as “First Dredge-Up” (FDU), or after the RGB, as “Second Dredge-Up” (SDU) and “Third Dredge-Up” (TDU). A highly evolved star may experience more than one episode of TDU, and it is also possible for FDU and TDU to occur without SDU. Our table cites α Tau A (Aldebaran) as a star that has undergone FDU. On the other hand, our table in its present state of development does not cite instances of SDU or TDU.

(2) The deducing of a star’s evolutionary stage from its observed MK type, as it makes its way off the MS toward, eventually, the AGB, is not always straightforward. In the case of the most massive stars (with masses greater than ~8 $M_\odot$ or ~10 $M_\odot$, and with death-by-supernova therefore impending, and with temperature evolution late in life at one or more stages proceeding in the sense OBAFGKM and at one or more stages proceeding in the contrary sense MKGFABO), temporary observed residence, as a Cepheid variable, on the Instability Strip (IS) raises the question (not always easy to answer) “Is this star making a first, a second, or a third crossing of the IS?” As pointed out in the table “Remarks,” this problem complicates, in particular, the analysis of that rather untidy Cepheid variable that is α UMi Aa (Polaris). For those stars massive enough to achieve core-helium fusion at some point in their lives, and not so massive as to die a supernova death (a condition met in our 323-star set by all the stars below ~8 $M_\odot$ or ~10 $M_\odot$), it sometimes proves difficult to distinguish residency on the theoretical RGB, residency in the theoretical Red Clump, and residency in the theoretical AGB from the available spectroscopy. Indeed the theoretical “Asymptotic Giant Branch” is so named because it corresponds in observational terms to a locus of MK-surface points running perilously close to, so-to-speak, asymptotically approaching, that just slightly redder locus of points that is the observational RGB.

SUBSECTION 5.9: Supplementary remarks on rotation with “Be phenomenon” and “shell” (in MK types O, B, A)

Some of our 323 bright MK-classified B stars have an “e” flag, for emission lines in spectroscopy. Some, and yet not all, such cases involve the important, and not yet well understood, “Be phenomenon.” Strictly speaking, the presently known “Be-phenomenon” stars in our set of 323 are at least the following 19 (in order of increasing RA): γ Cas A, α Eri (Achernar), ε Cas, η Tau Aa (Alcyone), η Ori Aa, ζ Tau (Tianguan), α Col A (Phact), κ CMa, β CMi A (Gomeisa), ω Car, p Car (HR4140), γ UMa A (Phecda), δ Cen Aa, μ Cen Aa, η Cen, δ Sco A (Dschubba), α Ara A, ζ Oph, and β Cep Aa (Alfirk). As we discuss again below, β Lyr Aa1 (Sheliak) may or may not constitute a twentieth case, and some doubt hangs additionally over γ Ara A (in our treatment, not a “Be phenomenon” star, because too evolved; but perhaps we are wrong). Further, closely related to the Be phenomenon is the spectroscopic (predominantly B-star) “shell” phenomenon. The amateur-spectroscopy essay in the Handbook current printed editions notes that the spectroscopic-shell phenomenon, and by implication the Be phenomenon, is a potentially fertile field for amateur spectroscopy. We accordingly supply here a general briefing on the Be phenomenon and its “shell” associate, highlighting the connection of both Be and shell with the often-troubling topic of rotation.

Although many of the most tempting amateur targets in the Be-phenomenon and “shell” fields are members of our 323-star set, we nevertheless discuss the Be and shell phenomena for the most part in general terms, without restriction to the set of 323. We hope thereby to maximize the value of our discussion, and in particular to stimulate an interest in Pleione, as a Be and sometimes-“shell spectrum” star not much fainter than our mag. ~3.55 cutoff.

Of all the non-cluster B stars in the galaxy, about 17% at some point in their lives present the “Be phenomenon,” with the phenomenon more prevalent at the hotter (near-O) than at the cooler (near-A) end of the B range. Within the overall set of galactic stars, the exceeding rare O stars are known to sometimes present the same phenomenon (with the term “Oe” star therefore used occasionally in the literature). In our set of 323, ζ Oph, as an O star with a photosphere almost, and yet not quite, cool enough to entail classification as a hot B, is an instance. Also within the overall set of galaxy stars, some A stars are known to present the Be phenomenon. Again, our 323-star set furnishes an instance, namely γ UMa A (Phecda): this star is of MK temperature type A0, and so is just barely cool enough not to fall into the B classification bin. Nevertheless, since the phenomenon (which we will soon describe in proper detail) occurs predominantly in the B stars, the term “Be phenomenon” is standardly applied to stars in all three of the O, B, and A observational MK temperature types.

Several qualifying comments are now necessary.

The Be-phenomenon stars are not to be confused with the “Herbig Ae/Be ‘stars’.” The latter are not stars in the strict sense, but instead are contracting starlike bodies that have not yet achieved starbirth, i.e. have not yet started core hydrogen fusion. In their present stage of development, they are continuing to heat up under gravitational contraction, and are (unsurprisingly for objects condensing out of ISM clouds) embedded in circumstellar dust.

A true “Be phenomenon star” need not currently have emission lines in its spectrum. It must, on the other hand, be known to have at some point in its past presented emission. In observational practice, the emission is always found to occur in at least one or more lines of the hydrogen Balmer series.

The condition of past-or-present emission, while necessary, is not in its turn sufficient. A supergiant in MK type B, with Balmer emission, is not a Be-phenomenon star. For a star to be Be-phenomenon, it must either lie on the theoretical MS or (as in the case of Be-phenomenon ζ Tau (Tianguan) in our table, observationally in MK luminosity class III) be evolved only modestly beyond the theoretical MS.

Also not harbouring a Be-phenomenon star is a theoretical-MS or near-theoretical-MS member of a binary system with mass transfer, in which the observed hydrogen Balmer emission comes from an incandescent mass-transfer stream. In the table, this is perhaps the case for β Lyr Aa1 (Sheliak), which certainly has such a mass-transfer stream. Confusingly, however, a “shell” spectrum is observed for Sheliak, and “shell” in the case of a young B star (as we explain below) is generally, or even inevitably, associated with the Be phenomenon. Perhaps all that can be said here is that Sheliak is a confusing case. (It has certainly been notorious over the decades, in one way or another, as a challenge to modelling.) The conceptual point remains that if, hypothetically speaking, emission in a young B star were to come from no source other than a mass-transfer stream, thanks to that star’s membership in a tight binary system, then that star, while being obliged to show the observation-driven “e” flag in its MK type, would not count as an instance of the Be phenomenon

This, then, concludes the qualifying comments. To recapitulate: the true Be-phenomenon stars are theoretical-MS or near-theoretical-MS stars with presently observed or historically observed emission lines, where the emission is not due to a mere mass-transfer process attributable to membership in a tight binary system.

The astrophysical task is now to determine what produces the emission. Emission must mean that the star has somehow managed to shed significant quantities of incandescent gas. Copious shedding cannot be attributed to stellar winds, since winds play only a minor role in mass-shedding for stars within or near the MS (except, perhaps, for the case of stars at the hottest end of the O range, where even the concept of time-spent-on-theoretical-MS is, as noted above, problematic). Our own Sun, for instance, as an MS star, sheds a mere tenth-of-a-trillionth of its mass per year.

The cause of the copious shedding has not yet been determined with confidence. It is possible that all Be stars are rapid rotators (although, as we remarked in subsection 5.4, spectroscopy, with its incorporation of “n” or “nn,” as occasionally appropriate, in an MK type, cannot by itself detect rotation when the star is oriented pole-on to Earth). On the other hand, there are many rapidly rotating theoretical-MS or near-theoretical-MS O, B, and A stars that do not present the Be phenomenon.

The following picture therefore suggests itself: if the star is a rapid rotator, and in addition possesses some mechanism “X” for launching photosphere gas from near its equator into its equatorial plane, then an incandescent disk forms, girdling the star, and registering as emission at the spectrograph. With the star a rapid rotator, it will not be a sphere but a rotationally somewhat flattened object, with local gravity in the photosphere somewhat lower at the equator than at the poles, and with launching into an equator-plane orbit consequently favoured. The observed hydrogen Balmer emission is on this picture a signature of hydrogen ionization in the disk, under a violent barrage of UV from the (hot, as O-or-B-or-A) photosphere: Balmer-lines hydrogen light is emitted as part of the process in which free electrons and hydrogen nuclei recombine, where a captured electron falls to the penultimate energy level from some higher level.

The equatorial-disk picture was first proposed in 1931. Now quite widely accepted is a “Viscous Decretion Disk” elaboration of this idea, introduced in 1991MNRAS.250..432L. “Decretion” proves a useful contrived astronomical term, created as an antonym for “accretion.” Accretion disks figure in various astrophysics contexts, for instance in such black-hole binaries as Cyg X-1 (material shed by the readily amateur-visible member of this binary falls first onto an accretion disk around the black-hole event horizon), and again in the case of starbirth, where material from the gestating ISM cloud forms an accretion disk around the protostar, in a process that might see the disk eventually transform itself into a bevy of exoplanets, with perhaps also a belt of small rocky asteroid-like bodies, and with some analogue of our Solar System’s zodiacal dust, all orbiting an infant star. Correspondingly, a “decretion disk” forms when an astronomical object (in our case the Be-phenomenon star) for one reason or another releases matter into orbit in its neighbourhood.

Although the dimensions of the hypothesized disk are not easily investigated, emission in the Balmer hydrogen-α line in the cases so far studied has been found to come from a disk on the order of 0.3 AU to 0.6 AU in radius. We seem to have here, in other words, one of the grandest of all theoretical-MS or near-theoretical-MS stellar spectacles.

Unfortunately, it is a spectacle that at best can be imaged only fuzzily, even with the most capable current optical interferometers. Let the Jupiter disk, of diameter ~50″, familiar from the small telescope, become a circular tea-tray 50 cm in diameter. The binaries resolvable in good seeing by the small telescope, at a separation of ~1″, thereby become a pair of points on that tray lying 1 cm apart. The most celebrated of the Be-phenomenon stars, γ Cas A, already noted as spectroscopically peculiar by the first stellar spectroscopist, Fr Angelo Secchi, in or shortly before 1866, lies at a distance of 600 ly from Earth. A disk of incandescent gas on the order of 0.5 AU in radius, or 1 AU in diameter, is seen at this distance as an object a mere 5 mas across. In terms of the tea-tray, this corresponds to an object around 50 microns wide, in other words to an object having the approximate width of a human hair.Consistent with the picture of gases launched by “Mechanism X” into circumstellar orbit is the 2007A&A...464...59M discovery that the gas in Be-phenomenon star α Ara A is in a normal central-gravitational-field (i.e. “Keplerian”) orbit, moving unconstrained by any such nongravitational forces as magnetism, and not possessing the kinetics of a mere stellar wind.

What, then, can “Mechanism X” be? It is possible that different Be-phenomenon stars have different gas-launching mechanisms. Outflows from the poles are not currently considered relevant. Nonradial pulsation, on the other hand, may play a role in at least some cases, as may also local magnetic phenomena at the low latitudes. (There is perhaps no known case of a Be-phenomenon star with a strong global magnetic field.) Helpfully, all hitherto scrutinized Be-phenomenon stars have been found to be pulsating variables, although in some cases the pulsation-produced luminosity variation is at the millimagnitude level or below, eluding detection by ground-based photometry. (In addition to facing possible very-low-amplitude variations, photometric monitoring of the stellar pulsation is confronted by the complication that the disk itself may vary photometrically (possibly with high amplitude).)

Nonradial pulsation aside, it is possible that in some cases, where the Be star is a member of a binary with tight orbit, or at any rate with an orbit possessing a tight periastron, the “X” role is played by the perturbing gravitational field of the companion.

Some Be-phenomenon stars have emission (from, on the currently accepted modelling, equatorial disks) which is, so far as the existing multidecade observational record goes, stable. Other Be-phenomenon stars, however, present emission lines only intermittently, in their years or decades of “outburst.” Two prominent instances of outburst-and-quiescence in our 323-star set are the already-cited γ Cas A and the recently active δ Sco A (Dschubba). Another well-known instance, although a little too dim for inclusion in the 323-star set, and sharing the notoriety of bright γ Cas A, is Pleione. This star, easy in binoculars as the northern neighbour of Atlas at the eastern extremity of the Pleiades, presented an emission-line outburst of uncertain commencement extending to 1903, and presented additional emission-line outbursts in the periods 1955–1972 and 1989–2005.

Where the disk is permanent, the “X” mechanism works steadily to launch fresh consignments of photospheric gas into orbit, i.e. to perpetuate the decretion. The ongoing launch compensates for the ongoing accretion of matter from at least the inner part of the disk back onto the photosphere. If the mechanism should for some reason cease to operate, decretion ceases, and yet accretion continues. This has the consequence that the disk vanishes (with, however, some of the outlying parts of the disk lost not to accretion onto the photosphere but to outflow, into the embedding ISM).

On some current modelling, a typical Be disk increases in thickness rather gently as one progresses outward (with radially directed tangents to the disk, as taken at the points where disk meets photosphere, yielding a tight “full-opening angle” of ~10˚). A further geometrical detail from some current modelling may also be noted: if, as is often the case, the Be-phenomenon star is a member of a binary not tight enough to produce mass transfer, and yet tight enough to produce a gravitational perturbation from the companion star, and if the Be-phenomenon star equatorial plane diverges somewhat from the orbital plane of the binary system, then the disk is warped.

We may now turn from Be to the related “shell spectrum” phenomenon. The term is somewhat unfortunate, being perhaps a relic from discussions in the early 20th century, when it was perhaps thought that an O or B or A star in or near luminosity class V could under the right circumstances surround itself not with an equatorial disk of gaseous ejecta (as on the currently accepted modelling) but with a literal “shell” of gaseous ejecta, in other words with an enclosing blanket. For better or worse, the term has stuck, surviving the acceptance of the disk morphology (and has nothing to do with thermonuclear-fusion shells in stellar interiors, as discussed in subsections 5.6, 5.7, and 5.8 of this essay). A “shell” spectrum in a rapid rotator, oriented equator-on to Earth, occurs when some lines are seen not in the expected broadened absorption typical of an equator-on rapid photosphere, but in, or also in, narrow absorption. Typically, though not inevitably, the unexpected narrow absorption lines occur as narrow absorption cores within Balmer emission.

On the current understanding, “shell” in this sense typically results when a Be star not only generates its (perhaps temporary) disk of equatorial ejecta but happens to be oriented more or less equator-on in relation to the spectrograph. Under these circumstances, part of the disk lies between photosphere and spectrograph, yielding the absorption. Since this part of the disk is moving more or less orthogonally to the line of sight, i.e. is neither approaching the spectrograph nor receding, its absorption lines escape the rotational broadening characteristic of absorption lines from the photosphere.

Although a Be-phenomenon star with equator-on orientation can, as just noted, be simultaneously in emission-line outburst and in “shell,” it sometimes happens that shell absorption is present in a Be-phenomenon star even after its emission has for the time being subsided. The Be-phenomenon star Pleione, in particular, had a shell spectrum without emission in the period 1938–1954, and then again for some years after 1973.

Rapid rotators fitting the definition of “shell spectrum” occur even somewhat outside our present domain of interest, the Be-phenomenon stars, with instances known even in type F, right down to the F5 “rotation break.” It remains the case, however, that “shell” is most prominently connected with the Be, as a phenomenon contemporaneous with a Be outburst or present in a star that at some earlier or later time is observed to be in Be outburst.

What, in this general Be-cum-”shell” field, are the possible lines of activity for the amateur spectroscopist?

On the humblest level (even with a visual spectroscope and no camera, as in the case of 1860’s Fr Angelo Secchi), it is possible to monitor theoretical-MS or near-theoretical-MS rapid rotators, to see whether emission is currently present or currently absent. The sudden onset of emission would be newsworthy of communication to AAVSO, to the LESIA laboratory at Paris-Meudon (as mentioned again below), or to other appropriate pro-am authorities.

On a less humble level, where spectrograms are taken, and are converted into intensity-against-wavelength plots, or "extracted one-dimensional spectra," with such professional astrophysical tools as IRAF, the evolution of emission-line and shell-absorption-line profiles could be tracked. In particular, where shell absorption is present simultaneously with emission, as in the (conveniently strong) hydrogen Balmer lines, duly equipped amateurs could examine from month to month whether emission is currently stronger on the violet, or on the contrary red side of the partitioning absorption.

Finally, we suggest in a speculative spirit that it might prove possible to keep a month-upon-month polarimetry log (although we do not ourselves know whether any amateurs in any country have attempted polarimetry, whether in a Be-phenomenon context or in other contexts): if the Be-phenomenon star is not seen pole-on, then some light from its photosphere will be scattered toward the polarimeter by free electrons in the disk and will therefore be linearly polarized.

The recent literature includes a long review article, 2013A&ARv..21...69R, on the Be phenomenon. The IAU Working Group on Active B Stars (a group whose domain of interest includes, and yet is not confined to, the Be and shell phenomena) has a homepage at activebestars.iag.usp.br/bstars, with a link to its newsletter materials, including a newsletter archive. The LESIA laboratory at the Observatoire de Paris-Meudon maintains the "BeSS Database" comprising Be-phenomenon stars, the Herbig Ae/Be "stars" briefly mentioned near the beginning of this subsection, and a "B[e]" category of supergiants, at basebe.obspm.fr/basebe.

APPENDIX: Glossary of acronyms and similar designation

The following is a glossary of the acronyms and similar designations used in our essay and table. We omit, as sufficiently obvious, a small handful of universally known acronyms (e.g. NASA), designations of chemical elements and chemical compounds (e.g. CO, for carbon monoxide. We do include a few designations of particular satellites or similar space missions (e.g. BRITE, MOST).

- AAT: Anglo-Australian Telescope (3.9 m, Siding Spring Mountain, New South Wales, Australia)
- AAVSO: American Association of Variable Star Observers
- AAVSO(VSX): AAVSO International Variable Star Index (www.aavso.org/vsx)
- ALMA: internationally funded Chile-based radio interferometer ("Atacama Large Millimetre/submillimetre Array")
- AMBER: spectro-interferometer at VLT ("Astronomical Multi-BEam combineR")
- AGB: asymptotic giant branch (as a region in the two-dimensional MK luminosity-versus-temperature stellar classification space)
- Astron. Alm.: *The Astronomical Almanac*, as the joint annual publication, in print and to a reduced extent online, of the United States Naval Observatory and HM Nautical Almanac Office; "Section H" (not necessarily always up to date in the online version) provides V magnitudes, B–V and V–I colours, and MK types for several hundred bright stars; Astron. Alm. particulars can be had from asa.hmnao.com/ and aa.usno.navy.mil/publications/asa
- AU: astronomical unit (the formal 2012 IAU definition is in effect a precisification, in the (SI) laboratory unit of metres, of the earlier epoch-of-Kepler AU concept; before 2012 the concept was defined in astronomical, as distinct from laboratory, terms–before 1976 as the half the sum of the Earth-to-Sun distance at perihelion and the Earth-to-Sun distance at aphelion, and from 1976 onward with a gravitation-theory precisification of that half-of-sum concept)
- BRITE: BRIte Target Explorer, a.k.a. Canadian Advanced Nanospace eXperiment 3 (CanX-3:

constellation of precision-astrometry satellites (6 attempted, 5 successfully deployed), as a Canada-Austria-Poland collaboration; first launch was in 2013)

- BSC5: Yale Bright Star Catalog, Version 5
- BSG: blue supergiant
- BTA-6: Bolshoi Teleskop Alt-azimutalnyi-6 *(*Большой Телескоп Альт-азимутальный-6, "Large Alt-Azimuth Telescope 6": 6-m telescope on north side of Caucasus Mountains, Russia)
- CADARS: Catalogue of Absolute Diameters and Apparent Radii of Stars (doi.org/10.1051/0004-6361:20000451)
- CHARA: the Mount Wilson optical interferometer (Center for High Angular Resolution Astronomy)
- CME: coronal mass ejection
- CNO cycle: the carbon-nitrogen-oxygen-catalyzed cycle under which the hotter stars fuse hydrogen into helium
- COAST: the Cambridge optical interferometer (Cambridge Optical Aperture Synthesis Telescope)
- CODEX: a series of computer codes for the numerical simulation of stellar atmospheres (Cool Opacity-sampling Dynamic EXtended)
- CORIOLIS: USA-but-not-NASA satellite launched 2003; mission involves not only instrumentation for Earth ocean-environs monitoring, but also solar-wind monitor SMEI (Solar Mass Ejection Imager)
- DAO: Dominion Astrophysical Observatory (on Vancouver Island, British Columbia)
- DR2: Data Release 2 (at *Gaia*)
- ESA: European Space Agency
- ESO: European Southern Observatory (multiple sites, in northern Chile)
- FDU: First Dredge-Up (as a stage in stellar evolution, soon after a star evolves out of the MS)
- FUV: far ultraviolet
- GCVS: General Catalogue of Variable Stars (Sternberg Astronomical Institute, Moscow)
- GTR: general theory of relativity
- Hp: a visible-light passband used for photometry at *HIPPARCOS*
- HM Nautical: "Her/His Majesty's Nautical" (for UK publications and UK agencies)
- HR diagram, HR plot: two-dimensional luminosity-versus-temperature plot for the members of some given population of stars; it is useful to distinguish the "observational" (phenomenological, MK-classification) and the "theoretical" HR diagrams
- HST: *Hubble Space Telescope*
- IAU: International Astronomical Union (Paris)
- IR: infrared
- IRAF: Image Reduction and Analysis Facility: a suite of software tools, for astronomical tasks including aperture photometry and the "extraction of one-dimensional spectra" from raw spectrograms, available free of charge from the National Optical Astronomy Observatory (USA); very widely used at North American professional observatories, and quite widely also, but in competition with MIDAS, at professional observatories outside North America: ast.noao.edu/data/software
- IS: Instability Strip (as a region in the two-dimensional luminosity-versus-temperature stellar-classification space)

- ISM: interstellar medium
- *IUE*: International Ultraviolet Explorer (space telescope: NASA, ESA, and United Kingdom; 1978–1996)
- LESIA: Laboratoire d'Études Spatiales et d'Instrumentation en Astrophysique (physically at Paris-Meudon): lesia.obspm.fr/
- LPV: long-period variable
- LSR: Local Standard of Rest (as reference frame for kinematics of bodies in our own galaxy)
- $M_{\odot}$: solar mass
- mas: milliarcsecond
- MK: Morgan-Keenan (two-dimensional phenomenological, non-theoretical, stellar classification scheme, with "MK luminosity classes" and "MK temperature types")
- MOST (*Microvariability and Oscillations of Stars/Microvariabilité et Oscillations STellaire*): Canadian space telescope for precision photometry, launched 2003
- MS: main sequence (as a region in the two-dimensional luminosity-versus-temperature stellar classification space; it is useful to distinguish the "observational MS," in other words the empirical MK luminosity class V, from the "theoretical MS")
- My: megayears
- NCP: North Celestial Pole
- NPOI: a US Naval Observatory facility (Navy Precision Optical Interferometer)
- NSV: New Catalogue of Suspected Variable Stars (Sternberg Astronomical Institute, Moscow)
- OBAFGKMLTY: the temperature-ordered sequence of MK types, with O the hottest and Y the coolest; until the discovery of brown dwarfs, in types L, T, and (very recently) Y, the sequence was simply OBAFGKM, recalled by 20th-century students with the unfortunate mnemonic "Oh Be A Fine Girl Kiss Me" (implementing gender-neutrality, and allowing for the three progressively cooler brown-dwarf types, one might instead propose "Oh Be A Fine Gymnast, Kiss Me Like This, Yowee"); outside this sequence are the special MK labels (marking gross chemical anomalies) W (for the Wolf-Rayet stars; these turn out to be hot, like O stars), C (for stars whose photospheres are rich in carbon; these turn out to be cool, like K or M) and S (for stars with chemically anomalous photospheres, these are in terms of spectral phenomenology intermediate between M and C, and turn out to be cool); C is the current label for a group that was in earlier decades divided into R and N: additionally, the special "D" and "P" flags are used, in a more colloquial MK spirit, for planetary nebulae hosts and white dwarf "stars"
- OHP: Observatoire de Haute-Provence (France)
- PA: position angle
- PTI: Palomar Testbed Interferometer
- $R_{\odot}$: solar radius
- R*: stellar radius (with reference to some given, reasonably spherical, star)
- R_{eq}: equatorial radius (with reference to some given rotationally flattened star)
- RGB: red-giant branch (as a region in the two-dimensional luminosity-versus-temperature stellar classification space)
- R_{pol}: polar radius (with reference to some given rotationally flattened star)

- RSG: red supergiant
- SAAO: South African Astronomical Observatory
- SB: spectral binary, whether double-lined or single-lined
- SB2: double-lined spectral binary
- SETI: search for extraterrestrial intelligence
- SI: Système International d'Unités; the internationally agreed system of second-metre-kilogram-ampere-kelvin-mole-candela laboratory units, at one time implemented with recourse to some physical artefacts (including most notoriously the "standard kilogram", in a vault at the Bureau International des Poids et Mésures in France), but since a 2019 decision defined in a way which can be reproduced by any duly equipped laboratory, independently of artefacts; SI, in its various iterations through the decades, is a 1960 precisification of the earlier internationally agreed "MKS system", from 1889
- SGB: sub-giant branch (as a region in the two-dimensional luminosity-versus-temperature stellar classification space)
- SMEI: Solar Mass Ejection Imager, as an instrument on the CORIOLIS satellite
- SN: supernova
- SNR: supernova remnant
- STIS: Space Telescope Imaging Spectrograph System (an instrument on HST)
- SWB: stellar-wind bubble
- UV: ultraviolet
- V: the visible-light passband in the UBVRI photometric passband system that best approximates the response of the human eye, as lying between the blue ("B") and red ("R") visible-light passbands
- VLT: a Chile-based facility of the European Southern Observatory (Very Large Telescope)
- VLTI: the interferometer at VLT
- VSX: AAVSO International Variable Star Index (www.aavso.org/vsx)
- WFC3: Wide Field Camera 3 (an instrument on HST)
- WFPC2: Wide Field and Planetary Camera 2 (an instrument on HST)
- WD: white dwarf
- WIRE: *Wide-field Infrared Explorer* (a.k.a. Explorer 75, a.k.a. SMEX-5); a NASA space telescope, 1999–2000
- WR: Wolf-Rayet (as a type of star)
- WDS: Washington Double Star Catalog: www.usno.navy.mil/USNO/astrometry/optical-IR-prod/wds/WDS
- ZAMS: zero-age main sequence (the subregion of the MS comprising stars that have just begun stable core-hydrogen fusion)

<u>History of recent revisions to both essay and table</u>

Recent revisions are tracked with UTC YYYYMMDDThhmmssZ timestamping, in the "major.minor.patch" version-numbering scheme common in software development.

- 20220313T223945Z/7.0.3: Added Title, page numbering, and fixed some tab errors
- 20220312T171817Z/7.0.2: Minor editing and formatting, particularly as regards subsection references in the essay, colour of URLs, italics or no-italics on names of space missions (mostly other than HST).
- 20220311T033032Z/7.0.1: Significant editing and formatting, particularly tabs and fitting long remarks into available space; corrected a few spelling errors. Removed all instances of "http://" as these are redundant in a URL; also removed italics on HST (only use italics on full text, i.e. *Hubble Space Telescope*).
- 20220303T210237Z/7.0.0: Performed sufficient updating of the 5.x.x version series to support not only the print edition of the 2022 Handbook, but also to support uploading to the online-version server. With the now-noted binarity of ζ CMa Aa,Ab, the count of Sample S (in the essay) was increased from 322 stars to 323 stars. Also in the essay, the former Section 4 was reassigned as Section 5, and a rather long Section 4 was inserted, discussing the astrophysics of binaries. Concomitantly with this addition of this essay section, astrometric detail was added, under "Remarks," to perhaps roughly one out of every four table entries, in a general review of the treatment of binaries, and underlining was added in the leftmost column of the table to flag cases in which a binary system possesses a published orbital solution. Apart from many small revisions in the table, extended "Remarks" treatments were inserted for two stars of special interest, δ Sco A (amateur photometry and amateur spectroscopy is particularly needed during and around the time of periastron passage, in 2022 May) and α Her Aa.
- 20210811T201642Z/6.0.0: Performed sufficient updating of the 5.x.x version series to support the print edition of the 2022 Handbook, but without sufficient updating to support uploading to the online-version server.
- 20210807T203107Z/5.2.0: Made various copy-edit corrections (such as insertion of missing punctuation, correction
  of a few spelling errors), and additionally on the side of scientific substance made a few corrections or amplifications (chiefly as follows: amplified the essay elucidation of "n", "nn" in MK types; improved an essay remark on rotation in stellar evolution; made essay correction regarding protracted-versus-brief membership of SGB; corrected essay list of Be-phenomenon stars (the phenomenon is not observed in Adhara); added "SGB" to glossary of acronyms; improved table discussion of exoplanet status for α Tau A (Aldebaran); corrected table magnitude range for α Ori Aa (Betelgeuse); corrected table typo for angular distance in α Cru AB (Acrux and companion; correct value is 3.5″, not 35″); corrected table typo for magnitude of η Oph B (correct value is 3.3, not 7.3); updated α PsA Aa (Fomalhaut) table entry to reflect the fact that HST-imaged "exoplanet" Dagon (2008) has now faded below the imaging threshold, and is therefore now believed to be an expanding, and therefore an increasingly tenuous, debris cloud rather than a true exoplanet); this version is a supplement to the 2021 Handbook, with the upcoming 6.x.x series intended to support instead the 2022 Handbook
- 20210217T042710Z/5.1.1: Made minor adjustments to tabs and spacing for paragraphs before creating online PDF.
- 20210216T161213Z/5.1.0: Made minor adjustments (small points of syntax, spelling, punctuation, or similar, with much bibcode error correction). Added a long paragraph with five methods for retrieving a full-text, all-illustrations PDF from a typical astronomical bibcode citation. Corrected a mistake of astrophysical substance, in the subsection 4.8 discussion of onset-of-helium-core-fusion (violence in the onset of core-helium fusion is characteristic of the less massive, not of the more massive, incipient fusers-of-core-helium) This yielded a work sufficiently updated to support uploading to the online-version server.
- 20210128T145046Z/5.0.0: Made major revisions of the 4.0.0 version series, by adding several

thousand words to the introductory online essay, with stellar-evolution background and a detailed briefing on the amateur-relevant “Be phenomenon” and “shell spectra” (and to a lesser extent by expanding “Remarks,” most notably for α Eri (Achernar), ζ Tau (Tianguan), and α Aql A (Altair); other work on “Remarks” included routine updates for such things as binary position angles and celestial-sphere distances, and also comparison of our MK types against MK types as assigned by Astron. Alm. for epoch 2021.5, with the MK discrepancies logged). The work was not yet sufficiently polished to support uploading to the online-version server.

- 20200815T190800Z/4.0.0: Performed sufficient updating of the 3.x.x version series to support the print edition of the 2021 Handbook, but without sufficient updating to support uploading to the online-version server.
- 20191231T235959Z~/3.x.x series: Supplemented previous editions of this online publication in various ways, most notably by adding the (rather prolix) results of (rather detailed) primary-literature inspections for o Cet Aa (Mira), α Umi Aa (Polaris), β Per Aa1 (Algol), α Tau A (Aldebaran), ε Aur A (Almaaz), α Ori Aa (Betelgeuse), γ Vel Aa, α Leo A (Regulus), α Vir Aa (Spica), ζ Oph, and α Lyr A (Vega).
- 20181231T235959Z~/2.x.x series: Supplemented the 1.x.x version series with some (rather detailed) primary-literature inspections for selected familiar bright stars, thereby expanding “Remarks.”

	Star Name	RA (2022.5) h m	Dec ° ′	m_V	$B-V$	MK Type	π mas	M_V	D ly	μ ″/y	PA °	V_{rad} km/s	Remarks	
	Sun			−26.75	0.63	G2 V		4.8	8 lm					
α	And Aa	0 09.6	+29 13	2.07	−0.04	B9p IV: (HgMn)	34	−0.3	97	0.214	140	−12 SB2†	the SB2 components α And Aa, α And Ab (period 96.7 d) are now interferometrically measured, yielding orbital value e = 0.5 or 0.6	**Alpheratz**
β	Cas A	0 10.4	+59 16	2.28†	0.38	F2 III	60	1.2	55	0.554	109	+12 SB†	slight var.: 2.25–2.29 in V, 0.1010 d second-brightest of the δ Sct variables (the brightest is α Aql A (Altair)) ¶ rapid rotator: 2011ApJ...732...68C finds the rotation to be > 90% of breakup rate, and radius at poles to be ~24% less than radius at equator, with β Cas A of mass ~2 $M_{\odot}$, seen nearly pole-on; β Cas A is notable for being cooler than typical rapid rotators, lying just barely on the rapid side of the F5 “rotation break,” and additionally is notable for being old enough to have evolved off the MS, having in its MS career been an A star rather than an F star (generally, rotation slows as an aging star increases in radius: but our table of bright stars does harbour at least one other such evolved rapid rotator in type F, namely θ Sco A); an envelope at this modest photospheric temperature is dominated by convection not only at the equator but even at the (~1000 K hotter) poles; consistently with this picture of an envelope everywhere convective, interferometry of β Cas A is found to yield results for low latitudes gravity darkening inconsistent with 1920’s von Zeipel law (the law is accurate only if an envelope is radiative); 2011ApJ...732...68C suggests that in its process of evolution off the MS (in which a core contracts, an envelope expands) β Cas A has been efficient in transferring angular momentum from core to envelope ¶ 2011ApJ...732...68C Fig. 4 presents imaging (as a single star, not as a binary) of β Cas A, from CHARA interferometry ¶ β Cas A is SB, of period 27 d, as yet unresolved, even in interferometry (so WDS is as yet unable	Caph

													to write "β Cas Ba", "β Cas Bb"); NPOI interferometry has, however, succeeded in assigning a limb-darkened diameter, ~2.1 mas	
γ	Peg A	0 14.4	+15 19	2.83†	−0.19	B2 IV	8	−2.6	400	0.009	168	+4 SB	slight var., β Cep type, 2.82–2.86 in V, 0.1518 d ¶ E(B–V) =+0.01	Algenib
β	Hyi†	0 26.9	−77 08	2.82	0.62	G1 IV	134.1	3.5	24.3	2.243†	82	+23†	possible exoplanet ¶ high space velocity (interloper from more remote galactic region?)	
α	Phe	0 27.4	−42 11	2.40	1.08	K0 IIIb	38.5	0.3	~85	0.426	147	+75 SB		Ankaa
δ	And Aa	0 40.5	+30 59	3.27	1.27	K3 III	~30.9	0.7	106	0.142	126	−7 SB†	one of the rare instances in which a resolved SB has been resolved with direct imaging, rather than with interferometry (although angular diameter of δ And Aa has been found through interferometry): 2015ApJ...809...11B reports a single 2014 observation via Palomar (5 m) Stellar Double Coronagraph, working in near IR with adaptive optics; angular separation was 0.4″, yielding Aa-to-Ab distance ~12 AU, with *e*~0.5 ; since system is large, period is long (~20000 d?) ¶ there seem (as of 2021 Sept.) to be no reports of interferometric resolution of this SB ¶ δ And Ab (mag. 10.0) is asserted in 2015ApJ...809...11BWD to be likely not a WD, as had been previously believed, but instead likely an MS star of MK type K (with the authors additionally noting, as one alternative possibility (in this case a remote alternative possibility), that δ And Ab might be a pair of low-mass MS stars)	
α	Cas A	0 41.8	+56 40	2.24†	1.17	K0 IIIa	~14.3	−2.0	230	0.060	122	−4 V?	AAVSO(VSX) millimag. variable (from starspots) ¶ limb darkening observed interferometrically (disk diameter 5.25 mas) ¶ α Cas D (mag. 9), PA 283°, 70″ (2018), is mere line-of-of-sight coincidence, not gravitationally bound to α Cas A	Schedar
β	Cet	0 44.7	−17 52	2.04	1.02	K0 III†	~33.9	−0.3	96	0.235	82	+13 V?	anomalous in being X-ray-bright and yet a slow rotator ¶ evolutionary status uncertain (helium core ignited already, or still contracting?) ¶ Astron. Alm. (epoch 2021.5) assigns MK type G9 III CH–1 CN 0.5 Ca 1	Diphda
η	Cas A†	0 50.5	+57 56	3.46	0.59	G0 V†	168	4.6	19.4	1.222	117	+9 SB?	B:7.51, K4 Ve, 13.4″, PA 62°→326°, 1779→2019 orbit 480 y; SB status has been asserted for η Cas A, and yet is said by WDS (as viewed 2021 Sept. 14) to be not confirmed; overall field is crowded with dim stars, the brightest of which, apart from η Cas A and η Cas B, are the well-separated η Cas G (mag. 9.5; 420″, PA 259° (2012); rectilinear-solution analysis of proper motions, 1852→2012, does not reveal any orbital motion) and the very widely separated η Cas H (mag. 8.5; 701″, PA 355° (2012); analysis of proper motions covers only 1991→2012, and rectilinear-solution analysis of proper motions seems unavailable as of at any rate 2021 Sept. 14) ¶ Astron. Alm. (epoch 2021.5) assigns MK type F9 V	Achird
γ	Cas A†	0 58.1	+60 50	2.15v†	−0.05	B0 IVnpe (shell) †	5	−4.2	600	0.026	98	−7 SB	var.:1.6–3.0 (V); B:10.9, 2.1″, PA 255°→259°, 1888→2002 orbit >1500 y ¶ first "Be phenomenon" discovery (Secchi, 1866); additionally the prototype for the γ Cas type of eruptive irregular variables; background on Be phenomena and γ Cas-type variability is given in www.aavso.org/vsots_gammacas; 2002ASPC..279..221H summarizes the observational history, including major shell-spectrum phases in 1935–1936 and 1939–1940; despite its historical importance, however, Cas A cannot safely be taken as a typical "Be phenomenon" star, since it presents the peculiarity of hard thermal X-ray emission (cf 2013A&ARv..21...69R p. 42, and also e.g. 2012A&A...537A..59N), derived from magnetic heating (perhaps from magnetic star-disk interaction, perhaps from disk intrinsic magnetic	

													field); rotationally flattened (period = 1.21 d, axial tilt=45°); one of only three Be-phenomenon stars so far observed (via polarimetry, not via interferometry) to produce ejecta disks with differing position angles at different outbursts (2013A&ARv..21...69R p. 42; the other two known instances of this geometrical variation are Pleione and 59 Cyg, both too faint to be in this Handbook table of brightest stars) ¶ as of at least 2007, AAVSO has called for amateur assistance with photometry: γ Cas A has been as bright as V mag. 1.6, as faint as V mag. 3; two recent AAVSO reports, from the same observer, working in the V band, are 2.16 (2022 Feb. 13) and 2.20 (2022 Jan. 10) ¶ dimming through ISM dust, ~0.35 mag.	
β	Phe AB†	1 07.1	−46 36	3.32	0.88	G8 III + G8 III	16	0.3:	~180	0.088	293	−1	AB similar, 0.6″, PA 26°→76°, 1891→2018 orbit 168 y, highly eccentric; masses nearly equal; A is mag. 4.1 and B is mag. 4.2	
η	Cet A+2P	1 09.7	−10 04	3.46	1.16	K1.5 III CN1†	26.3	0.6	124	0.257	123	+12V	¶ Astron. Alm. (epoch 2021.5) assigns MK type K2^{-} III CN 0.5	
β	And A	1 11.0	+35 44	2.07†	1.58	M0 IIIa†	17	−1.8	200	0.209	123	+3 V	slight var.? (AAVSO(VSX): 2.01–2.10 in V) ¶ Astron. Alm. (epoch 2021.5) assigns MK type M0^{+} IIIa	Mirach
δ	Cas A	1 27.3	+60 21	2.66	0.16	A5 IV	32.8	0.2	99	0.301	99	+7 SB†	slight var. 2.68–2.76 in V (as ecl.) now discounted? δ Cas A, as an (unresolved) SB, has been asserted to be eclipsing ¶ E(B–V) =+0.27	Ruchbah
γ	Phe	1 29.3	−43 12	3.41v	1.54	K7 IIIa†	14	−0.9	230	0.209	185	+26 SB	SB period 193.85 d; also var.: 3.39–3.49 in V, 194.1 d ¶ Astron. Alm. (epoch 2021.5) assigns MK type M0^{-} IIIa	
α	Eri	1 38.5	−57 07	0.45†	−0.16	B3 Vnpe (shell?)†	23	−2.7	140	0.095	114	+16 V	ecl. slight var.: 0.40–0.46 in Hp, 1.263 d slight variable in λ Eri class (pulsation? or, rather, starspots?) ¶ brightest of the "Be phenomenon" stars (but first recognized as such only recently, ~1976); an active Be phase that began between 2012 Dec. and 2013 early Jan. (after a period of inactivity, with α Eri presumably diskless, over the previous 7 years) was first noted in amateur spectroscopy, in Brazil; Balmer hydrogen-α line indicates a slow, steady buildup of the Be disk, over a period of ~1.6 y, with polarization suggesting that disk was slightly less dense in 2014 than it had been in 2013; 2017A&A...601A.118D, a case study of α Eri, presents for the first time in astrophysics images of a disk forming around a Be-phenomenon star (with H-band (IR) emission from the disk extending to an outer radius of between 1.7 and 2.3 stellar equatorial radii, in good agreement with current computations in the general theory of the Be phenomenon); it is possible that plane of α Eri disk is inclined to plane of stellar equator; rapid variations in polarization indicate that in addition to its disk, α Eri possesses rings, due to episodic ejections of gas consignments from its photosphere ¶ α Eri is a notably rapid rotator (< 2.1 d) within the (currently small) population of stars interferometrically resolved; period in or near the disk phases varies, either because gas is injected ("decreted") from photosphere into Be disk or because Be-disk gas is re-accreted onto photosphere; interferometry as performed when α Eri is temporarily without its Be disk reveals oblateness (cf e.g. www.eso.org/public/unitedkingdom/news/eso0316/) ¶ although 2008A&A...484L..13K reports companion, at Dec. 2007 angular separation < 0.15″, WDS has not, as of 2020 Nov., asserted binarity; the orbital motion of this companion seems to not be correlated with the repeated formation and disappearance of the Be disk	**Achernar**
τ	Cet A+6P	1 45.1	−15 49	3.49	0.73	G8 V†	~274.0	5.7	11.9	1.921†	296	−16† V		

													mass < 1 $M_\odot$ (unusual in Sample S, although typical in Population P) ¶ high space velocity, low metallicity: interloper from thick galactic disk ¶ on original Frank Drake (1960) SETI target list	
α	Tri A	1 54.4	+29 41	3.42	0.49	F6 IV	52	2.0	63	0.234	177	−13 SB	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	Mothallah
β	Ari A†	1 55.9	+20 55	2.64	0.16	A4 or A5 Va†	56	1.4	59	0.148	138	−2 SB2†	β Ari B (mag. 5.2 in IR or Johnson R band or similar) is SB companion of β Ari A, and yet AB has also been resolved interferometrically as an exceedingly tight binary (15 measurements, 1988→2005, with two quite similar published orbital solutions (*e* is high in both solutions, at ~0.9; A-to-B distance is 0.08 AU min, 1.2 AU max; period is 107 d)) ¶ 2003AJ....126.2048G discusses MK type	Sheratan
ε	Cas	1 56.0	+63 47	3.35	−0.15	B3 IV:pe (shell?)†	8	−2.2	400	0.037	121	−8 V	instance of "Be phenomenon" ¶ He-weak (cp α And, α Tel)	Segin
α	Hyi	1 59.5	−61 28	2.86	0.29	F0n III–IV†	45	1.1	72	0.265	84	+1 V	rapid rotator (< 30 h) ¶ metal-rich	
γ	And A	2 05.3	+42 26	2.10	1.37	K3 IIb†	9	−3.1	400	~0.065	~139	−12 SB	B: 5.0, B9 V, 9.6″ (2019); C: 6.5, A0 V; BC 0.2″ BC orbit 63.7 y ¶ limb darkening observed interferometrically (disk 6.80 mas) ¶ Astron. Alm. (epoch 2021.5) assigns MK type K3$^-$ IIb	Almach
α	Ari +1P	2 08.4	+23 34	2.01†	1.15	K2 IIIab	~49.6	0.5	66	0.240	128	−14 SB	slight slow irregular variability, 2.00−2.03 in V ¶ calcium weak	**Hamal**
β	Tri	2 10.9	+35 06	3.00	0.14	A5 IV	26	0.1	130	0.154	105	+10 SB2†	period of this unresolved SB is 31.39 d, with orbit (at least two solutions published) rather elongated (*e*=0.4 or 0.5; inter-component distance possibly 0.17 AU min, 0.42 AU max) ¶ IR excess (circumstellar matter? possible harbinger of planetesimals)	
o	Cet Aa†	2 20.5	−2 53	6.47v†	0.97	M5–10 IIIe†	11†	1.7	300†	0.238	178	+64 V	LPV, 2−10.1; next max expected 2022 mid-July recent o Cet Aa maxima Nov. 2019 (V~2.3), Sept.–Oct. 2020 (V~3.0), Aug. 2021 (V~2.6), AAVSO reports V band mag. 8.06 on 2022 Jan. 31: 2009ApJ...691.1470T discusses variability, including variation in dominant (333 d) pulsation period and the question of longer-period variations ¶ times of maxima are, and times of minima are not, independent of wavelength: minima are at least coarsely correlated with maximum diameter of o Cet Aa ¶ prototype of the AGB variables, mass ~1 $M_\odot$: the first O-rich AGB star with a CI detection (2018A&A...612L..11S) ¶ physical radius ~2 AU in visual, ~4 AU in IR, still greater upon taking instead the "radio photosphere," which itself increases in radius as progressively longer radio wavelengths are selected: 2015ApJ...808...36M draws parallels with α Ori Aa, attributing radio inhomogeneities in both cases to convective cells (and cf also 2016A&A...592A..42K, which summarizes some recent radio work) ¶ Ab (VZ Cet) WD, 10.4, 0.5″, orbit ~500–600 y; what is seen as o Cet Ab (with professional equipment), and indeed probably seen as itself variable (over and above the long-period variability o Cet Aa) may be not the WD itself, but an accretion disk around the WD, captured from the o Cet Aa wind ¶ nearest instance of (weak) symbiotic binarity, and the only symbiotic to be observed in all wavelength regimes from X-ray to (mm, also cm) radio; interferometry (in IR) is available from VLT, and *CASSINI* has yielded (via Saturn-ring occultations) tomographically	Mira†

													recovered imaging (2016MNRAS.457.1410S); *GALEX* has found bow shock, tail (length 13 ly) in ISM: mass-loss rate ~2.5e-7 $M_{\odot}$ /y; asymmetric atmosphere is discussed in 2016MNRAS.457.1410S ¶ 2018A&A...620A..75K reports dust trail linking Aa,Ab (consistently with other reports of Aa-to-Ab mass transfer) ¶ 2016A&A...592A..42K (2017A&A...599A..59K) discusses o Cet Aa dust nucleation generally, with reference to aluminum (resp. titanium) species: in o Cet Aa, it is silicates that dominate the spectrum (in contrast with less-evolved stars, in which alumina features are spectrally dominant); 2016A&A...590A.127W discusses SiO gas, o Cet Aa inner dust shells: it seems still an open question whether o Cet Aa dust formation is cyclic, as part of the photometrically evident pulsation cycle, or proceeds independently of the pulsations ¶ X-ray emission from o Cet Aa was reported in 2005 (2005ApJ...623L.137K, as the first X-ray detection from an AGB star), and OH, SiO maser emission has also been reported (cf, e.g. 2017MNRAS.468.1703E); further, 2015A&A...577L...4V asserts a hot spot, proposing magnetic activity as the cause ¶ 2016A&A...590A.127W summarizes history of modelling: models generally agree that near o Cet are alternating circumstellar layers of infall and outflow, and that at greater radii is an accelerating outflow, from dust-driven winds: recent observations have tended to agree with overall results from running CODEX (e.g. 2014A&A...565A.119S) ¶ 2016A&A...586A..69P discusses discrepancies in distance determinations (350 ly, 380 ly, 340 ly, and (least reliable?) *HIPPARCOS* 300 ly) ¶ Aa,Ab orbit would, if better known, yield improved total mass of Aa,Ab system, thereby advancing the overall theory of AGB stars ¶ protoplanetary disk was detected around Ab in 2007 ¶ Astron. Alm. (epoch 2021.5) assigns MK type M5.5–9e III ¶ Fabricius noted variability in 1956; Hevelius proposed the name Mira in 1642 ¶ for entry-level briefing-with-bibliography, cf www.aavso.org/vsots_mira2,updpdating www.aavso.org/vsots_mira; and for summary of recent primary literature, cf first section of 2016MNRAS.460..673N	
γ	Cet A†	2 44.5	+3 20	3.54	0.09	A2 Va	41	1.5	80	0.207	225	−5 V	B: 6.23, 2.0″, PA 283°→299°, 1825→2015 orbit ≥ 320 y	Kaffaljidhma
θ	Eri A	2 59.1	−40 13	3.28	0.17	A5 IV	30	0.5	100	0.057	293	+12 SB2	B: 4.35, A1 Va, 8.2″, PA 82°→90°, 1835→2020	**Acamar**
α	UMi Aa†	3 00.7	+89 21	1.97†	0.64	F5–8 Ib	7.5†	−3.6	430†	~0.046	~105?	−17 SB	slight Cep. var, 4.0 d; B: 9.1, F3 V, 18.4″ (2016) the brightest of the Cepheids, but not a classical Cepheid, matching instead the "s-Cepheid" light-curve phenomenology of 1995A&A...303..137B ¶ AAVSO(VSX) as viewed 2021 Jan. 15, and again as viewed 2022 March 3, gives V-mag. range 1.97–2.00, period ~3.97 d: period is increasing ~4.4–4.9s/y, with sudden change around 1963, and with CORIOLIS satellite suggesting a further recent change: period change is often in Cepheid theory linked to evolution, but this may not be the whole story here (in particular, pulsation-driven mass loss through stellar wind, as affirmed by some recent authors (denial also published) would increase the rate of period change) ¶ pulsation mode (1st overtone? 2nd? fundamental?), evolutionary history (1st crossing of IS? or 3rd crossing?), and distance are controverted by various 2010-through-2018 authors (we here use *Gaia* DR2 distance for α UMi B as a proxy, assuming with the current literature that B is indeed gravitationally bound with Aa,Ab) ¶ α UMi Aa is first Cepheid with mass determined through purely dynamical means (via the Aa,Ab orbit: Aa is single-lined SB, and Aa, Ab have been resolved with HST, as first announced (0.17″) in 2008AJ....136.1137E	**Polaris**

													(orbit ~30 y)) ¶ 2000A&A...360..399W warns that peculiarity of the α UMi Aa light curve makes this particular Cepheid perhaps not a suitable anchor point for determining the overall Cepheid period-luminosity relation (essential though a determination of that relation is for calibrating the cosmic distance scale) ¶ α UMi Aa is important as a case study in the “Cepheid mass discrepancy” problem (Cepheid masses deduced from pulsation periods are found to be too low in comparison with masses from stellar-evolution modelling) ¶ strictly a three-star system, UMi Aa+ UMi Ab+UMi B; Aa,Ab has period 29.6 y, inter-component distance 6.7 AU min, 27 AU max, 17 AU average (orbital solution has been published); B experiences Aa,Ab as essentially a point mass, with period ≥ 42,000 y, separation at least 2400 AU; B is mag. 9.1, at angular separation 18″ (no orbital solution published) ¶ α UMi Aa, Ab,B is approaching NCP: closest approach will be 14′, in ~2105 ¶ B has E(B–V)=0.0 ¶ 2018ApJ...863..187E summarizes recent work	
α	Cet	3 03.5	+4 11	2.54†	1.63	M2 III†	13	−1.9	250	0.078	188	−26 V	slight var., 2.45–2.54 in V rather mild variable, of the “giant irregular” type ¶ radio source (due to stellar wind) ¶ notably deficient in carbon ¶ Astron. Alm. (epoch 2021.5) assigns MK type M1.5 IIIa	**Menkar**
γ	Per <u>Aa,Ab</u>†	3 06.4	+53 36	2.91v†	0.72	G8 III† + A2? V?	13	−1.5	240	0.006	175	+3 SB2†	composite spectrum, orbit 14.6 y, next eclipse 2035 eclipse duration < 2 weeks; eclipse variation is significantly above threshold of naked-eye detection, with AAVSO(VSX) giving V-mag. range 2.91−3.21 (ranges in Johnson U and B bands are still larger) ¶ angular separation of Aa,Ab is always ≤ 0.3″, putting this binary, for part of its orbit, just within the limits of feasibility for traditional micrometer astrometry; orbit is highly elliptical, with *e*=0.79 on the better of two published orbital solutions; amenability of this binary to spectroscopic orbit solution (via radial velocities: radial-velocity data are available since ~1900) and to interferometry make it a useful candidate for stellar-mass studies, and therefore a useful test for theoretical predictions of luminosity classes and spectral types from given masses (1999A&A...348..127P)) ¶ Aa is mag. 3.6 and Ab is mag. 3.8 ¶ Astron. Alm. (epoch 2021.5) assigns MK type G5 III	
ρ	Per	3 06.6	+38 56	3.32v†	1.53	M4 II	11	−1.6	310	0.167	129	+28	semiregular variable: 3.3–4.0 in V, ~50 d, ~120 d, ~250 d period ~50 d, with possibly also a longer period	
β	Per<u>Aa1</u>†	3 09.6	+41 02	2.09v†	0.00	B8 V †	36	−0.1	90	0.003	119	+4 SB	Aa=compos. spectrum Aa1,2 ecl,;2.09-3.30 in V, 2.9 d Algol Aa2 is K2IV? ¶ in older terminology, β Per Aa1 = β Per A, β Per Aa2 = β Per B, β Per Ab = β Per C: but WDS, following the current terminology, uses the “B” and “C” for other purposes (since there are optical neighbours B,C,D,E,F,G,H; all are between 5″ and 100″ from the Aa1,Aa2,Ab triple, and all are fainter than mag. 10); system is hierarchical, with outlying Ab experiencing the close Aa1,Aa2 pair (Aa1−Aa2 distance is just 14.14 $R_{\odot}$) as essentially a point mass; angular separation between Aa1,Aa2 and Ab is ~0.1″ (WDS 1973, 2010); orbital solutions have been published both for the tight binary which is is β Per Aa1,Aa2 and for the wider binary which is β Per Aa,Ab ¶ among the most visually prominent of the eclipsing binaries, and for theoreticians the most familiar of the semidetached binaries (i.e., binaries in which one of the two Roche equipotential surfaces	

is fully occupied, the other not)
¶ Aa2 is tidally locked, in a rapid circular orbit with Aa1; the consequent rapid spin of Aa2 causes dynamo action in Aa2 convection zone, with Aa2 consequently having complex magnetosphere (mass-transfer stream possibly even deflected out of Aa1,Aa2 orbital plane by magnetics; 2012ApJ...760....8R; Aa2 has additionally a meridional coronal loop, approximately as high as the diameter of Aa2 (the size exceeds what has been anticipated from modelling) believed pointing at all times to Aa1), X-ray emission, varying radio morphology (double-lobed when radio-brightest) and CME episodes (2017ApJ...850..191M suggests the 1997 Aug. 30 superflare event supplies "arguably the best candidate" for a non-solar CME)
¶ the (unsteady) Aa2-to-Aa1 mass transfer, while ongoing, and indeed responsible for an annulus around Aa1, is no longer copious (in contrast with the copious transfer still present in, e.g. β Lyr)
¶ it is not the (now modest) unsteady mass transfer, but possibly instead the Applegate mechanism (1992ApJ...385..621A), implicating a stellar magnetic activity cycle, which dominates the Aa2,Aa1 period variation (increase-decrease-increase cycle, not quite strictly periodic, 32 y: there are additionally period modulations of 1.9 y and 180 y; as viewed 2021 Jan. 28, AAVSO(VSX) asserts period 2.86736 d, but as viewed 2022 Mar. 3 gives updated value 2.86734 d); full amplitude of the Aa1 Aa2 period variation is ~0.8 s; such alternating period changes in binaries are still not, however, well understood
¶ it is the (several My ago rapid and copious) mass transfer that resolves the "Algol paradox" of a lower-mass more evolved (in this case, sub-giant) star in orbit with a higher-mass less evolved (indeed MS) star; masses are well known in this particular case: 2015MNRAS.451.4150K, having disentangled the β Per Aa1, Aa1, Ab spectra, determines their masses within plus-minus 2%, corroborating 2012ApJ...752...20B
¶ β Per Aa2 elemental abundances below corona and flare (investigated in 2015MNRAS.451.4150K) are of special interest, since mass transfer has stripped off Aa2 outer layers, opening the Aa2 interior to spectroscopic inspection
¶ 1983ApJ...273L..85K reports discovery of Chandrasekhar eclipse-induced stellar limb polarization from β Per Aa1, in a wide optical passband
¶ MK type K2 IV is assigned to Aa2 in at least 3 recent papers, whereas the older 1993ApJ...410..808L has the slightly hotter MK type K0 IV; what is essential here is the agreed "IV" (as opposed to "V"), indicating evolution of this (secondary) star off the MS (and Astron. Alm. (epoch 2021.5) assigns MK type B8 V for Aa1, as we do here, and for one companion (is this Aa2, or is it Ab?) the uncertainty-flagged "F:," without luminosity class)
¶ β Per Ab, spectrally

													Am with some F1V characteristics, orbits the β Per Aa1,Aa2 binary with period ~680 d, without eclipsing ¶ 2012ApJ...752...20B presents CHARA interferometry (~0.5 mas, H (near-IR) band) of the β Per Aa1,Aa2,Ab system (finding orbital plane of Ab to be nearly perpendicular to Aa1,Aa2 orbital plane), and also summarizes earlier interferometry; an approx 55-frame animation from this paper can be conveniently viewed at en.wikipedia.org/wiki/Algol ¶ E(B-V)=+0.03 ¶ 2013ApJ...773....1J suggests that β Per system variability is documented in the "Cairo Calendar" papyrus (New Kingdom, dated to 1271–1163 BCE); al-Sufi (Persia, ca 964 CE) is, however, silent on question of β Per variability ¶ AAVSO has briefing notes, with some history, at www.aavso.org/vsots_betaper; 1910ApJ....32..185S is the discovery paper for β Per Aa1,Aa2 secondary minimum, from the beginnings of photoelectric-cell photometry; 1998A&AT...15..357P analyzes "Algol paradox" history; arxiv.org/pdf/astro-ph/0611855.pdf, "Appendix B.", is a tabular history of β Per-pertinent investigations from antiquity to 1999; in this same K. Wecht 2006 Lehigh Univ PhD thesis, Table 2.5.1 summarizes 1966-through-1983 observational coverage, as tabulated in the less Web-accessible 1986 work of Budding	
α	Per A	3 25.9	+49 56	1.79†	0.48	F5 Ib	~6.4	−4.2	510	0.035	138	−2 V	in open cluster Melotte 20, a.k.a. Collinder 39 near edge of HR diagram IS (slightly too hot to be a straightforward Cepheid)	**Mirfak**
δ	Eri	3 44.3	−9 41	3.52v	0.92	K0 IV	111	3.7	29.5	0.749	353	−6	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	
δ	Per Aa	3 44.5	+47 51	3.01†	−0.12	B5 IIIn†	6	−3.0	500	0.050	149	+4 SB	slight var., γ Cas type, 2.99–3.04 in V (but WDS indicates possible variability of the type presented by α CVn A (a star discussed later in this table); it is variability of the γ Cas type that is asserted by AAVSO(VSX)) ¶ cluster affiliation is controverted ¶ Astron. Alm. (epoch 2021.5) assigns MK type B5 III ¶ E(B–V) =+0.04	
γ	Hyi	3 46.9	−74 10	3.26	1.59	M2 III	15.2	−0.8	~214	0.126	24	+16	evolutionary status is uncertain	
η	Tau Aa	3 48.8	+24 10	2.85	−0.09	B7 IIIne†	8	−2.6	400	0.048	156	+10 V? SB	brightest member of Pleiades ¶ η Tau Aa,Ab is known as a binary both from spectroscopy and from occultation (but this tight pair seems not to have been resolved interferometrically, at any rate as of 2021 Oct. 25; it is therefore on the strength of occultation alone that WDS is able to write "η Tau Aa", "η Tau Ab"), with Ab of mag. 4.6; additionally, in this crowded part of the celestial sphere, WDS catalogues as neighbours of the tight Aa,Ab binary B, C, Da, Db, E, F, G, H; of these neighbours to Aa, not only Ab, but also B, C, Da, and Db are brighter than mag. 10, and possess astrometric data from as early as the 19th century (except that the tight Da,Db pair was first resolved in the 21st century, through speckle interferometry); 1972JBAA...82..431K describes the 18.6-year 1940-through-2050 cycle of lunar occultation possibilities ¶ rapid rotator, with "Be" and "shell-spectrum"	Alcyone

													phenomena (BS5: “rotationally unstable”), making this star an appropriate target for periodic low-cadence (e.g. once-yearly) amateur- spectroscopy monitoring ¶ significant dimming by ISM dust; E(B–V)=+0.03	
ζ	Per A†	3 55.6	+31 57	2.84	0.27	B1 Ib	4	−4.0	800	0.011	150	+20 SB	B: 9.16, B8 V, 12.8″, PA 205°→208°, 1824→2020 orbit ≥ 50,000 y ¶ since the SB which is ζ Per A is not as yet resolved, even in interferometry, WDS is not as yet able (at any rate as of 2021 Oct. 25) to write “ζ Per Aa”, “ζ Per Ab”; ζ Per B experiences ζ Per Aa,Ab as essentially a point mass, and is too slow in its orbit to yield spectroscopic (radial-velocity) data; further, WDS gives, as celestial-sphere neighbours of the wide and slow ζ Per Aa,Ab−B pairing, ζ Per C, D, and E, with only E (marginally) brighter than mag. 10, at mag. 9.96 (and with the A,E angular separation, as measured in 2015, wide, at 120″) ¶ significant dimming by ISM dust; E(B–V) =+0.33 (pronounced reddening)	
γ	Eri A	3 59.1	−13 27	2.97	1.59	M1 IIIb†	16	−1.0	200	0.129	151	+62	Ca, Cr weak ¶ Kaler, at stars.astro.illinois.edu, writes, “must be one of the least- studied of the cooler bright stars” ¶ Astron. Alm. (epoch 2021.5) assigns MK type M0.5 IIIb Ca–1	Zaurak
ε	Per A†	3 59.4	+40 04	2.90†	−0.20	B0.5 IV†	5	−3.6	600	0.028	149	+1 SB2	B: 8.9, B9.5 V, 9.1″, PA 10°→10°, 1821→2015 orbit ≥ 16,000 y; since the SB2 has not yet been resolved, even in interferometry, WDS is not as yet able (at any rate as of 2021 Oct. 25) to write “ε Per Aa”, “ε Per Ab”; ε Per B experiences the unresolved SB2 which is ε Per A as essentially a point mass, and is too slow in its orbit to yield spectroscopy (radial- velocity) orbital data; further, WDS gives, as celestial-sphere neighbours of the wide and slow ε Per AB pairing, ε Per C and D, with D a little brighter than mag. 10 (at mat 9.25), and at the wide angular separation of 163″ from ε Per A; the unresolved SB2 which is ε Per A does have published orbital solutions, with period 14.069 d (indicating a tight pairing), and with *e*=0.5 ¶ slight variable, possibly of β Cep type (2.89–2.91 in V; one of the most extreme spectroscopic variables (periods 2.27 h and 8.46 h)) ¶ E(B–V) =+0.10	
λ	Tau	4 01.9	+12 33	3.41v†	−0.10	B3 V	7	−2.4	480	0.017	209	+18 SB2†	ecl.: 3.37–3.91 in V, 3.953 d; secondary is A4 IV AAVSO(VSX) as viewed 2021 Jan. 16, and again at 2022 March 3, gives period 3.9529478 d ¶ shape distortion (mutual tides), reflection effect, some evidence of mass transfer	
α	Ret A	4 14.7	−62 25	3.33	0.92	G8 II–III	20.2	−0.1	162	0.065	40	+36 SB?		
ε	Tau Aa +1P†	4 29.9	+19 14	3.53	1.01	K0 III†	22.2	0.3	150	0.113	110	+39 V?	in Hyades; Aa,Ab 0.2″ (2005), mags. ~3.6, ~6.0 as of 2021 Oct. 26, WDS records just one (interferometric) measurement of the ε Tau Aa,Ab binary ¶ metal-rich ¶ first known instance of a planet-host in an open cluster; unusually massive among the currently known planet-hosts ¶ Astron. Alm. (epoch 2021.5) assigns MK type G9.5 III CN0.5	Ain
													[THIS STAR ONLY IN ONLINE VERSION OF TABLE]	
θ	Tau Aa†	4 30.0	+15 55	3.40†	0.18	A7 III	22	0.1	150	0.112	104	+40 SB†	system Aa-plus-Ab is a.k.a. θ^2 Tau; in Hyades companion in elongated orbit, with published orbital solutions (period is 140.728 d; *e*=0.73; 0.23 AU min, 1.3 AU max); we here state mag. for Aa,Ab combined light (separately these are mag. ~3.7, mag. ~4.9, making Aa alone	Chamukuy

													a little fainter than what this table consider a “bright star”, even though the naked-eye bright point which is θ Tau Aa,Ab is the brightest Hyades member; one of the components in the Aa,Ab pair, typically presumed to be Aa, is a δ Sct variable, with V range somewhat less than 0.1 mag., for which 12 periods are known, 1.64 h to 2.22 h, the ranges in some cases small (0.5 millimag., 30 millimag.); this is one of the intensely studied cases of δ Sct variability); the SB system θ Tau Aa,Ab forms a wide pair with the bright SB system θ Tau Ba,Bb, a system a.k.a. θ^1 Tau (mag. ~3.9; astrometry of Aa,Ab with respect to Ba,Bb is 340″→348″, PA 346°→347°, 1800→2016); not only the Aa,Ab but also the Ba,Bb pair has a published orbital solution; Aa,Ab and Ba,Bb are in turn gravitationally bound, with each of the two tight pairs in this quadruple of stars experiencing the other tight pair as essentially a point mass; the entire Aa,Ab,Ba,Bb system, notably including at least three of the four individual masses, has been much studied since the 1990s, drawing on data from occultation, spectroscopy, and interferometry, even though the various challenges include some troublesome rotational broadening of spectral lines, since both Aa and Ab are rapid rotators; determination of masses, plus (helpfully, even *HIPPARCOS*-independent, via orbit model) determination of distances, confers on this Aa,Ab,Ba,Bb system, as on several other stars in the Hyades, a specially enhanced stellar-astrophysics significance, as isochrone-anchored data points for plotting the empirical (i.e. the theory-free) mass-luminosity relation, and therefore for constraining models of stellar evolution (1997ApJ...485..167T recapitulates the strategic position as follows: “The Hyades is unique in this respect. In no other case have dynamical masses been determined over a range covering much of the main sequence /…/ for stars of the same age and known chemical composition.”) ¶ the θ Tau Aa name “Chamukuy”, IAU-official since 2017, is the Yucatec Mayan name for a small bird	
α	Dor A†	4 34.5	−55 00	3.30	−0.08	A0p V: (Si)†	19	−0.3	169	~0.059	~79?	+26	A: 3.8; B: 4.3, B9 IV; 0.2″ (2019); orbit 12 y orbit very elongated, with A-to-B distance 1.9 AU min, 17.5 AU max ¶ Astron. Alm. (epoch 2021.5) assigns MK temperature type A0p Si (and does not assign an MK luminosity class)	
α	Tau A +1P†	4 37.2	+16 33	0.87	1.54	K5 III†	49	−0.7	67	0.199	161	+54 SB	slight irregular var., 0.86–0.89 in V ¶ foreground star, not true Hyades member; among the nearest of the red giants; evolution has proceeded beyond the “FDU” stage which accompanies helium-core contraction on RGB ¶ 49 lunar occultations occurred over the period 2015 Jan. 29/2018 Sep. 03 (and yet there is a surprisingly large scatter in the occultation determinations of α Tau angular diameter; 1972JBAA...82..431K describes the overall 18.6-year 1940-through-2050 cycle of lunar occultation possibilities) ¶ in contrast with its celestial-sphere neighbour α Ori, α Tau is of modest mass (with recent literature variously offering ~1.2 $M_{\odot}$, ~1.3 $M_{\odot}$, ~1.5 $M_{\odot}$): Appendix C of 2018ApJ...865L..20F tabulates values for mass, luminosity, radius, age, and several other parameters, on the strength of five separate 2008-through-2012 spectroscopy investigations ¶ 2013A&A...553A...3O reports “MOLsphere” (molecule-harbouring atmosphere) inhomogeneities, from VLTI/AMBER, thereby helping advance the still poorly	Aldebaran

													understood topic of RGB mass loss (especially in a context in which dust condensation might appear not to play a significant role; in general, it is RGB mass loss that is puzzling, AGB mass loss that is straightforward) ¶ recent literature proposes oscillations, and also proposes rotational modulation from modest photospheric-activity features (with possibly an activity cycle (2015A&A...580A..31H): the features could be (cool) starspots, but could alternatively be large convection cells; the general topic of activity in K giants is not yet well understood) ¶ Astron. Alm. (epoch 2021.5) assigns MK type $K5^{+}$ III ¶ although 2019A&A...625A..22R casts doubt on 2018ApJ...865L..20F, 2015A&A...580A..31H exoplanet assertion, exoplanet is asserted in NASA exoplanet catalogue (as viewed 2021 Aug. 07)	
π^3	Ori A	4 51.1	+7 00	3.19	0.48	F6 V	124	3.7	26.3	0.464	89	+24 SB2		Tabit
ι	Aur	4 58.5	+33 12	2.69v	1.49	K3 II	7	−3.2	500	0.016	155	+18V	var.: 2.63–2.78 in V; maybe "+2P" (brown dwrfs?) ¶ X-ray "hybrid star" (unusual combination of (hot) corona, cool wind) ¶ dimming by ISM dust, ~0.6 mag.	Hassaleh
ε	Aur A	5 03.6	+43 51	3.03†	0.54	F0Iab?†	~2†	−8.0:	~1450†	~0.003	n.a.	−3 SB†	slight ecl.: 2.92–3.83 in V, ~27.1 y (dim ~700 d) more formally, period is 9896.0 ± 1.6 d: as again discussed twice below, there are spectroscopic, as distinct from photometric, phenomena indicative of an eclipsing mass before the onset, and continuing after the end, of the photometric eclipse ¶ Astron. Alm. (epoch 2021.5) assigns MK type A9 Ia ¶ ε Aur B MK type is ~B5 V ¶ ε Aur ranks among the longest-period eclipsing binaries (exceeding even V383 Sco, with period 13.5 y; WW Vul, with period 13.9 y; and VV Cep, with period 20.3 y: the current long-period record, however, is held by TYC 2505-672-1, at ~69.1 y (with dimming ~3.45 y)) ¶ determination of orbit elements has proven troublesome, with 2012A&A...544A..91M urging caution even in respect of recent careful studies ¶ it is remarkable that, even though the eclipsing entity is physically very extended (because the eclipse is protracted), and even though orbital dynamics indicates that the entity is quite massive, nevertheless no visible radiation from an eclipsing body is readily observable (i.e. it is remarkable that this SB is essentially a single-lined SB) ¶ although the (notably protracted) ε Aur eclipse is largely flat-bottomed, nevertheless even during eclipse the (dimmed) spectrum of the primary can be seen, with no visible-wavelength colour preference in the attenuation (except that there are absorption lines, as from a semi-transparent atmosphere around the eclipsing mass, at the start and the end of the dimming); the 1937ApJ....86..570K explanation, postulating a large semitransparent totally eclipsing mass, with the non-selective opacity due to scattering off free electrons, is now universally abandoned in favour of the 1953AJ.....58..219K and 1965ApJ...141..976H hypothesis of an almost edge-on (2010ApJ...714..549H) cool opaque gas-dust low-mass disk or disk-like entity (spiral arm? cf. 2013PASP..125..775G) (rotating while orbiting, and several AU in diameter, presenting a	Almaaz

temperature gradient ~550 K to ~1150 K (representing
the portions respectively
farthest from and closest to the primary star),
and in terms of its vertical
development not a (thick) hockey puck
but a (thin) wafer, of much larger
radius than the primary star;
2013PASJ...65L...1S gives
evidence for clumping in the disk;
2015ApJS..220...14K raises the
possibility that the disk is slightly tilted out of the
binary-system orbital plane, with consequent precession),
shrouding a B-type
star (B5V?) or star pair (the more dramatic
hypothesis of a shrouded black hole is not now
generally favoured: 2010AJ....140..595W, e.g. reports
null result from X-ray search), with the disk geometry
making the eclipses of the primary star,
as observed from Earth, only partial;
the disk may have been formed
by mass transfer from the primary star,
and indeed 2013PASP..125..775G
and 2018MNRAS.479.2161G
report putative spectroscopic
detection of narrow mass-transfer stream;
the former paper stresses that the
detection of rare-earth elements
within the putative stream
spectrum (an indication that the primary
is highly evolved?) now poses a fresh puzzle,
in a system traditionally classed as puzzling
¶ 2012ApJ...748L..28H and
2012MNRAS.423.2075M discuss the question of gas-to-dust
ratio in the disk; 2015ApJ...798...11P,
2012MNRAS.423.2075M, and 2010ApJ...714..549H
suggest not-very-small values
in the distribution of dust-particle
diameters, with the first
two of these three papers
suggesting carbonaceous chemistry;
additionally, 2011AJ....142..174S
spectroscopy finds CO absorption bands,
symptomatic of sublimation, with indications that
large particles dominate
¶ 2013ARep...57..991P and 2013PASP..125..775G
document indications
that the structure of the disk does not greatly
change from one eclipse to the next
¶ the brightening around mid-eclipse has
in the post-1970 papers repeatedly been
attributed to a central opening in the postulated disk:
however, (*a*) dissenter 2011A&A...530A.146C
has instead suggested intrinsic variability in the primary
(which indeed has various quasi-periods or periods,
with 67 d and 123 d prominent,
with also variations
in radial velocity, and (unblended) spectral line width,
and other periodic or quasi-periodic behaviour, including
possible orbitally excited
non-radial pulsation; there seems as
yet, however, to be no extensive asteroseismology),
and (*b*) dissenter 2011A&A...532L..12B
has instead suggested
forward scattering by disk dust (a line of thought now
supported by the key imaging-and-modelling paper
2015ApJS..220...14K)
¶ *HIPPARCOS* yields π possibly < 2 mas, distance ~2000 ly;
we now, however, choose to relinquish
the *HIPPARCOS* determination,
made at the limit of *HIPPARCOS* capabilities,
in favour of 2019IBVS.6258....1P,
which deduces from *Gaia* DR2

π =2.4144 ± 0.5119 mas, and goes
on to deduce from this, via supplementary (not straightforward , Bailer-Jones et al. 2018AJ....156...58B) considerations what we express here as "~1450 ly"
¶ section 1 of 2012A&A...546A.123G and section 1 of 2012A&A...544A..91M summarize past controversies regarding mass of primary (low or high?), stemming from the difficulty in determining distance (2012A&A...546A.123G assigns a high distance, ~4900 ly, and consequently favours a high mass value, ~20 $M_{\odot}$; however, several post-2009 papers instead assign a modest mass to the primary, suggesting various values within the range ~2 $M_{\odot}$ – ~6 $M_{\odot}$: 2014MNRAS.445.2884M, e.g. suggests 2.5 $M_{\odot}$ for primary, 5.4 $M_{\odot}$ for secondary (and suggests disk diameter 8.9 AU)): evolutionary status of the primary has been correspondingly controverted (post-AGB star, now of modest mass, with much past shedding of mass, and consequent accumulation of the low-mass opaque disk around the secondary (a view taken by various papers, including recently 2019IBVS.6258....1P) or, rather, an evolutionally earlier supergiant (cf 2012JAVSO..40..647K), even perhaps of high mass? – but it is clear that the primary is at any rate sufficiently evolved to have left the MS, and there are indications that it is pulsating and a wind source; angular diameter is 2.1 mas)
¶ most recent photometric eclipse started 2009 Aug. 12, ended 2011 Aug. 23 ±15 d;
next secondary (shallow, for the casual observer elusive) eclipse is possibly 2025 Dec. 20 through 2028 Mar. 29;
next (deep, easy observable) primary photometric eclipse starts in 2036;
monitoring even outside both the primary eclipse and the secondary eclipse is useful, in part because of intrinsic variations in the primary star (cf 2012JAVSO..40..647K); in part because the postulated dense disk has an extended "atmosphere" yielding (e.g.) Hα absorption even outside photometric eclipse (2011A&A...530A.146C), with spectral premonitions starting ~3 y before the onset of the photometric eclipse; and in part because the opaque primary-star-eclipsing disk is potentially liable to thermal changes, visible in mid-infrared outside primary and secondary eclipse (2011AJ....142..174S)
¶ the Kloppenborg et al. CHARA interferometric imaging of the eclipsing disk is perhaps the single largest 21st-century advance in ε Aur studies: 2010Natur.464..842G supplies journalistic background, including a recapitulation of 2010ApJ...714..549H modelling; 2010Natur.464..870K is the formal Kloppenborg et al. discovery paper (with the first spatially resolved image for any eclipsing binary during eclipse); and 2015ApJS..220...14K is a Kloppenborg-et-al update, with additional interferometry, now including also PTI and NPOI (and supplying also an overall history of ε Aur studies)
¶ news sources include mysite.du.edu/~rstencel/epsaur.htm

													(Prof. R. Stencel, Univ of Denver, on the Kloppenborg-2010 team) and twitter.com/epsilon_Aurigae; 2012JAVSO..40..618S summarizes the 2009-2011 campaign from an AAVSO perspective; an 18-paper archive, of NSF-supported ~2009-through~2011 AAVSO eclipse campaign, is at www.aavso.org/citizen-sky-epsilon-aurigae-papers ¶ since the eclipsing companion of ε Aur A has not yet been resolved, even in interferometry, WDS is not as yet (at any rate as of 2021 Oct. 27) able to write "ε Aur Aa", "ε Aur Ab"; WDS does, on the other hand, catalogue celestial-sphere neighbours B,C,D,E,F,G,H,I,J,K; all of these are fainter than mag. 10, except for ε Aur E (at the wide angular separation of 207″ from ε Aur A, and at mag. 9.6 just barely clearing our mag.-10 threshold for comment-worthiness of a bright-star neighbour on celestial sphere)	
ε	Lep	5 06.4	−22 21	3.19	1.46	K4 III	15	−0.9	210	0.076	164	+1	evolutionary status is uncertain: RGB or AGB?	
η	Aur	5 08.1	+41 16	3.18	−0.15	B3 V†	13	−1.2	240	0.075	155	+7 V?	slight rotating ellips. var?: 3.16–3.18 in V, 2.5617 d spectral variations also suggested ¶ weak magnetic field detected, ~2× strength of Earth's dipole field	Haedus
β	Eri A	5 09.0	−5 04	2.78†	0.16	A3 IVn	36	0.6	89	0.112	228	−9	unexplained brightening episode, over 2 h, by ~3 mag, in 1985 (recalling the 1972 unexplained brightening of ε Peg)	Cursa
μ	Lep	5 13.9	−16 11	3.29v?†	−0.11	B9p IV: (HgMn) †	18	−0.5	190	0.050	109	+28	var?.: 2.97–3.41 in V?, 2 d? variable of α^2 CVn type? (variability so far unconfirmed, and no CVn-class-appropriate magnetic field detected yet?) ¶ among the brightest of the Hg-Mn stars ¶ Astron. Alm. (epoch 2021.5) assigns MK temperature type B9p HgMn (and does not assign an MK luminosity class) ¶ X-ray emission noted from putative companion, at angular distance 0.93″	
β	Ori A†	5 15.6	−8 11	0.18†	−0.03	B8 Ia	4	−6.9	900	0.001	69	+21 SB	B: 6.8, B5 V, 9.7″ (2017); C: 7.6; BC: 0.1″ A−BC orbit ≥ 25,000 y, BC orbit ~400 y ¶ variable in the α Cyg class (non-radial pulsator): 0.17–0.22 in Hp ¶ E(B–V) =+0.00	Rigel
α	Aur Aa,Ab†	5 18.4	+46 01	0.08	0.80	G6:III + G2:III	76	−0.5	43	0.433	170	+30 SB2	composite; Aa: 0.7, Ab: 0.9; Aa,Ab < 0.1″ Aa,Ab are resp. mags. 0.08, 0.18 ¶ under IAU rules, "Capella" designates Aa, not Ab ¶ orbit of the Aa, Ab SB is 140.0 d, with orbit exactly or very nearly circular; Aa,Ab is the first binary with orbit studied interferometrically (Anderson-Pease, Mt Wilson, 1910: the binary is informally known as "The Interferometrist's Friend"); full system, however, appears to be α Aur Aa+Ab+H+L, where H and L are red dwarfs, of respective mags. 9.99 and 13.5, sharing the proper motion of Aa+Ab and perhaps possessing further gravitationally bound companions (with α Aur B,C,D,E,F,G,I,J,K being mere line-of-sight coincidences; of these 9 celestial-sphere neighbours to "The Interferometrist's Friend" which is the Aa,Ab SB, only G, at wide angular separation of 522″ from Aa,Ab, is brighter than mag. 10 (at mag. 8.10); additionally, WDS documents the celestial-sphere neighbours M,N,O,P, of which only the sparsely observed M and N are brighter than mag. 10 (at mag. 6.29 and mag. 9.84, respectively); the crowding, surely with abundant line-of-sight coincidences, in this region of the celestial celestial sphere is perhaps to be expected, given the celestial-sphere adjacency of	**Capella**

													α Aur to the Milky Way ¶ α Aur Ab is in rapid evolutionary transition, currently crossing the Hertzsprung Gap ¶ system is among the brightest of X-ray sources	
η	Ori Aa†	5 25.6	−2 23	3.35v†	−0.24	B0.5 Ve†	3	−4.0	1000	~0.004?	n.a.	+20 SB2	ecl.: 3.31–3.60 in V, 8.0 d; A: 3.6; B: 4.9, 1.8″ (2019) η Ori Aa is unresolved SB2, strictly 7.9904 d, eclipsing, to be informally thought of as “the Aa1, Aa2 binary” (but since the SB2 is not as yet resolved, even in interferometry, WDS is not as yet able to formally catalogue the stars as η Ori Aa1, η Ori Aa2, and uses the informal terminology only in a note); Ab experiences “Aa1”, “Aa2” as essentially a point mass; the Ab-with-“Aa1,Aa2” period is 9.44 y, spectroscopically SB but not SB2 (with at least one orbital solution published); in a very slow orbit with the three-star “Aa1,Aa2 binary”-plus-Ab system, and so in turn experiencing that triple as essentially a point mass, is η Ori B (mag. 4.9; AB astrometry is 0.9″→1.9″, PA 87°→77°, 1848→2020, with orbit ≥ 2000 y) ¶ one or other of “Aa1”, “Aa2” is a pulsator, with period 8 h (possible β Cep variability? but one or other of the Aa components is an instance of the “Be phenomenon”, and additionally a shell spectrum has been observed from one or other of “Aa1”, “Aa2”) ¶ Astron. Alm. (epoch 2021.5) assigns MK type B1 IV ¶ BSC5: “expanding circumstellar shell”	
γ	Ori A	5 26.3	+6 22	1.64	−0.22	B2 III†	13	−2.8	250	0.015	212	+18 SB?	BSC5: “expanding circumstellar shell”	**Bellatrix**
β	Tau	5 27.7	+28 37	1.65	−0.13	B7 III†	24	−1.4	130	0.175	173	+9 V SB	BSC5: “expanding circumstellar shell” ¶ lunar occultations possible as far N as southern California: we use simply the designation “β Tau”, even though WDS has “β Tau Aa”; there is a single observation of an alleged “β Tau B”, of undocumented mag., at angular separation 33.40″, from 1898, and additionally a single observation of an alleged “β Tau Ab”, of mag. 2.40, at angular separation just 0.10″, possibly an occultation observation, from 1930; this sparse observational material suggests to us that a lunar-occultation campaign (for observers at suitable latitudes, such as in southern California, although not in Canada or in northern Europe) might be worth undertaking, perhaps to yield a null result rebutting the 1898 and 1930 suggestions of multiplicity, perhaps to yield a non-null result confirming at least one of these two suggestions of multiplicity ¶ often, but not invariably, classified as Hg–Mn star: Mn 25× solar (and Ca, Mg only ~0.12× solar: radiative lofting, gravitational settling) ¶ E(B–V)=0.00	**Elnath**
β	Lep A†	5 29.2	−20 45	2.81	0.81	G5 II	~20.3	−0.6	160	0.086	183	−14 V?	B: 7.5, 2.7″, PA 268°→10°, 1875→2017 ¶ β Lep B is possibly variable ¶ duplicity now suspected also in β Lep A, through 2002 adaptive-optics observation at Haleakala: separation 2.58″	Nihal
δ	Ori Aa†	5 33.2	−0 17	2.25v†	−0.18	O9.5 II	5	−4.4	700	0.001	137	+16 SB	a good marker of celestial equator ¶ Aa is an Algol-type eclipsing binary, not as yet resolved (so WDS is not as yet, at any rate as of 2021 Oct. 30, able to write “δ Ori Aa1”, “δ Ori Aa2”) with period 5.73249 d, yielding a combined-light variation reported at AAVSO(VSX) (for the entire Aa,Ab system? or just for the unresolved Aa SB?) over the mag. range 2.20−2.32 (in the V band); WDS assigns to the unresolved Aa SB the MK type O9.5 II, and to the single Ab	Mintaka

													star (mag. 3.76) the MK type (with luminosity class tentative) “O9.7 III:”; angular separation of the unresolved Aa pair and the single star Ab has increased over recent decades (0.2″→0.3″, 1978→2019); at least one orbital solution has been published for Aa,Ab; there is also a bright celestial-sphere neighbour, δ Ori C, at mag. 6.8 (AC astrometry 50″→56″, PA 0°→4°, 1777→2017) ¶ yielded first detection of ISM (Hartmann, 1904, through non-moving Ca line in the SB) ¶ E(B–V)=+0.07	
α	Lep A	5 33.7	−17 48	2.58	0.21	F0 Ib†	1.5	−6.6	2000	0.004	72	+24	evolutionary status unclear (has helium fusion already started in core?); helium-fusion past yields now abundances N 5× solar, Na 2× solar	Arneb
β	Dor	5 33.8	−62 29	3.76v†	0.64	F7–G2 Ib	3.2	−3.7	1000	0.013	4	+7 V	Cepheid variable: 3.41(?)–4.08 in V, 9.84 d period not quite constant; evolutionary status uncertain ¶ observed by FUSE, XMM-*Newton* missions	
λ	Ori A†	5 36.4	+9 57	3.39	−0.16	O8 IIIf	3	−4.2	~1100	0.004	216	+34	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]** B: 5.45, B0 V, 4.3″, PA 45°→44°, 1779→2019 WDS, citing 1985A&AS...60..183L, remarks that B may be a mere line-of-sight coincidence; WDS also gives, in addition to two faint celestial-sphere neighbours, neighbours λ Ori D (mag. 9.6; AD angular separation was 78″ in 2012) and λ Ori F (mag. 9.2; AF angular separation was 151″ in 2012) ¶ the most prominent member of Collinder 69 ¶ within gas ring 150 ly in diameter (possibly, but not certainly, remnant from a Type II supernova) ¶ E(B–V)=+0.12	Meissa
ι	Ori Aa†	5 36.5	−5 54	2.75	−0.21	O9 III	~1.4	−6.5	2000	0.001	108	+22 SB2	Aa,Ab 0.1″ (2016), mags. 3.0, 6.3 B: 7.3, B7 IIIp (He wk), 12.5″, PA 134°→146°, 1779→2018, orbit ≥ 700,000 y; there is additionally a celestial-sphere neighbour ι Ori C, at mag. 9.8, with AC angular separation 49″ in 2002; ι Ori Aa is SB, not as yet resolved (so WDS is not as yet, at any rate as of 2022 Feb. 4, able to write “ι Ori Aa1”, “ι Ori Aa2”), with published orbital solutions (29.134 d, *e*=0.76, 0.11 AU min, 0.8 AU max): the elongated orbit, and a disparity in ages, suggest duplicity through many-body interaction-with-expulsion, rather than through the cogenesis which is usual for a binary; the ι Ori Aa, ι Ori Ab pairing does not for its part possess published orbital solutions ¶ colliding winds make ι Ori A a strong X-ray source ¶ ι Ori B is variable ¶ brightest member of Sword asterism ¶ E(B–V)=+0.07	Hatysa
ε	Ori A	5 37.4	−1 11	1.69v†	−0.18	B0 Ia	2	−7.2	2000	0.002	118	+26 SB	var. of α Cyg type, 1.64−1.74 in V supergiant, nonradially pulsating; superposition of many oscillations yields a light curve without a single obvious period ¶ luminosity (etc) controverted: Crowther (2006) 275,000 $L_{\odot}$, Searle (2008) 537,00 $L_{\odot}$, Puebla (2015) 832,000 $L_{\odot}$ ¶ E(B–V)=+0.08	**Alnilam**
ζ	Tau	5 39.0	+21 09	2.97v†	−0.15	B2 IIIpe (shell) †	7	−2.7	400	0.020	175	+20 SB†	var., eclips. (and γ Cas?), 2.80–3.17 in V, 133.0 d γ Cas variability would be consistent with the observed Be-phenomenon-cum-shell, and is accepted by AAVSO(VSX), although not accepted throughout the literature ¶ the primary in the SB pairing is one of the best-known “Be phenomenon” stars, and is possibly one of the keys to the solution of currently unsolved Be-phenomenon problems; consistently with the shell-spectrum history, the disk is just 5° away from being seen edge-on	Tianguan

													(2013A&ARv..21...69R, p. 58n); although the disk gases move in Keplerian orbits, their orbits are not circular, and consequently the material has some nonzero radial velocity even at the midpoint of transit; a further consequence of this kinematics is that the orbiting gas is less dense at apastron than at periastron; shell spectrum underwent three full cycles of V/R variation from 1997 to 2010, with these cycles generally taken as making the precession, under gravitational influence of the elusive SB companion, of a one-armed density wave within the Be disk (for geometry and time variations of disk, cf Fig. 7 of 2010AJ....140.1838S, Fig. 8 of 2015A&A...576A.112E); however, in more recent years, the V/R cycling has been absent; precession notwithstanding, the disk has been observed to be stable, and therefore must be being fed by decretion from the host-star photosphere at a nearly constant rate; as a step toward the eventual discovery of the excitation structure of some conveniently observable Be-phenomenon disk, 2012ApJ...744...19K reports spectro-interferometry from two different ζ Tau primary-star radii, in hydrogen Brackett γ and in a set of hydrogen Pfund lines (while drawing also on hydrogen Balmer α data from previous literature); the emission is found to originate at roughly the same disk radius for hydrogen Balmer α and hydrogen Brackett γ, and at a smaller radius for the hydrogen Pfund lines; the 2012ApJ...744...19K ζ Tau study provides some observational support for the viscous decretion-disk, Keplerian-rotation model prevalent in recent Be-phenomenon theorizing, and additionally supports the density-wave-in-disk hypothesis for the observed V/R cycles; modelling efforts are ongoing, with 2015A&A...576A.112E serving as a progress report ¶ its rapid rotation and Be-phenomenon and shell-spectrum histories notwithstanding, the ζ Tau primary has already evolved some distance off the MS, to "giant" stage (in general, giants are not expected to be rapid rotators); 2012ApJ...744...19K assumes an equatorial radius of 7.7 $R_\odot$ ¶ the nature of the elusive low-flux? ~1 $M_\odot$ SB companion, of period 133.0 d, is unknown (could even be a neutron star); separation (with orbit nearly circular) is ~1.17 AU; since interferometry seems so far to have failed to resolve the companion, WDS, at any rate as of 2022 Feb. 4, is constrained to write "ζ Tau" rather than "ζ Tau A" and "ζ Tau B"; the elusive companion may be producing a truncation in the Be-phenomenon disk, in the sense of a radical change in the dependence of disk density on radius (2013A&ARv..21...69R) ¶ under IAU rules, the name "Tianguan" applies only to the primary, not to the entire SB system	
α	Col A	5 40.5	−34 04	2.65†	−0.12	B7 IVe†	12	−1.9	260	0.025	176	+35 V?	slight var. of γ Cas type, 2.62−2.66 in V rapid rotator, with mass loss to disk, and so an instance of the "Be phenomenon"; Hα is variable, and Hβ profile varies rapidly; nevertheless, the Be disk is stable (unlike, e.g. the Be disk of γ Cas), indicating that the process of decretion-from-photosphere is in this case proceeding at a constant rate ¶ E(B–V) =0.00	Phact
ζ	Ori Aa	5 41.9	−1 56	1.74	−0.20	O9.5 Ib	4	−5.0	960	0.005	58	+18 SB2†	B: 3.7, B0 III, 2.4″, PA 152°→167°, 1822→2017 orbit ≥ 1500 y; B is a very rapid rotator, and also is possibly a β Cep variable, with departures from the classic rotational-broadening spectral profile (2013A&A...554A..52H is first presentation of high-quality spectroscopy for ζ Ori B); there is additionally a celestial-sphere neighbour ζ Ori C, of mag. 9.6 (AC astrometry: 56″→58″, PA 8°→10°,	Alnitak

													1839→2017), said by WDS to be possibly a physical companion; ζ Ori Aa,Ab now possesses an orbital solution from a combination of high-resolution spectroscopy (Aa,Ab is now found to be SB2, not merely SB) with full-period NPOI interferometry campaign (2013A&A...554A..52H; period is 7.3 y); since minimum distance of the O-star ζ Ori Aa from the B1 IV star ζ Ori Ab is 9.5 AU, this system is a good candidate for mass determinations (helpfully for the determinations, the distance, even at periastron, precludes a significant mass-transfer stream linking Aa with Ab; and in fact ζ Ori Aa is not only the brightest of the O-type stars in our visual sky, but is the first O-giant to have been assigned a mass via orbital computation); although there is no mass transfer between ζ Ori Aa and ζ Ori Ab, nevertheless Aa does eject mass vigorously, consistently with its membership in MK type O; at least one orbital solution has also been published for Aa,Ab-with-B (where ζ Ori B experiences the Aa,Ab binary as essentially a point mass) ¶ we here take the distance D from the Solar System (in our notation, and writing here to just 2 significant figures) 960 ly suggested in 2013A&A...554A..52H on the basis of orbital solution (the suggestion is made with 7% uncertainty, and additionally with the caveat that ζ Ori B photometry would indicate a larger D, in our notation ~1300 ly); from this D we deduce the corresponding value of π, as 4 mas; *HIPPARCOS* 2007 stated instead a different D (in our notation, and to just one significant figure, D = 700 ly) ¶ 2013A&A...554A..52H suggests age of ~7 My (but elsewhere a still lower age, below 4 My, has been suggested) ¶ E(B–V) =+0.09	
ζ	Lep	5 48.0	−14 49	3.55	0.10	A2 Vann†	~46.3	1.9	~70.5	0.015	266	+20 SB?	rapid rotator (period ~0.2 d or ~0.3 d) ¶ has debris disk, has first known extrasolar asteroid belt ¶ Astron. Alm. (epoch 2021.5) assigns MK type A2 Van ¶ approached to within ~4 ly or ~5 ly of Sun ~1 My ago	
κ	Ori	5 48.8	−9 40	2.07†	−0.17	B0.5 Ia†	5	−4.4	600	0.002	131	+21 V?	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]** slight var. of α Cyg type, 2.04−2.09 in V evolutionary status unclear, high mass-loss rate ¶ carbon-deficient (with metallicity otherwise unremarkable) ¶ E(B–V) =+0.07	Saiph
β	Col	5 51.8	−35 46	3.12	1.15	K1.5 III†	37.4	1.0	87	0.408†	8	+89† V	although high space velocity indicates that β Col is an interloper from outside galactic thin disk, nevertheless this star is richer than Sun in the elements beyond He	Wazn
α	Ori Aa†	5 56.4	+7 25	0.45v†	1.50	M2 Iab†	7†	−5.5	500†	0.030†	68	+21† SB	semireg., late-type supergiant var.: ~0–1.7 in V variability was discovered by J. Herschel in 1839; the latest minimum, early in 2020, at ~1.7 in V, sank below even the minima of 1927 and 1941 (each ~1.2); journalism on this 2020 event includes www.sciencenews.org/article/betelgeuse-star-dim-supernova-death-what-happened; three currently offered explanations are dust cloud from mass ejection, (gigantic) starspot, and fortuitous coincidence of minima from three separate cyclical variations; recovery began 2020 Feb. 22, with a rise to ~0.3 in V by 2020 late April; AAVSO reports mag. 0.65 in V on 2021 Jan. 11,	**Betelgeuse**

0.60 in V on 2022 Feb 4;
2018A&A...615A.116M suggests on
basis of magnetic variations
a scenario on which evolution of giant
photospheric convective cells, generating
magnetism through local
dynamos, is responsible for the
observed long secondary ~2100-day photometric period;
there are additionally ~200- ~400-day photometric
periodicities, plus a stochastic variation
ascribed to photospheric granulation
¶ Astron. Alm. (epoch 2021.5) assigns MK type
M1–M2 Ia–Iab
¶ brightest star in IR sky, also brightest in bolometric sky
¶ nearest RSG (contrast with o Cet, as AGB);
greatest angular diameter of almost any star other than Sun
(near-IR limb-darkened disk ~42 mas;
but R Dor, having approx 1/3 radius
of α Ori, is less distant, and
so attains still greater angular diameter);
en.wikipedia.org/wiki/List_of_largest_stars
supplies context, giving radii for many supergiants;
reduction of α Ori angular diameter
over period 1993/2009 has been asserted
¶ 2017AJ....154...11H reviews the
longstanding α Ori
distance problem: parallaxes, including *HIPPARCOS*,
labour under the difficulty of
accurately determining photocentre of visually
extended object, awkwardly harbouring
even plumes and hotspots;
we now give in our table these authors' values for π
(rounding from their 4.51 mas) and
by implication for D (strictly 717 ly ± 20%)
¶ very slow rotator (true period difficult;
8.4 y has been suggested)
¶ 2019A&A...628A.101H announces
dust halo with inner radius 1.5 R*;
2016A&A...585A..28K locates 3 R* as the interface between
hot-gas and more outlying dust envelopes
¶ CO shells inner 50 R* to 150 R*, outer as far as 250 R*
¶ runaway star, exceeding local
speed-of-sound in ISM: bow shock 6′–7′,
from stellar wind meeting ISM,
plus linear bar at 9′ (it has been suggested that the
bar is a relic of collapsing wind from a previous BSG phase,
and it also has been suggested that the bar is a feature
intrinsic to the embedding ISM, unconnected with
any α Ori Aa wind)
¶ although RSG
stars pose a more serious mass-loss problem for astrophysics
than do the AGB stars,
since it is not immediately clear
what mechanism is lifting RSG
stellar material above the photospheres
(convection? pulsation? magnetics?), there is now a possible
partial resolution in this particular case:
2018A&A...609A..67K, using ALMA,
finds α Ori anisotropic mass loss, with plume of ejecta;
the authors suggest that plume is associated with strong
"rogue" convection cell, observable as photospheric hot spot
(in contrast with the cool spots encountered on
such MS stars as the Sun)
¶ progenitor mass possibly ~20 $M_\odot$ (making
α Ori very massive),
age since arrival on ZAMS possibly
8.0-8.5 My (making α Ori very young)
¶ present evolutionary status of α Ori uncertain:
has this RSG previously been a BSG?
(and 2017MNRAS.465.2654W
suggests history may have been
complicated by a stellar merger)

													¶ α Ori is SN Type II-P progenitor, the core collapse being due within, (perhaps much within) 1 My: although SN will plateau for several months, yielding a star visible even in daytime, with the brilliance of a quarter Moon or full Moon, the SN radiation from so distant a source will not constitute a terrestrial biohazard ¶ *Sky & Telescope.* feature article 2019-05 on α Ori can usefully be supplemented with Fig. 13 from 2018A&A...609A..67K (multi-wavelength composite, showing ejecta plume condensing to dust at a few R*, and showing also two areas of local photospheric magnetic activity): AAVSO has backgrounder at www.aavso.org/vsots_alphaori ¶ WDS documents the putative detection, from a small amount of work in speckle interferometry, of two close companions (and is therefore compelled to write "α Ori Aa", "α Ori Ab", and "α Ori Ac"; since we reproduce WDS designations, we are in turn obliged to refer to Betelgeuse not as α Ori A but as α Ori Aa); however, since the WDS-documented speckle interferometry observations are from no later than 1983, and since current interferometry detects no close companions, it is now likely that Betelgeuse is unperturbed by any other star (the very faint WDS-catalogued stars α Ori B,C,D, E,F,G,H;I,J all lie on the celestial sphere at large angular separations from Betelgeuse, with the smallest angular separation, between A and B, measured as a quite wide 38″ in 2014; of these faint celestial-sphere neighbours, even the brightest, α Ori E, shines at a mere mag. 11)	
β	Aur Aa,Ab	6 01.2	+44 57	1.90†	0.08	A1 IV + A1 IV	~40.2	−0.1	81	0.056	269	−18 SB2	slight ecl.: 1.89–1.98 in V, 3.96 d (mags. equal) orbit is found in the published orbital solutions to be either circular or nearly circular (possibly *e*=2.8e−06); β Aur Aa,Ab was spectroscopically identified as a binary in 1890 (and is said to be only the third binary ever to be spectroscopically identified); orbit has been studied interferometrically since 1990s ¶ under IAU naming rules, "Menkalinan" denotes Aa, not Ab	Menkalinan
θ	Aur A	6 01.3	+37 13	2.65†	−0.08	A0p II: (Si)†	~19.7	−0.9	166	~0.086	~149	+30 SB	B: 7.2, G2 V, 4.2″, PA 7°→304°, 1871→2019 orbit ≥ 1200 y, with AB distance ≥ 185 AU ¶ A is magnetic, and an oblique rotator; there are abundance anomalies in photosphere patches, with Si and Cr 10× and 100× solar, respectively; consistently with rotation, θ Aur A is weakly var., period 1.37 d, mag. 2.62−2.70; although it is tempting to suggest α CVn-type variability, 2007A&A...464.1089S finds that observed variations in Hα, Hβ, and Hγ profiles cannot be modelled with photosphere inhomogeneities, and instead proposes changes in atmospheric pressure structure, as ions moving in the star's magnetic field undergo Lorenz-force deflections; the slight θ Aur A photometric variability is not, at any rate as of 2021 Nov. 4, documented in AAVSO(VSX)	Mahasim
η	Gem A†	6 16.2	+22 30	3.31v†	1.60	M3 III	8	−2.0	400	~0.064	~259	+19 SB	var.: 3.1−3.7 in V, 2979 d; B: 6.2, 1.8″ (2020) (AB astrometry in detail: 1.1″→1.8″, PA 300°→258°, 1881→2020); a potentially confusing blend of two variabilities: the unresolved single-line SB which is η Gem A is an Algol-type eclipsing system, with each eclipse lasting several weeks (most recent minimum around 2020 Oct. 22, with mag 3.766 in V band reported at AAVSO; next eclipse may therefore be expected to begin late in 2028); additionally, however, one or the other component of this SB presents semiregular Mira-type instability, with	Propus

													one or more periods, average period 234 d ¶ liable to lunar, and also to very rare planetary, occultations (making η Gem A not only an SB, but also an occultation binary)	
ζ	CMa Aa,Ab	6 21.2	−30 04	3.02†	−0.16	B2.5 V	9.0	−2.2	360	0.008	61	+32 SB†	SB is recently resolved, as ζ CMa Aa, ζ CMa Ab, with just 4 observations (from 2019 and 2020) documented in WDS as of 2021 Nov. 4 (Aa,Ab angular separation surely much less than 1″); SB period is 675 d; WDS asserts mags. 3.6, 3.8; the pairing with ζ CMa B (mag. 7.8) is wide (167″→170″, PA 338°→340°, 1833→2016) ¶ variability has been claimed somewhere in what is now resolved as the ζ CMa Aa,Ab pair (with membership claimed in the β Cep pulsator class: AAVSO(VSX) has flagged this as a suspected variable, but is (at any rate as of 2021 Nov. 4) unable to give a magnitude range)	Furud
β	CMa A	6 23.7	−17 58	1.98†	−0.24	B1 II–III	7	−3.9	~490	0.003	256	+34 SB	slight var., β Cep type, 1.97–2.00 in V, 0.25130 d (we give here the AAVSO(VSX) period and V-mag. range, as viewed 2021 Aug. 07); the brightest of the β Cep pulsators; has multiple modes, with beat period 50 d; it is not known why ε CMa, while physically similar, is not a pulsator ¶ near the boundary of the “Local Bubble” ISM cavity ¶ E(B–V) =+0.01	Mirzam
μ	Gem A	6 24.3	+22 30	2.87v	1.62	M3 IIIab	14	−1.4	230	0.124	153	+55 V?	semiregular variable: 2.75–3.02 in V ¶ on AGB ¶ subject to lunar occultations	Tejat
α	Car	6 24.5	−52 43	−0.62	0.16†	A9 Ib†	11	−5.5	~310	0.031	41	+21	visible both in X-ray (magnetically heated corona; also rapid rotator, strongly convective) and in radio ¶ evolutionary status not fully clear, and colour unusual in its luminosity class ¶ Astron. Alm. (epoch 2021.5) assigns MK type A9 II	**Canopus**
ν	Pup	6 38.4	−43 13	3.17	−0.10	B8 IIIn†	9	−2.1	370	0.004	186	+28 SB	rapid rotator, with period < 1.7 d shell spectrum has been suggested, with “central quasi-emission peak” (cf 1999A&A...348..831R) ¶ distance was ~27 ly 3.6 My ago	
γ	Gem Aa	6 39.0	+16 23	1.93	0.00	A1 IVs	30	−0.7	110	0.057	166	−13 SB†	γ Gem Aa,Ab is SB system in highly elongated orbit, known historically as an occultation binary (the brightest ever to be observed in an asteroid occultation: 381 Myrrha, in 1991) and as SB, but also reported in 2014 as resolved with adaptive optics at the USA military facility “Starfire Optical Range”, thereby facilitating study of component masses (2014AJ....147...65D, Fig. 6: this author finds period 12.634 y, *e*=0.89, in good agreement with period and eccentricity from other published orbital solutions for γ Gem Aa,Ab); average Aa,Ab distance is ~8.5 AU; Ab is mag. ~7.5 ¶ E(B–V) =+0.03	Alhena
ε	Gem A	6 45.3	+25 06	3.06	1.38†	G8 Ib	4	−4.0	800	0.014	204	+10 SB	unusually yellow in the general population of supergiants ¶ among the few supergiants liable to lunar and planetary occultations ¶ celestial-sphere neighbour ε Gem B (112″→110″, PA 93°→94°, 1825→2013) is mag. ~9.6	Mebsuta
α	CMa A†	6 46.1	−16 45	−1.44	0.01	A0mA1 Va†	~379	1.5	8.6	~1.339	~204	−8 SB	B: 8.5, WDA; 11.2″ (2020); orbit 50.1 y separation 8.2 AU min (3″), 31.5 AU max (˘11″, in 2019); although α CMa AB, now among the more celebrated visual binaries, was first resolved in 1862 (Alvan G. Clark, using 18.5-inch refractor), binarity was conjectured as early as 1834 (by Bessel, on the grounds of variability in proper motion); cdsarc.u-strasbg.fr/viz-bin/getCatFile_Redirect/?-plus=-%2b&B/wds/./notes.dat, as notes for “WDS 06451-1643”, has vivid and extended discussion of 19th-century struggles with obtaining accurate angular-separation and PA	**Sirius**

													measurements for the α CMa AB system; as of 2021 Nov. 6, WDS reports the existence of 2061 measurements, for the timespan 1862→2020 (in contrast with the more difficult, and less well observed, α CMi AB system, where the secondary is again a WD, and the primary is again an intrinsically luminous (~7 Suns; Sirius radiates with the power of ~25 Suns) and nearby and notably hot star, but where smaller angular separation and greater magnitude difference pose a greater challenge: as of 2021 Nov. 6, WDS reports the availability of just 99 α CMi AB measurements); α CMa B is the brightest WD in the visual sky (its nearest competitors being o^2 Eri B, at mag. 9.5, and α CMi B, at mag. 10.8); www.atnf.csiro.au/outreach//education/senior/astrophysics/binary_types.html compares and contrasts a visual-wavelengths image of the α CMa AB pair (with A much brighter than B) and a *CHANDRA* X-ray image of this same pair (with B much brighter than A; A is said to appear in the image largely through reflecting UV emitted by the much hotter B) ¶ *IRAS* detected IR excess, a signature of dust (rather unexpected in a binary) ¶ Fe abundance of α CMa is ~2× or ~3× solar ¶ α CMa B is unusually massive for a WD (1.02 $M_{\odot}$; Chandrasekhar Limit is, however, 1.4 $M_{\odot}$; spectral type of α CMa B is DA (= hydrogen-only)) ¶ E(B–V) =-0.03	
ξ	Gem	6 46.6	+12 52	3.35	0.44	F5 IV	56	2.1	58.7	0.223	211	+25 V? †	possibly SB, with components of ~equal mass ¶ rapid rotator (but just barely over the internal-structure transition, or "F5 rotation break," that causes some stars to rotate rapidly, others to experience braking through magnetics and winds) ¶ X-ray source (suggesting significant corona)	Alzirr
α	Pic	6 48.4	−61 58	3.24	0.22	A6 Vn kA6†	~34	0.9	100	0.252	345	+21	rapid rotator; shell, with time-varying spectral absorption features ¶ X-ray emission suggests a companion, otherwise undetected	
τ	Pup	6 50.5	−50 39	2.94	1.21	K1 III	18	−0.8	180	0.077	154	+36 SB†	SB period 1066.0 d, separation ~3 AU, orbit of low eccentricity; since the SB is not as yet resolved, even in interferometry, WDS is not as yet able to write "τ Pup A" and "τ Pup B"	
κ	CMa	6 50.7	−32 32	3.50v†	−0.12	B1.5 IVne†	4.9	−3.0	700	0.010	293	+14	var., γ Cas type, 3.40–3.97 in V (was at faint end of its range before 1963; AAVSO reports visually ~3.3 in 2021 Jan.); an instance of the "Be phenomenon"	
ε	CMa A†	6 59.5	−29 00	1.50	−0.21	B2 II	8.0	−4.0	410	0.004	68	+27	B: mag.7.5 (7.9″, PA 161°→162°, 1850→2008) AB distance 900 AU, period at least 7500 y ¶ brightest known source of extreme UV (~75 nm) in Earth's night sky; hydrogen Lyman α (121.6 nm) observed by NASA *OAO-3* ¶ E(B–V) =+0.02	**Adhara**
σ	CMa A	7 02.6	−27 58	3.49v	1.73	K7 Ib†	3	−4.2	1100	0.008	308	+22	irregular var.: 3.41–3.51 in V authorities are in some disagreement on MK type (possibly M, rather than K) **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	Unurgunite
o^2	CMa	7 04.0	−23 52	3.02†	−0.08	B3 Ia	1	−6.6	3000	0.004	329	+48 SB	slight var., α Cyg type, 2.98–3.04 in Hp, 24.44 d the α Cyg vars. are non-radial pulsators ¶ E(B–V) =+0.03	
δ	CMa	7 09.3	−26 26	1.83	0.67	F8 Ia†	2	−6.6	2000	0.005	317	+34 SB	¶ slow rotator (possibly ~1 y); N 2× solar, Na 6× solar	Wezen
L_2	Pup A	7 14.2	−44 41	4.42v†	1.33	M5 IIIe	16	0.4	210	0.342	18	+53 V?	semireg. late-type var., 2.6–8.0 in V, 140.6 d	HR2748
π	Pup Aa	7 17.9	−37 08	2.71v†	1.62	K3 Ib	4	−4.3	800	0.012	303	+16	B: 7.9, 66″, PA 214°→213°, 1826→2009 Aa,Ab system has received only 2 observations, both in 1991 (0.7″ and again 0.7″, PA 148°	

													and 152°), with Ab at mag. ~6.5 ¶ semiregular variable.: 2.70–2.85	
δ	Gem A	7 21.5	+21 56	3.50	0.37	F0 IV†	54	2.2	60	0.018	237	+4 SB†	B: 8.2, K3 V, 5.5″, PA 198°→229°, 1822→2018 orbit 1200 y ¶ lunar occultations possible; planetary occultations possible-yet-rare: since the SB which is δ Gem A is not as yet resolved (even interferometrically) with astrometry, WDS is not as yet (at any rate as of 2022 Feb. 5) able to write "δ Gem Aa", "δ Gem Ab"; SB period is 6.129y; a companion (the secondary component in the SB?) has, however, been noted through occultation ¶ in evolutionary transition, having completed stable core-hydrogen fusion ¶ Astron. Alm. (epoch 2021.5) assigns MK type F0 V+ **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	Wasat
η	CMa A	7 25.0	−29 21	2.45v†	−0.08	B5 Ia	2	−6.5	2000	0.007	325	+41 V	B: 6.8, 178″ (2010) is mere optical companion ¶ variable in α Cyg class of non-radial pulsators; AAVSO(VSX) as viewed 2021 Jan. 28 and again 2022 March 3, gives mag. range 2.36–2.50 in V, period 4.70433 d ¶ strong wind; ejected circumstellar mass inferred from IR excess ¶ E(B–V) =+0.02	Aludra
β	CMi A	7 28.4	+8 15	2.89†	−0.10	B8 Ve†	~20.2	−0.6	~162	0.064	234	+22	rapid rotator, possibly ~1 d, with modest variability in the hydrogen Balmer emission; disk of ejected matter has diameter ~4× diameter of β CMi itself (BSC5: "rotationally unstable"); an instance of the "Be phenomenon"; although GCVS and AAVSO(VSX) assertion of γ Cas-type variability has not been corroborated, 2007ApJ...654..544S reports, using MOST, millimagnitude "slowly pulsating B-type" variability; AAVSO(VSX) as viewed 2021 Jan. 16 gives V-mag. range 2.84–2.92, but as viewed 2022 March 3 the narrower range 2.89–2.90; in contrast with e.g. the Be-phenomenon star γ Cas A, the Be disk is in this case considered very stable (2013A&ARv..21...69R), indicating constancy in the process of decretion from the host-star photosphere; 2012ApJ...744...19K reports confirmation of Keplerian rotation in the Be disk (an important follow-on to the discovery of Keplerian rotation in Be-phenomenon star α Ara A); 2019ApJ...875...13H finds that, contrary to an earlier claim that β CMi A is SB with period 170.4 d, there is so far no convincing evidence of SB status; the question is significant for overall Be theory, in that duplicity has been suggested as a possible mechanism generating the "Be phenomenon" in at least some stars	Gomeisa
σ	Pup A	7 29.9	−43 21	3.25†	1.51	K5 III	17	−0.6	190	0.198	342	+88 SB†	B: 8.8, G5: V, 22.1″, PA 90°→74°, 1826→2015 orbit ≥ 27,000 y, A-to-B distance ≥ 1300 AU; SB is eclipsing, of β Lyr type, with orbit 257.8 d, with very modest alternating primary (0.04 mag) and secondary (0.03 mag) minima; the SB primary component shows slow irregular variability ¶ system has high space velocity	
α	Gem A†	7 36.0	+31 50	1.93	0.03	A1mA2 Va	63	0.9	52	~0.254	~234	+6 SB	orbit 445 y; max = 6.5″, in 1880;	
α	Gem B†	7 36.0	+31 50	2.97	0.03	A2mA5 V:	63	2.0	52	~0.254	~234	−1 SB	min = 1.8″, in 1965; 5.5″ (2019); separation 71 AU min, 138 AU max; C mag. 9.8; AC 70″, PA 162°→163°, 1822→2017, orbit ≥ 14,000 y; orbital solutions have been published both for the binary α Gem AB and for the much wider binary involving C (with C experiencing AB as a point mass); C has variable-star name YY Gem (an eclipsing binary, and additionally a variable of the BY Dra class, with flaring); not only C, but also each of A, B is itself SB, making ABC a hierarchical 6-star system	Castor

													(Kaler at stars.astro.illinois.edu/sow/castor.html writes, “certainly the sky’s ranking sextuple”); en.wikipedia.org/wiki/Castor_(star) has a diagram summarizing this sextuple hierarchy, on the basis of 2012MNRAS.423..493H; since the A SB is not yet resolved (even interferometrically) and since the B SB is not yet resolved (even interferometrically), WDS is not yet able to write “Aa,Ab” and is not yet able to write “Ba,Bb” ¶ Castor-Pollux comparison is a helpful test of naked-eye night colour response	
α	CMi A†	7 40.5	+5 10	0.40†	0.43	F5 IV–V	285	2.7	11.5	~1.259	~215	−3 SB	B: mag. 10.8, WD; 3.8″ (2014); orbit 40.84 y (PA 286° in 2014; this is the most recent astrometry available in WDS as of 2022 Feb. 5) with e=0.4; 2015ApJ...813..106B gives α CMi AB periastron distance as 9.1 AU; α CMi B is visually the third-brightest WD in the sky (overtaken by o^2 Eri B, at mag. 9.5, and by α CMa B, at mag. 8.5); 2015ApJ...813..106B is a recent study of masses, with orbital solution, drawing on observations beginning in the 19th century and including 1995−2014 HST data (Fig. 4 makes the problem of orbit-fitting, from good recent data and less good historical data, vivid), and discussing also the possible evolutionary history of the system (past mass transfer may be a complicating factor) ¶ α CMi A radiates with the power of ~7 Suns (and so is not dramatically unlike α CMa A, which radiates as ~25 Suns) ¶ asteroseismology of α CMi A is somewhat uncertain (MOST mission 2004 did not find pulsations, and yet WIRE mission 1999 and 2000 did) ¶ the WD α CMi B is physically unlike the WD α CMa B, attaining only ~0.2 of the density of α CMa B, and being of a rare spectral type DQZ6.5 (elusive in ground-based spectroscopy: but at long last *(a)* helium-not-hydrogen character of spectrum was noted at HST through filter photometry in a set of bands running from 1600 Å to 7828 Å, and *(b)* detailed spectroscopy was performed at HST through STIS camera, over the range 1800 Å to 10000 Å; in this spectral type, H features are absent and C, Mg, and Fe features are present) ¶ from an astrometrist’s perspective, the α CMi AB system contrasts with the less difficult, and consequently better measured, α CMa AB system, in which the secondary is again a WD, and the primary is again an intrinsically luminous and nearby and notably hot star, with the magnitude difference less severe and the typical angular separation greater: as of 2021 Nov. 6, WDS reports the existence of just 99 α CMi AB astrometry measurements, as against the much larger tally of 2061 α CMa AB measurements; Bond et al.write, in 2015IAUDS.187....1M, “Charles Worley, double-star observer at the USNO, asserted to two of us, more than two decades ago, that he was the only living astronomer who had seen Procyon B with his own eye” ¶ WDS documents, for the α CMi AB system, celestial-sphere neighbours α CMi C,D,E,F,G,H; of these, only E and G are brighter than mag. 10, at mags. 9.2 and 8.8 respectively (at the large respective angular separations, from α CMi A, of 467″ (2009) and 356″ (2012))	Procyon
β	Gem A+1P†	7 46.7	+27 58	1.16	0.99	K0 IIIb	97	1.1	33.8	0.628	266	+3 V	the nearest of the giants; unusual in being a giant known to harbour an exoplanet (and the brightest known exoplanet host in Earth’s sky); as of 2015,	Pollux

													exoplanet is IAU-named “Thestias” ¶ subject to rare lunar occultations, for observers S of Earth’s equator ¶ Castor-Pollux comparison is a helpful test of naked-eye night colour response	
ξ	Pup A†	7 50.2	−24 55	3.34	1.22	G6 Iab–Ib†	3	−4.5	1200	0.005	260	+3 SB†	full system comprises an unresolved tight binary (an SB) and widely separated ξ Pup B, with B experiencing the unresolved SB which is ξ Pup A as essentially a point mass (B is mag. 13; 4.6″→5.1″, PA 189°→191°, 1899→1964; orbit ≥ 26,000 y?) ¶ SB primary has high metallicity, with exact evolutionary status uncertain ¶ SB primary is near, but is a little too cool to lie within, the HR diagram IS	Azmidi
χ	Car	7 57.4	−53 03	3.46†	−0.18	B3 IV(p?)†	7	−2.3	500	0.035	304	+19 V	Si II anomalous strength now discounted ¶ suggestion of variability now discounted, via *HIPPARCOS* ¶ the MK luminosity class “IV” (phenomenologically “giant”) notwithstanding, χ Car is in astrophysical terms in the last part of its stable core-hydrogen-fusion phase; Astron. Alm. (epoch 2021.5) assigns MK temperature type B3p Si without assigning an MK luminosity class	
ζ	Pup	8 04.4	−40 04	2.21†	−0.27†	O5 Iafn†	3.0†	−5.4	1080†	0.034†	299	−24† V?	blue supergiant ¶ rapid rotator (1.78 d), despite ~2300 km/s stellar wind (in which spiral structure was announced in 2017 by BRITE mission team), with mass loss rate > 1e-6 $M_\odot$ /y ¶ high space velocity (impelled by past nearby supernova? or, rather, impelled by multibody gravitational interactions in its stellar birth family?); possibly ejected from Trumpler 10 OB association ¶ distance has been controverted ¶ He, N overabundant ¶ has been suspected of being a variable of the α Cyg type ¶ E(B–V) =+0.04	Naos
ρ	Pup A	8 08.5	−24 22	2.83v†	0.46	F2mF5 II: (var)†	51.3	1.4	64	0.095	299	+46 SB	var.: 2.68–2.87 in V, 0.14 d prototype of the “ρ Pup stars” (these combine δ Sct variability with Am-like abundance anomalies); main period is ~3.3 h (0.15 mag.); photosphere temperature is notably low in the overall population of stars presenting δ Sct variability ¶ Astron. Alm. (epoch 2021.5) assigns MK type F5 (Ib–II)p ¶ IR excess (circumstellar ring, at separation 50 AU?)	Tureis
γ	Vel Aa†	8 10.2	−47 24	1.75	−0.14	O7.5 III-I†	3†	−5.9	~1100†	0.012	330	+35 SB2†	eruptive WR var.:1.81−1.87 in V; Aa is a.k.a. γ^2 Vel ¶ Aa is a double-lined SB not as yet resolved, even in interferometry (so WDS is not as yet able to write “γ Vel Aa1”, “γ Vel Aa2”), but with published orbital solutions (period 78.5 d, *e*=0.4 or 0.5); faint Ab (mag. 13.4) is poorly known, with just one astrometry result, from 1997 (Aa,Ab angular separation 4.7″); distance between components of the unresolved γ Vel Aa SB is 0.8(?) AU min, 1.6 AU max); γ Vel B, a.k.a. γ^1 Vel, is itself a resolved SB pair (so in WDS γ Vel Ba, γ Vel Bb: period is 1.48 d); AB astrometry is 43″→41″, PA 222°→221°, 1826→2017 ¶ the (carbon-rich) WR component in the unresolved SB which is γ Vel Aa is of spectral type WC8, and is the nearest and visually brightest of all WR stars (presenting “a unique opportunity to spatially resolve a WR wind by means of interferometry” (2007A&A...464..107M)), and is an exceptionally massive WR (9.0 $M_\odot$; but at birth, > 30 $M_\odot$); this SB is the best studied of all O-WR binaries: in the SB pair it is the WR component, rather than the O component, that dominates spectrally (although we assign an O type, and Astron. Alm. (epoch 2021.5) rather similarly assigns	

the uncertainty-flagged MK type
"O9 I:", since the V-band light is overwhelmingly
from the more massive (18.5 $M_{\odot}$) O-type
component), making the γ Aa SB the "Spectral
Gem of the Southern Skies", and a notable sight within
broader "Vela complex" (dominated by the
Gum Nebula, within which lie the Vela SNR, the
IRAS Vela shell, and the Vela pulsar; some literature,
including 2011A&A...525A.154S, indeed proposes
intersection between the Vela SNR and a γ Vel
SWB, taking the *IRAS* Vela shell as marking the meeting
of SNR and SWB)
¶ like η Car (bright to mag. ~0 for several years after
1837, but now too faint, and now too lacking in firm
future-outburst prognoses, to qualify for the RASC
Handbook "Brightest Stars" list), the unresolved
γ Vel Aa SB is a colliding-wind pair
(2017MNRAS.468.2655L Fig. 1 sketches the collision
geometry), and in consequence is a UV and X-ray source
(and in consequence may also possibly resemble η Car
in being a γ-ray source
(cf 2017ApJ...847...40R; as of
at any rate 2017, it seems that
no other colliding-winds-binary
stellar γ-ray sources are known));
it is the wind from the (WR) secondary
that dominates, with mass-loss
rate at least 100× greater than for the (O-type) primary;
the WR wind may feature some clumping,
but is to a good approximation
spherically symmetric until it
encounters the O-star wind;
orbital motion of the two SB components
around their centre of mass
yields a spiral structure in the wind-collision area,
particularly salient at periastron
¶ 2017ApJ...847...40R summarizes recent observations
of the Aa Vel SB, in radio and IR
and optical, including interferometry,
noting inter alia discrepancies in the available determinations
of mass-loss rates from the WR star (a copious 3e-6 $M_{\odot}$ /y?
or a still more copious 8e-5 $M_{\odot}$ /y?)
¶ notable among recent observational studies are
2017MNRAS.468.2655L (VLTI/AMBER
near-IR spectro-interferometry,
with also 3-D hydrodynamic modelling)
and 2012MNRAS.427..581R
¶ likely destiny of γ Vel Aa
(WR) secondary is as (exotic) stripped-core
SN (same prognosis as for η Car; this
contrasts with α Ori, which
will for its part instead explode as a (not exotic)
hydrogen-spectrum SN)
¶ dust emission is absent (even though
formation of circumstellar dust is common
in stars that, like the γ Vel Aa (WR) secondary, undergo
copious mass outflow)
¶ distance ~1200 ly, in contrast with our
~1100 ly, has also been recently asserted,
on basis of VLTI/AMBER
¶ we take MK type for γ Vel Aa from
1999A&A...345..163D (as what
must be considered an emendation
of our (slightly cooler)
Garrison-approved MK type
from earlier editions of this table;
admittedly, MK determination of γ Vel Aa is still
difficult, because the raw spectrum
is a composite comprising
not only the O and WR stars, but also emission from the
wind-collision zone)
¶ Ba,Bb appear in combined light as mag. 4.1; there

													are additionally wide celestial-sphere neighbours C (mag. 7.3; AC separation was 62″ in 2009) and D (mag. 9.2; AD separation was 94″ in 2000) ¶ neither the traditional Suhail al-Muhlif nor the modern Regor (devised within NASA, to commemorate 1967 fire victim Roger Chaffee) is presently IAU-approved name for any of the five stars γ Vel Aa primary, γ Vel Aa secondary, Ab, Ba, Bb	
β	Cnc A+1P	8 17.7	+9 07	3.53	1.48	K4 III†	11	−1.3	300	0.068	224	+22 V?	“barium star,” with Ba abundance ~6× solar, presumably as contamination from defunct companion (but no companion remnant has been found) ¶ Astron. Alm. (epoch 2021.5) assigns MK type K4 III Ba 0.5	Tarf
ε	Car A	8 23.0	−59 35	1.86v†	1.20	K3:III	5	−4.5	600	0.034	311	+2	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]** B is possibly ecl., with AB mag. 1.82−1.94 in V (system flagged by AAVSO(VSX), at any rate when viewed 2021 Nov. 8, as eclipse candidate; cf further 2004AJ....127.2915P, Table 5, which treats B as an unresolved binary: or is it that the eclipses (if real at all), involve not an unresolved “ε Car B primary”, “ε Car B secondary” pair, but more straightforwardly the ε Car A, ε Car B pair?); full WDS-catalogued system is ε Car A (mag. 2.2) and ε Car B (mag. 3.9), with ε Car B not further resolved, and with the ε Car AB pairing reported by WDS as very tight, the angular separation being just 0.4″ in 2019; the AB system is sparsely observed (WDS documents just 6 satisfactory astrometry measurements, 1991→2019); AAVSO general database (as distinct from AAVSO(VSX)) seems to have no record of observations, the 1959AJ.....64..127G report of variability notwithstanding; Kaler comments at stars.astro.illinois.edu/sow/avior.html, with reference to the overall paucity of observations, that “if there is a stellar category of ‘bright stars getting no respect’, [ε Car A] probably holds the record” ¶ B is of MK type (uncertainty-flagged) “B2: V” ¶ the IAU-official name “Avior” for ε Car A is of uncertain etymology, and yet its origins are known: here, as also with α Pav A (IAU-officially “Peacock”), the name stems from the 1930s RAF Air Almanac project, which directed HM Nautical Almanac Office that no air-navigation star was to be left nameless	**Avior**
o	UMa A+1P	8 32.1	+60 38	3.35†	0.86	G5 III	~18.2	−0.3	~179	0.172†	231	+20†	slight var.?: 3.30?–3.36 in V? ¶ currently in rapid evolutionary transition, crossing the Hertzsprung Gap ¶ despite high space velocity, a member of the galaxy thin disk	Muscida
δ	Vel Aa†	8 45.3	−54 48	~1.95v†	0.04	A1 Va	40	0.0	81	~0.107	~164	+ 2 V?	Aa,Ab brightest known ecl. binary (1.95−2.43, V) orbital period 45.15023 d, primary (resp. secondary) eclipse duration 0.587 d (resp. 0.91 d), with eclipse indicated by photometry to be total; system is Algol type, with average Aa,Ab distance 90.61 AU, resolved both interferometrically and with VLT adaptive optics; 2007A&A...469..633K offers orbital solution (0.23≤e≤0.37, with orbit inclination to plane of sky near one or other of the two extremes 87.5°, 92.5°), discusses masses, finds unexpectedly large stellar diameters (so both stars are evolved?); B experiences Aa,Ab as essentially a point mass, the Aa,Ab-with-B orbital period being 142 y (at least one orbital solution has been published); B is mag. 5.6, 2.2″→0.8″ (min. angular separation was in 2000), PA 177°→195°, 1894→2019 ¶ WDS documents faint celestial-sphere neighbours C,D,E, and also a sparsely observed close brighter celestial-sphere neighbour F (mag. 5.8, with just 3 observations 1991→1999)	Alsephina

ε	Hya A†	8 48.0	+6 20	3.38	0.68	G5:III	25	0.4	130	~0.232	259	+36 SB†	composite A: 3.8; B: 4.7, 0.2″ (2018); C: 7.8, 2.9″ (2020); B is of poorly known MK type "A:"; orbital solutions have been published both for AB (15.09 y) and for the much wider AB+C system (C experiences AB as essentially a point mass; period is 590 y) ¶ C is unresolved SB, with orbital period 9.9 d ¶ our π, D are from 2018 *Gaia* parallax, which is known to ±2% (since π is stated as 20.7182 ±0.3925 mas), and which we take to supersede 2007 *HIPPARCOS*	Ashlesha
ζ	Hya	8 56.6	+5 52	3.11	0.98	G9 II–III	~21†	−0.4	~157	0.101	279	+23	¶ Astron. Alm. (epoch 2021.5) assigns MK type G9 IIIa	
ι	UMa A†	9 00.7	+47 57	3.12	0.22	A7 IVn	~68.9	2.3	47.3	~0.491	~244	+9 SB†	A+BC 2.4″, PA 349°→90°, 1831→2017 A+BC orbit 818 y; BC 0.9″, period ~39 y; both the AB binary system and the AB,C binary system (in which C experiences the ι UMa A, ι Uma B binary as essentially a point mass) possess published orbital solutions; A is itself SB, orbit 4028 d, making this a quadruple system; the system is not, as in many cases of multiplicity, hierarchical and stable, but kinematically unstable (disruption in ~0.1 My?); B is mag. 9.9 M1 V and C is mag. 10.1 M1 V; since the A SB has not yet been resolved, even interferometrically, WDS is not yet able to write "ι UMa Aa", "ι UMa Ab" ¶ AAVSO(VSX) reports the ι UMa system as presenting some kind of as-yet-poorly-studied slight and rapid variability, with range 3.12−3.18 in V (and WDS indicates a possible line of research by indicating, in a note, that the ι UMa A SB pairing presents δ Sct variability)	Talitha
λ	Vel A	9 08.8	−43 31	2.23v†	1.66	K4 Ib–IIa†	6.0	−3.9	540	0.028	299	+18	semireg. var.: 2.14–2.30 in V ¶ probably on or approaching AGB, but could still be on RGB ¶ Astron. Alm. (epoch 2021.5) assigns MK type K4.5 Ib ¶ has slow wind, whose origins are said to be poorly understood	**Suhail**
a†	Car	9 11.6	−59 04	3.43†	−0.19	B2 IV–V	7	−2.3	500	0.022	312	+23 SB2†	slight ecl.: 3.41–3.44 in V ¶ orbit 6.74 d, with light curve indicating tidal distortion; since the SB is as yet unresolved, even in interferometry, WDS is not yet able to write "a Car A," "a Car B" ¶ there is some uncertainty whether observable light is solely from primary, or whether primary and secondary make approximately equal contributions ¶ not to be confused with α Car	HR 3659
β	Car	9 13.4	−69 49	1.67	0.07	A1 III	28.8†	−1.0	113	0.191	305	−5 V?	rapid rotator (< 2.1 d), despite having finished stable core hydrogen fusion ¶ quasi-periodic variation, ~0.5 h, in hydrogen Balmer lines	**Miaplacidus**
ι†	Car	9 17.7	−59 22	2.21†	0.19	A7 Ib	4.3	−4.6	800	0.022	302	+13	slight var.: 2.23–2.28 in V ¶ despite being slow rotator, has magnetic activity (as inferred from X-rays) ¶ not to be confused with l (letter el) Car	Aspidiske
α	Lyn A	9 22.4	+34 18	3.14†	1.55	K7 IIIab	16	−0.8	~203	0.224	274	+38	B: 8.8, 223″, PA 33°→43°, 1823→2016 suspected var., mag. 3.12–3.17 (beginning to evolve into a Mira?)	
κ	Vel	9 22.8	−55 06	2.47	−0.14	B2 IV–V	6	−3.8	600	0.016	315	+22 SB†	IAU name "Markeb" is not to be confused with "Markab" (the IAU name for α Per A) ¶ orbit 116.65 d, average separation possibly ~1.1 AU; since the SB has not yet been resolved (even interferometrically), WDS is not yet able to write "κ Vel A" and "κ Vel B" ¶ mass loss rate ~1e–9 $M_{\odot}$/y ¶ system is X-ray source ¶ ISM absorption has varied over the years (ISM cloud in transit?)	Markeb†
α	Hya A†	9 28.7	−8 45	1.99†	1.44	K3 II–III†	18	−1.7	180	0.038	336	−4 V?	poss. slight var., underobserved (1.93−2.01 in V?) ¶ slow rotator (possibly 2.4 y), with Ba mildly overabundant ¶ asteroseismology has been studied	Alphard

													¶ α Hya B (mag. 9.7; 284″, PA 55°→155°, 1833→2015) might be a true binary component (with orbit ≥ 870,000 y, separation ≥ 15,700 AU)	
N	Vel	9 31.9	−57 08	3.16†	1.54	K5 III	13.6	−1.2	240	0.033	280	−14	slight semiregular var., 3.12–3.18 in V, 82.0 d ¶ evolutionary status uncertain (helium core fusion impending, or already ended?)	HR 3803
θ	UMa A	9 34.4	+51 34	3.17	0.48	F6 IV†	74.2	2.5	44.0	1.088	241	+15 SB†	luminosity class, and also SB status, have been controverted, with postulated SB companion remaining undetected in speckle interferometry	
o	Leo Aa†	9 42.4	+9 47	3.52	0.52	F9 III + A5m	24.2	0.5	135	0.148	255	+27 SB2†	IAU name “Subra” applies only to o Leo Aa; Aa,Ab is a tight binary, with angular separation ~4 mas, but nevertheless now interferometrically resolved (2001AJ....121.1623H; from this paper we take MK type, and also our π-cum-D (as a distance derived from comparing seen angular size of orbit with orbit physical size, yielding a result in good agreement with the 2007 *HIPPARCOS* trigonometric parallax, and yet with smaller uncertainty); period is 14.498 d, with orbit nearly circular, distance between the two stars 0.165 AU; our assertion of eclipsing in earlier editions of the Handbook was erroneous, although it is true that the orbit is seen rather close to edge-on; lunar occultation as reported in 1978AJ.....83.1100A failed to split Aa,Ab ¶ o Leo Aa is a rare instance of a star that has ended core hydrogen fusion, and yet in which the convection typical of an evolved star has not removed the chemical peculiarities possible in a core-hydrogen fuser (where the still-quiet atmosphere facilitates radiative lofting and gravitational settling)	Subra
l†	Car	9 45.9	−62 37	3.69v†	1.01	F9–G5 Ib	2	−4.7	2000	0.015	302	+3 V	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]** Cepheid variable: 3.28–4.18 in V, 36 d AAVSO(VSX) as viewed both 2021 Jan. 18 2022 March 2 gives period 35.551609 d; an exceptionally luminous, and consequently exceptionally slow, Cepheid (compare both the visual brightness and the intrinsic luminosity with less dramatic δ Cep A (in this table), η Aql A (Okab; in this table), and ζ Gem Aa (Mekbuda; almost, but not quite, bright enough for inclusion in this table): Kaler remarks that “if Carina had been in the northern hemisphere, the collection of these variables might well have been called the ‘Carinids’”); radius, in its pulsation cycle, has been measured as 160 $R_\odot$ min, 194 $R_\odot$ max ¶ circumstellar envelope of ejected matter, radius 10 AU–100 AU ¶ lower-case ell Car; not to be confused with i (lower-case i) Car (HR 3663), ι Car (HR 3699), L Car (HR 4089), I (upper-case i) Car (HR 4102) (and note additionally that Bayer nomenclature does not use the label “λ Car”)	HR 3884
ε	Leo	9 47.1	+23 40	2.97†	0.81	G1 II	13.2	−1.4	250	0.047	259	+4 V?	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]** poss. slight var., underobserved (2.95−3.04 in V?) possible pulsation as in Cepheids ¶ slow rotator, period possibly as long as 200 d ¶ currently residing in the Hertzsprung Gap? ¶ variability has been studied (cf Andrievsky 1998; pulsation as in Cepheids?) ¶ the Arabic or quasi-Arabic name Algenubi (more classically, al Ras al Asad al Janubiyyah et al.) is not presently IAU-official	
υ	Car A	9 47.7	−65 11	2.92	0.29	A6 II	2.3†	−5.3	~1400†	0.028	307	+14	A: 3.01; B: 5.99, B7 III, 5.1″, PA126°→128°, 1836→2015 orbit ≥ 19,500 y, separation ~2000 AU ¶ the duplicity causes parallax to be poorly known	
φ	Vel A	9 57.7	−54 41	3.52	−0.07	B5 Ib	2.0	−4.9	1600	0.014	285	+14		
η	Leo A	10 08.6	+16 39	3.48	−0.03	A0 Ib†	3	−4.5	1300	~0.003	n.a.	+3 V	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]** B: 8.4, 0.4″, PA 84°→239°, 1937→2015	

¶ mass-loss rate ~5e-8 $M_{\odot}$/y (> 10,000× solar mass-loss rate); BSC5: "chromospheric shell"
¶ we follow WDS in asserting duplicity (with WDS indicating that the A,B pair has been split both through lunar occultation and in speckle interferometry, and documenting 7 observations 1937→2015; the assertion is, admittedly, questioned by Kaler at stars.astro.illinois.edu/sow/etaleo.html)

α	Leo A†	10 09.6	+11 51	1.36	−0.09	B8 IVn†	41	−0.6	79	0.249	271	+6 SB†	**Regulus**

α Leo A is SB orbit 40.11 d, with the secondary in the pair that is α Leo A now detected (2011IBVS.5987....1R reported null photometry result from MOST, at the high precision of ~0.5 millimag., but a spectroscopic detection is reported in 2020ApJ...902...25G; since the secondary is not yet resolved, even interferometrically, WDS is not yet able, at any rate as of 2022 March 3, to write "α Leo Aa" and "α Leo Ab")
¶ the primary in α Leo A is an exceptionally rapid rotator (15.9 h), making the star an oblate spheroid (R_{pol} ~3.14 $R_{\odot}$, but R_{eq} ~4.16 $R_{\odot}$) and rendering the photosphere equator ~3000 K cooler than the photosphere poles (and possibly inducing meridional flow in the envelope);
this is the first rotating star not in an eclipsing binary system to have its gravitational low-latitudes darkening detected, and the first to have its inclination angle and low-latitudes darkening measured through a direct application of spectroscopy-constrained interferometry (inaugural science run of CHARA, 2005ApJ...628..439M);
in contrast with pole-on rapid rotators such as α Lyr A, the α Leo A primary is seen nearly equator-on;
in 2011ApJ...732...68C, Fig. 5 presents an image as fitted to CHARA interferometry (the luminosity contours display the perturbing effect of limb darkening upon the rotation-induced gravity darkening; since the poles are near the limb, the brightest regions, as viewed from Earth, do not quite coincide with the poles);
the aperture-synthesis imaging of 2017NatAs...1..690C Fig. 5 displays the photosphere temperature variation (a joint consequence of limb darkening and oblateness), along with oblateness and axis orientation;
according to 2017NatAs...1..690C, the α Leo A primary (*i*) has attained 96.5% of its breakup speed (earlier literature had suggested 86%), and (*ii*) is the first rapid rotator found to exhibit Chandrasekhar rotation-induced stellar limb polarization (the related phenomenon of eclipse-induced stellar limb polarization was admittedly detected earlier, with β Per, as reported in 1983ApJ...273L..85K)
¶ despite the large rotation-induced latitude variation in photospheric effective temperature, at all latitudes the envelope is radiative (since the photospheric effective temperature, even at the equator, never falls so low as to approach the ~8300 K radiation-to-convection transition value); rotation-induced meridional circulation, on the other hand, disturbs the usual radiative-equilibrium picture of a radiative envelope (2011ApJ...732...68C, p. 11a); since meridional circulation transfers angular momentum, the envelope cannot be presumed to be in solid-body rotation
¶ the rapid rotation, the membership in MK type B, and the near-MS evolutionary status notwithstanding, the question of Be-phenomenon behaviour is answered in the negative by 2005ApJ...628..439M; the authors do, however, remark on p. 446 that the historical record contains a lone report of marginal hydrogen Balmer-α emission, from February 1981

												(might amateur-spectroscopist monitoring now be advisable?) ¶ 2011ApJ...732...68C revises the mass of the primary upward, offering 4.15 $M_{\odot}$ in place of the 2005ApJ...628..439M determination of ~3.5 $M_{\odot}$ ¶ A+BC almost unchanged since 1779 (179″; PA 307°→304°, 1779→2019); AB is nevertheless known to be a true binary pairing, rather than a mere line-of-sight coincidence; AB distance ≥ 4200 AU, orbit ≥ 125,000 y; BC combined light is mag. ~8.2; BC is no longer underobserved (PA: 89°→94°, 4.0″→2.20″, 1867→2019, with orbit ≥ 880 y) ¶ a puzzling discrepancy between the ages of the α Leo SB primary and α Leo B (surely condensed from the same ISM cloud, at the same time) is perhaps to be explained by the peculiarities in the evolution of rapid rotators ¶ we adopt here the MK classification of 2003AJ....126.2048G, while recalling that earlier editions of our RASC brightest-stars table used instead B7 Vn, essentially in accordance with 1953ApJ...117..313J; Astron. Alm. (epoch 2021.5) likewise assigns MK type B7 Vn ¶ the α Leo system is occasionally occulted by Mercury, Venus (e.g. 1959 Jul. 07, 2044 Oct. 01), Moon (e.g. 2017 Sep. 18; 1972JBAA...82..431K describes the 18.6-year 1940-through-2050 cycle of possibilities), and asteroids (e.g. 166 Rhodope 2005 Oct. 19 (2008mgm..conf.2594S reports GTR effect of light bending, not only from general solar gravitational field but also from Rhodope field), 163 Erigone 2014 Mar. 20 (cloud-defeated 2014 Erigone campaign is documented at occultations.org/regulus2014)) ¶ E(B–V) =+0.01	
ω	Car	10 14.3 −70 09	3.29†	−0.07	B8 IIIn†	9.5	−1.8	340	0.037	281	+7 V	or "IIIne"; shell star ¶ rapid rotator (< 1.2 d, ~85% of breakup speed); instance of "Be phenomenon"; photometric variation (cp γ Cas, δ Sco, …) might be expected, and yet seems undocumented; BSC5 does report variable hydrogen Balmer-α	
q	Car A	10 17.8 −61 27	3.39†	1.54	K3 IIa†	5.0	−3.1	660	0.026	286	+8	slight irregular variable: 3.36–3.44 in V ¶ metallicity is uncertain ¶ evolutionary state is uncertain (has core already started He fusion?) ¶ Astron. Alm. (epoch 2021.5) assigns MK type K2.5 II	HR 4050
ζ	Leo A	10 17.9 +23 18	3.43	0.31	F0 IIIa†	12	−1.2	270	0.020	110	−16 SB	B (6.0, 331″ , PA 337°) is mere optical companion (since A,B parallax discrepancy is large, 12 mas for ζ Leo A, but 33 mas for the decidedly less distant ζ Leo B) ¶ Astron. Alm. (epoch 2021.5) assigns MK type F0 III ¶ in rapid evolutionary transition, currently residing in Hertzsprung Gap	Adhafera
λ	UMa	10 18.4 +42 48	3.45	0.03	A1 IV†	24	0.3	140	0.186	256	+18 V	despite MK luminosity class "IV", has not yet finished core hydrogen fusion ¶ mildly metallic, being insufficiently metallic to warrant MK "Am" ¶ seems mild IR excess (indicating circumstellar debris)	Tania Borealis
γ	Leo A +1P†	10 21.2+19 44	2.61†	1.13	K1 IIIb Fe–0.5†	26	−0.3	130	~0.333†~118		−37† SB	4.7″ (2020), PA 99°→127°, 1820→2020 (510.3 y);	Algieba
γ	Leo B	10 21.2+19 44	3.16†	1.42	G7 III Fe–1†	26	0.2	130	~0.346†~118		−36† V	max = ~5″, around 2100 separation ≥ 170 AU, orbital parameters not yet well known ¶ A, B are of mildly unequal masses, and therefore are of mildly disparate evolutionary stage (Kaler	

												stars.astro.illinois.ed: “best understood as being in different stages of gianthood”; cf this same source for further discussion of the uncertainties in various γ Leo parameters, including the respective mags of A and B) ¶ γ Leo A “+ 1P” is an exception to the tendency for exoplanets to be found around the more metallic stars (but the “+1P” could be modelled as a brown dwarf); and indeed even “+2P” is now considered possible ¶ high space velocity of the γ Leo AB pair, plus their low metallicity, suggests system is interloper from more remote galactic region ¶ WDS documents celestial-sphere neighbours γ Leo Ca, Cb, D,E,F, of which all but Ca and Cb are fainter than mag. 10; Ca and Cb, a tight pair first split in 1981, shine with a combined light of mag. 9.64, the contribution coming almost entirely from Ca (the faint Cb has been detected only at 750 nm), and are widely separated from A (341″, at PA 288°, in 2016); the AC pairing is found through analysis of the respective proper motions to be a mere line-of-sight coincidence (211″→341″, PA 294°→288°, 1851→2016) ¶ γ Leo AB, and indeed also the next “Sickle” star ζ Leo, serve to mark the radiant of the Leonids meteor shower	
μ	UMa	10 23.7+41 23	3.06[†]	1.60	M0 IIIp[†]	14	−1.2	230	0.089	293	−21 SB[†]	slight var., 3.03–3.10 in V AAVSO(SVX), viewed 2021 Jan. 16, considers the system to be presenting both eclipse variability resembling the β Lyr case and slow irregular evolved-star variability ¶ SB period 230 d; since the SB is not yet resolved, even interferometrically, WDS cannot yet write “μ UMa A” and “μ UMa B”; ¶ Ca II emission ¶ Astron. Alm. (epoch 2021.5) assigns MK type M0 III ¶ Kaler (stars.astro.illinois.edu-sow-taniaas.html) terms this “a rare ‘hybrid star’” (in the sense of blowing both a fast-and-thin wind and a slower-and-dense wind), and additionally notes the puzzle posed by X-ray emission in the presence of cool photosphere	Tania Australis
p	Car	10 32.8−61 48	3.30v	−0.09	B4 Vne[†]	7	−2.6	500	0.021	304	+26	variable of γ Cas type, 3.22–3.55 in V and instance of the “Be phenomenon” ¶ fast rotator; BSC5: shell; variable hydrogen Balmer-line profiles	HR 4140
θ	Car	10 43.8−64 31	2.74	−0.22	B0.5 Vp	7	−3.0	460	0.022	303	+24 SB[†]	chemically anomalous; of three published orbital solutions for this unresolved SB, the two which seem most reliable (from 1995, 1988) assert periods of 2.20 d, 2.13 d respectively (with *e* values 0.0, 0.24 respectively); the short SB period, even given the low *e*-values (i.e. the lack of severely close periastron passages) suggests mass transfer could be the culprit in the anomalies ¶ since the SB is not as yet resolved, even interferometrically, WDS is not yet able to write “θ Car A,” “θ Car B” ¶ the primary is the brightest of the “blue stragglers”; at stars.astro.illinois.edu/sow/thetacar.html, Kaler discusses difficulties in determination of the primary’s temperature and of its (short) rotation period ¶ E(B–V) =+0.06	
μ	Vel A[†]	10 47.7−49 32	2.82	1.07	G5 III[†]	28	−0.1	~117	0.083	131	+6 SB	A: 2.82; B: 5.65, 2.3″, PA 55°→58°, 1880→2019 period variously given as 116.24 y (Hoffleit) and 138 y (Heintz); A-to-B distance possibly 8 AU min, 93 AU max, 51 AU average ¶ A is itself an SB, not as yet resolved, even in interferometry (so WDS is not as yet able to write “μ Vel Aa”, “μ Vel Ab”) ¶ μ Vel B is of uncertainty-flagged MK type “F8: V” ¶ one or the other component of the SB which is μ Vel A is in rapid evolutionary transition, having recently finished core hydrogen fusion	

												¶ one or other component of the SB which is μ Vel A is magnetic, and an X-ray emitter, with hot corona, and with violent 2-day X-ray flare detected in 1998 by IUE	
ν	Hya	10 50.7−16 19	3.11	1.23	K2 III[†]	23	−0.1	144	0.220[†]	25	−1[†]	slow rotator (but ≤ 619 d) ¶ low metallicity and high space velocity suggest interloper, born outside Sun's neighbourhood ¶ Astron. Alm. (epoch 2021.5) assigns MK type K1.5 IIIb Hδ–0.5	
β	UMa	11 03.2+56 16	2.34	0.03	A0mA1 IV–V	~40.9	0.4	80	0.088	68	−12 SB	debris disk first detected via IR excess, now marginally resolved by *Herschel Space Observatory* (2010A&A...518L.135M)	Merak
α	UMa A[†]	11 05.1+61 38	1.81[†]	1.06	K0 IIIa	27	−1.1	120	0.139	255	−9 SB	A: 1.86; B: 5.0, A8 V, 0.8″ (2017), PA 342° orbit 44 y ¶ the α UMa AB system has also a widely separated, not very faint, celestial-sphere neighbour α UMa C, at mag. 7.19: 384″→370″, PA 205°→205°, 1800→2015 ¶ the first cool star found to have multimodal oscillations (WIRE camera; 2000ApJ...532L.133B suggests fundamental mode 6.35 d) ¶ the most distant of the seven Big Dipper stars (and, like η UMa at the other extreme of the Big Dipper, not a member of the same-age association that is the UMa Moving Group)	**Dubhe**
ψ	UMa	11 10.9+44 23	3.00	1.14	K1 III	22.6	−0.2	145	0.068	246	−4 V?	slow rotator (but ≤ 2.6 y)	
δ	Leo A	11 15.3+20 24	2.56[†]	0.13	A4 IV	56	1.3	58	0.193	132	−20 V	possibly slight (δ Sct type?) var., 2.54−2.57 in V ¶ rapid rotator (< 0.5 d)	Zosma
θ	Leo	11 15.4+15 18	3.33	0.00	A2 IV[†]	~19.8	−0.2	165	0.099	217	+8 V	rotation rather slow for MK type A (but < 9 d); quiet atmosphere renders Ca, Sc underabundant, and Fe, Sr, Ba overabundant; Ca II K-line is variable ¶ IR excess (debris disk?)	Chertan
ν	UMa A	11 19.7+32 58	3.49	1.40	K3 III Ba0.3[†]	~8.2	−1.9	400	0.039	317	−9 SB	B: 10.1, 7.5″, PA 145°→149°, 1827→2018 orbit ≥ 12,000 y; AB distance ≥ 950 AU; in addition to celestial-sphere neighbour ν UMa B and the very faint celestial-sphere neighbour ν UMa C, WDS lists the not-very-faint, but widely separated, celestial-sphere neighbour ν UMa D (mag. 8.88; angular separation from A was 287″ in 2015, at PA 267°) ¶ Astron. Alm. (epoch 2021.5) assigns MK type K3⁻ III **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	Alula Borealis
ξ	Hya Aa	11 34.1−31 59	3.54	0.95	G7 III	~25.2	0.5	130	0.214	259	−5 V	we give mag.value for ξ Hya Aa, Ab combined light; ξ Hya Aa is of mag. 3.7, and ξ Hya Ab of mag. 5.8; WDS documents just one astrometry measurement for this tight pair, from 2001 (angular separation 0.1″) ¶ the CORALIE spectrograph on the 1.2 m Euler telescope at La Silla (ESO) in 2002 demonstrated for the first time the feasibility of asteroseismology for a highly evolved star: ξ Hya Aa has left the theoretical MS, being now near the theoretical SGB-RGB transition, and yet CORALIE was able to find solar-like oscillations (with periods, however, of 2.0 h to 5.5 h, in contrast with the "five-minute oscillations" in the Sun; ξ Hya Aa, being evolved, is larger than the Sun, and so its starquakes face larger propagation distances; on the other hand, where the solar "five-minute oscillations" involve speeds of ~15 cm/s or ~20 cm/s, in the case of ξ Hya Aa the observed speeds attain values only a little below 2 m/s (such a refined stellar radial-velocity measurement is now possible as a kind of spinoff from advances in exoplanet-search spectrograph engineering); both radial and non-radial oscillations are found in ξ Hya Aa; the pertinent ESO press release is at www.eso.org/public/news/eso0215/, and the modern ξ Hya Aa asteroseismology literature	

												in the journals starts with the "Letter" which is 2002A&A...394L...5F **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**
λ	Cen Aa†	11 36.8−63 09	3.11	−0.04	B9.5 IIn†	8	−2.4	400	0.034	258	−1 V	despite possible fast rotation (< 2.7 d?), Fe is overabundant, with Si and C mildly underabundant ¶ at stars.astro.illinois.edu/sow/lambdacen.html, Kaler discusses questions of visual binarity (λ Cen Aa, Ab, B: but WDS documents only a single observation of λ Cen Ab, from the year 2000, and just 3 observations of the faint λ Cen B, over the period 1937→2015)
β	Leo A	11 50.2+14 27	2.14†	0.09	A3 Va	91	1.9	36	0.511	257	0 V	**Denebola** very slight var., range 0.025 in V; AAVSO(VSX) classifies as variable type "DSCTC" (low-amplitude δ Sct), but cautions that this type is now considered obsolete; the assertion of δ Sct variability in any case now may be erroneous ¶ rapid rotator (< 0.65 d) ¶ debris disk resolved by *Herschel Space Observatory* (2010A&A...518L.135M), disk structures differentiated with ground-based interferometry (2010ApJ...724.1238S) ¶ WDS documents very faint celestial-sphere neighbours β Leo B,C, and additionally the not-so-faint D (mag. 8.5; AD astrometry is 298″→240″, PA 204°→190°, 1833→2019) and the tight (0.5″) pair Ea,Eb (mags. 6.5 and 6.6; this pair is separated by 2″ from A; but for the "AE" pairing WDS has just a single measurement, from 2009, and likewise for the Ea,Eb pairing WDS has just a single measurement, again from 2009)
γ	UMa A	11 55.0+53 34	2.41	0.04	A0 Van†	39	0.4	83	0.108	84	−13	Phecda rapid rotator: although in MK temperature class A, nevertheless an instance of the "Be phenomenon" (the term "Ae star" is sometimes used for this rare category) ¶ γ UMa A with an invisible companion (or is this, rather, the sparsely observed pairing of γ UMa A with (visible, mag. 8.2 γ UMa B?) is said in en.wikipedia.org/wiki/Gamma_Ursae_Majoris to be an astrometric binary, of period 20.5 y (and we are also not fully certain on the SB situation: although we have in the past applied the flag "SB" (is γ UMa A a binary with companion unresolved?), we now withdraw the flag, on the strength of 2010NewA...15..324G (which proposes an astrometric binary-system orbit for γ UMa AB, but additionally writes, "Spectroscopic duplicity of this star mentioned in some catalogues seems to be a mistake: it [sc the alleged duplicity] could not have been detected because of a large rotational velocity [with rotational broadening, therefore, of the spectrum lines]")) ¶ E(B–V) =0.00
δ	Cen Aa†	12 09.5−50 51	2.58v	−0.13	B2 IVne†	8	−2.9	400	0.050	262	+11 V	var. of γ Cas type, 2.51–2.65 in V ¶ rapid rotator (< 1.3 d), with shell spectrum; 2008A&A...488L..67M summarizes recent research, and as part of a wider VLTI investigation into the "Be phenomenon" not only discusses the circumstellar ejecta, but also reports discovery of binarity (Ab at angular distance 68.7 mas); WDS also documents celestial-sphere neighbours δ Cen B, δ Cen C; δ Cen B (a Be star) is mag. 4.4, with astrometry 221″→269″, PA 325° →325°, 1847→1999; δ Cen C is mag. 6.4, with astrometry 218″→217″, PA 227°→227°, 1847→1999
ε	Crv	12 11.3−22 45	3.02	1.33	K2 III†	~10.3	−1.9	320	0.072	278	+5	slow rotator (but ≤ 3.9 y) ¶ metals somewhat overabundant ¶ evolutionary status uncertain (core-helium fusion starting, in progress, or finished?)

												¶ Astron. Alm. (epoch 2021.5) assigns MK type K2.5 IIIa ¶ the etymologically Arab name "Minkar" is of merely modern origin, and is not currently IAU-official	
δ	Cru	12 16.4−58 52	2.79†	−0.19	B2 IV†	9.4	−2.3	350	0.037	254	+22 V?	slight var., β Cep type, 2.78–2.84 in V, 0.15 d ¶ rapid rotator (<1.3 d; BSC5: "expanding circumstellar shell")	Imai
δ	UMa A	12 16.5+56 54	3.32	0.08	A2 Van	40.5	1.4	81	0.104	86	−13 V	possesses debris disk, of unusually low radius (Wyatt et al 2007; Pointing-Robertson drag?)	Megrez
γ	Crv	12 17.0−17 40	2.58†	−0.11	B8 III†	21	−0.8	154	0.160	278	−4 SB	slight var., 2.56−2.60 in V, of unknown type ¶ rather rapid rotation notwithstanding (BSC5: "expanding circumstellar shell"), Hg and Mn are overabundant, with some other elements underabundant (but rotational line broadening makes abundance determinations difficult); Astron. Alm. (epoch 2021.5) assigns MK temperature type B8p Hg Mn, and does not assign an MK luminosity class	**Gienah**
α	Cru A†	12 27.9−63 13	1.25†	−0.20	B0.5 IV	10	−3.7	~320	0.037	251	−11 SB	5.4″ (1826); 3.5″ (2020)	**Acrux**
α	Cru B†	12 27.9−63 13	1.64†	−0.18	B1 Vn	10	−3.3	~320	0.037?	251?	−1	PA 114°→111°, 1826→2020 orbit ≥ 1300 y, AB distance ~430 AU (the WDS "U" flag for AB, indicating that the pairing is a mere line-of-sight coincidence, does not seem supported in the most obvious part of the secondary literature (Kaler, Wikipedia)); A is SB pair (not as yet resolved, even in interferometry, so WDS cannot as yet write "α Cru Aa", "α Cru Ab"), with period 75.78 d, distance between components ~0.5 AU min, ~1.5 AU max; C (itself an SB pair, unresolved), at mag 4.8, has been said to be imperfectly sharing AB proper motion (AC astrometry 92″→89″, PA 216°→203°, 1750→2020), and so is possibly (not assuredly) gravitationally bound with the putatively gravitationally bound three-star AB system (if C is bound, then period > 130,000 y, with distance from AB ≥ 9,000 AU); the other celestial-sphere neighbours of the unresolved α Cru A SB pair are fainter than mag. 10; the IAU-official name "Acrux" applies to the "primary" (the more luminous, and more massive) component of the unresolved α Cru A pair; that unresolved pair has also been called α^1 Cru, with α Cru B correspondingly called α^2 Cru; it is a little puzzling that the WDS note for this α Cru ABC putatively 5-star system is silent on SB status of α Cru C, while suggesting instead (on strength of 1967 Batten work at DAO) possible SB status for α Cru B ¶ the AB duplicity makes the two individual mag. determinations "Aa-with-Ab-combined light", "B" somewhat controverted	
δ	Crv A†	12 31.0−16 38	2.94	−0.01	B9.5 IVn	~37.6	0.8	87	0.252	237	+9 V	B:8.3, K2 V, 24″, PA 216°→216°, 1782→2020 although A,B have common proper motion, disparity in age estimates has caused binarity to be questioned; Kaler, accepting binarity (he proposes period ≥ 9400 y) suggests that the δ Crv AB system is young, and that B (radiating less powerfully than A) is a post-T-Tauri star (i.e. a star which, although already stably burning core hydrogen, has nevertheless not yet succeeded in clearing away its surrounding dust)	Algorab
γ	Cru A	12 32.4−57 14	1.59†	1.60	M3.5 III†	37	−0.6	89	0.267	174	+21	slight semiregular var., 1.60–1.67 in V although has been classified as semiregular var., at least 6 pulsation periods have been documented ¶ the nearest of the M giants, radius > 0.5 AU; evolutionary status uncertain (is core He fusion now finished?) ¶ cause of the observed Ba overabundance is unknown (undetected evolved companion?) ¶ γ Cru B (celestial-sphere neighbour through mere line-of-sight coincidence) is mag. 6.4, with AB astrometry 93″→133″, PA 41°→24°, 1826→2018; γ Cru C is mag. 9.7, with AC astrometry 155″→168″, PA 86°→68°, 1879→2018	**Gacrux**
β	Crv	12 35.6−23 31	2.65†	0.89	G5 II†	22	−0.6	146	0.057	179	−8 V	slight var. (2.60−2.66 in V) reported or suspected	Kraz

												¶ slow rotator (but ≤ 180 d) ¶ possibly in evolutionary transition (He core about to ignite?) ¶ assertion of weak Ba-star status is perhaps erroneous ¶ Astron. Alm. (epoch 2021.5) assigns MK type G5 IIb	
α	Mus Aa	12 38.6 −69 16	2.69†	−0.18	B2 IV–V	10.3	−2.2	320	0.042	252	+13 V	slight var., β Cep type?, 2.68–2.73 in V, 0.090 d classification of the low-amplitude variability as β Cep has been questioned ¶ rapid rotator (< 2 d) ¶ we give mag. for combined light of α Mus Aa and its very close, less luminous, companion α Mus Ab: the pair is sparsely observed, perhaps with nothing beyond the 2013MNRAS.436.1694R report that binarity has been discovered (Sydney interferometer: two 2020 observations, with angular separations ~10 mas, ~16 mas); individual mags. are 2.8, 5.5	
γ	Cen A†	12 42.8 −49 05	2.95	−0.02	A1 IV	25	−0.1	130	~0.194	~267	−6	orbit 84 y; 0.4″ (2010),	
γ	Cen B †	12 42.8 −49 05	2.85	−0.02	A0 IV	25	−0.2	130	~0.194	~267	−6	0.3″ (2019); max = 1.7″; AB distance 8 AU min, 67 AU max, 37 AU average; γ Cen D (mag. 3.85), despite its large physical distance from, and large angular separation from, the γ Cen AB pair (1.72 ly, ~1°) is likely gravitationally bound to γ Cen AB, experiencing that binary as essentially a point mass; γ Cen D is a.k.a. τ Cen ¶ Arabic name Muhlifain, for γ Cen A, is not currently IAU-official	
γ	Vir AB†	12 42.8 −1 34	2.74	0.37	F1 V + F0mF2 V	85	2.4	39	~0.619	~276	−20	A: 3.48; B: 3.53; 0.8″ (2007); 3.0″ (2020) orbit 169 y; separation 5 AU min (most recently 1836 and 2005), 81 AU max, 43 AU average, with plane of orbit inclined 31° to plane of sky; for discussion of orbit, with observations plot showing error bars (binary astrometry being now old enough to archive data for one full orbit), cf Kaler at stars.astro.illinois.edu/sow/porrima.html; the γ Vir AB pair has, in addition to some faint celestial-sphere neighbours, the widely separated, but not very faint, neighbours γ Vir E (mag. 8.9) and γ Vir F (mag. 9.5), neither of which is gravitationally bound to the AB pair; AE astrometry is 255″→260″, PA 186°→167°, 1851→2016, and AF astrometry is 482″→422″, PA 269°→267°, 1909→2015 ¶ lunar occultations possible, planetary occultations possible-yet-rare	Porrima
β	Mus Aa	12 47.7 −68 14	3.04	−0.18	B2 V†	~9.6	−2.1	340	~0.043†	~258	+42† V	A: 3.52; B: 3.98, 1.0″, PA 317°→58°, 1880→2019 orbit 194 y; average AB distance uncertain (101 AU, or only ~80 AU?); orbit map, showing error bars, given by Kaler at stars.astro.illinois.edu/sow/betamus.html, with Kaler's accompanying discussion of orbit-modelling problems, underscores limitations in current β Mus AB knowledge; 2013MNRAS.436.1694R reports splitting β Mus A as tight binary system β Mus Aa, β Mus Ab, in two 2010 observations with Sydney interferometer: Ab is mag. 6.6, and Aa,Ab angular separation is < 50 mas ¶ β Mus A is rapid rotator (< 1 d) ¶ β Mus B is of MK type B2.5 V ¶ a runaway system, in the sense of presenting a high velocity relative to the general galactic rotation	
β	Cru A	12 49.1 −59 49	1.25†	−0.24	B0.5 III†	12	−3.4	300	0.046	249	+16 SB†	slight var., β Cep type, 1.23–1.31 in V, 0.24 d SB period 1828.0 d, *e*=0.38 (with distance between the SB components 5.4 AU min, 12.0 AU max); Kaler at stars.astro.illinois.edu/sow/mimosa.html discusses other possible companions, including an X-ray visible, and yet optically invisible, object interpreted as a pre-MS star; in gross optical-astronomy terms, the unresolved β Cru A SB has celestial-sphere	Mimosa

											neighbours B,C,D,E, of which only C is brighter than mag. 10; AC astrometry is 384″→373″, PA 23°→23°, 1826→2000 ¶ β Cru A is believed to be a rapid rotator (possible ~3.6 d) ¶ β Cru A is a multiperiodic β Cep variable ¶ β Cru A, its MK luminosity class "III" notwithstanding, is only about halfway through its career of stable-core hydrogen fusing	
ε	UMa A	12 55.0+55 50	1.76†	−0.02	A0p IV: (CrEu)	~39.5	−0.3	83	0.112	94	−9 SB?	slight var., α^2 CVn type, 1.76–1.78 in V, 5.1 d the brightest of the Ap stars (in the specific case of ε UMa A, the magnetic-dipole axis is believed to be nearly perpendicular to rotation axis, yielding Cr bands nearly perpendicular to equator; dipole strength is unusually low) (but it has also been suggested that a substellar companion of mass ~14.7× Jupiter, at average inter-component distance 0.055 AU, orbit 5.1 d, rather than a 5.1-d stellar rotation, is the source of the observed variability period); WDS, as viewed 2021 Nov. 17, documents just 2 astrometry measurements for the ε UMa B companion of ε UMa A; the discovery of ε UMa B, via speckle interferometry, is announced in 1978MNRAS.183..701M; we at the Handbook do not know whether this discovery of ε UMa B constitutes a resolving of the putative SB	**Alioth**
δ	Vir A	12 56.7 +3 17	3.39†	1.57	M3 III†	16	−0.5	~198	0.473†	264	−18† V?	slight var., semiregular, 3.32−3.40 in V multiperiod pulsator ¶ high space velocity relative to galactic neighbours ¶ evolutionary status uncertain (helium fusion recently started, or already finished?) ¶ Astron. Alm. (epoch 2021.5) assigns MK type M3$^+$ III	Minelauva
α	CVn A†	12 57.1+38 12	2.85v†	−0.06	A0 Vp (SiEu)	28	0.1	110	0.241	283	−3 V SB	B:5.5, F0 V, 19.5″, PA 234°→230°, 1777→2018 orbit ≥ 8300 y (common proper motion indicates true binarity); separation ≥ 675 AU; prototype for the α^2 CVn var. class (chemically anomalous photospheric regions yielding spectroscopic variability, and with magnetism yielding large spots; in the particular case of α^2 CVn, rotation period is 5.46939 d, with consequent spot-driven mag. range 2.84–2.98); the α CVn A SB is not as yet resolved, even in interferometry (so WDS cannot as yet write "α CVn Aa", "α CVn Ab"), and also α CVn B is an as-yet-unresolved SB (so WDS cannot as yet write "α CVn Ba", "α CVn Bb") ¶ two correct, potentially confusing, designations are α CVn A (signalling that this is the brighter of the binary pair) and α^2 CVn (signalling that α^1 crosses the local meridian before α^2, lying further W); the Latin "heart-of-Charles" designation, official at IAU as of 2016, honours the "martyr king" Charles I (although Charles II is sometimes cited in error)	Cor Caroli
ε	Vir A	13 03.3+10 50	2.85	0.93	G9 IIIab†	29.8	0.2	110	0.275	274	−14 V?	one of the most notable X-ray sources in our table (X-ray luminosity, although far below α Aur, is nevertheless almost 300× solar) ¶ Astron. Alm. (epoch 2021.5) assigns MK type G8 IIIab	Vindemiatrix
γ	Hya A	13 20.2−23 17	2.99	0.92	G8 IIIa	~24.4	−0.1	134	0.081	121	−5 V?	slow rotator (but ≤ 240 d) evolutionary state uncertain (core-helium fusion impending, or already in progress?)	
ι	Cen	13 21.9−36 50	2.75	0.07	A2 Va†	55	1.5	59	0.352	256	0	rapid rotator (< 2d) ¶ low metallicity ¶ debris disk (unusually luminous, given evolutionary state of ι Cen)	
ζ	UMa Aa†	13 24.8+54 48	2.23	0.06	A1 Va†	40	0.1	90	0.123	100	−6 SB2†	B:3.88, A1mA7 IV–V, 15″; period >5000 y? (more precisely 13.9″→14.6″, PA 143°→153°, 1755→2020); not only is ζ UMa AB a true binary; it is now additionally argued (controversy possibly continues) that the pair ζ UMa Ca (Alcor), Cb is gravitationally bound to the pair AB (Bob King, *Sky and Telescope* 2015 March 25); ζ UMa Aa, ζ UMa Ab are an interferometrically resolved SB2, seen nearly edge-on, with at	Mizar

														least three orbital solutions published (20.5385 d or 20.5386 d, *e*=5.54 or 5.53, angle subtended by semimajor axis as projected onto plane of sky ≈10 mas (apod.nasa.gov/apod/ap970219.html, as NASA “Astronomy Picture of the Day” for 1997 Feb. 19, depicts an NPOI-derived ζ UMa Aa,Ab orbit), with 2010NewA...15..324G suggesting, on the basis of astrometry perturbations, a possible further unseen body); ζ UMa B is an unresolved SB, period 175.6 d, with highly elliptical orbit; although the old, widely repeated claim (cf Heard *ApJ* 1949) that ζ UMa C is binary is shown in www.leosondra.cz/en/mizar/ to be unfounded, binarity is now established (with WDS accordingly writing “ζ UMa Ca”, “ζ UMa Cb”, at mags. 4.0 and 8.0 respectively, typical angular separation 1″, with 5 satisfactory astrometry measurements over the period 2007→2009); Cb is a mid-M red dwarf, very notable as one of the few cases of a red dwarf detected as gravitationally bound to an A star (Ca is of MK type A5 Vn); the IAU-official name “Alcor” applies to ζ UMa Ca; www.leosondra.cz/en/mizar/ should be consulted also *(a)* for details on ζ UMa multiplicity-studies history, including Galileo and Michelson (Leos Ondra, citing inter alia Fedele 1949, seems to establish that it was Galileo’s pupil Castelli, rather than (as widely asserted) Riccioli, who discovered Mizar’s visual duplicity) and *(b)* for a 15′ map documenting around 20 of the stars in the field, including mag. 7.6 “Stella Ludoviciana” (“Sidus Ludovicianum”, in WDS ζ UMa D), a mere line-of-sight coincidence on the celestial sphere, too distant from ζ UMa ABC to be gravitationally bound to this 6-star system (Aa,Ab; B binary-as-yet-unresolved, Ca,Cb): WDS additionally documents, as gravitationally bound to the ζ UMa ABC 6-star system, E (mag. 6.9), F (mag. 9.9), G (mag. 8.2), and H (mag. 8.6) ¶ Astron. Alm. (epoch 2021.5) assigns MK type “A1 Va$^+$ (Si)” ¶ in the mythology of the Mi’kmaq and the St Lawrence Seaway Iroquois, as presented at www.aavso.org/myths-uma, α, β, γ, and δ UMa are a bear at various seasons of the year passant, rampant, and expired (its four paws upward in death), pursued in the warm months by seven hunters, but once the nights are cold by a remaining above-horizon three, these persistent three being ε UMa, Mizar, and ζ UMa, with Alcor the middle hunter’s cooking pot, awaiting bearmeat; www.aavso.org/myths-uma gives further detail, offering also a speculation about a possible bear mythology shared by Siberian and North American Paleolithic peoples, in the epoch of the Bering Strait land bridge	
α	Vir Aa†	13 26.4	−11 17	0.98†	−0.24	B1 V†	13	−3.4	250	0.052	234	+1	SB2†	slight var., 0.96–1.00 in V, 4.0145 d very tight (< 1″) system resolved interferometrically (and in occultation?) into α Vir Aa (mag. 1.3), α Vir Ab (mag. 4.5), α Vir Ac (mag. 7.5); as SB2 (primary-to-secondary distance 0.12 AU; the geometry is close to achieving a grazing eclipse), the brightest of the rotating ellipsoid variables (by definition no eclipse, and by definition with the SB’s total presented luminous area varying, through geometrical asymmetry, as the orbital motion proceeds); the Aa,Ab orbit is highly eccentric; Aa was measured in 1975 to lie 0.50″ from Ac ; although the Aa,Ab pair is at all times very close, an angular distance of 0. 1″ is reported from 1975;	**Spica**

Aa (a rapid rotator, at ~0.3 breakup speed) is itself
a pulsating variable of the β Cep type (0.1738 d;
shortly after the ~1970 discovery
of the β Cep variability,
photometric and spectroscopic
variations were present;
the photospheric variations soon ceased,
but the spectroscopic
(radial-velocity, i.e. pulsational)
variations continued;
2016MNRAS.458.1964T, incorporating
precision MOST photometry, reports
for Aa one radial and two
non-radial pulsation modes,
with one of the non-radial modes tidally induced)
¶ in an early application of intensity interferometry,
1971MNRAS.151..161H argues with
the example of α Vir Aa,Ab
that given supporting spectroscopy and photometry,
orbit and distance of a
double-lined SB can be deduced
(the SB distance notably
without recourse to trigonometric parallax,
since distance can be deduced by comparing
the angular and the physical dimensions of the
ascertained orbit)
¶ the tidal-interactions studies
2016A&A...590A..54H and
2013A&A...556A..49P stress
the importance of the α Vir Aa,Ab
double-lined SB for critically testing the
(astrophysically foundational) assumption
that the individual components *x, y* of a binary,
of determined masses, rotation
periods, and chemical compositions,
resemble in their photospheres,
and even in their interiors,
solitary stars *x', y'* possessing
the same masses, rotation periods,
and chemical compositions
(could tidal effects, e.g. change
internal temperature structure?);
additionally, the tidal effects in the α Vir Aa,Ab SB
are judged in 2009ApJ...704..813H to be
responsible for large-scale shearing
horizontal photospheric motions,
spectroscopically observable as
modifiers of line profiles
(but 2016MNRAS.458.1964T questions
the judgement)
¶ assignment of individual MK types
to Aa, Ab is challenging:
the rather-unevolved-B MK types
(1971MNRAS.151..161H B1.5 IV-V + B3V,
2007AAS...211.6301A B0.5 III-IV +
B2.5-B3V) are in any case
consistent with rather high masses
(10.9 $M_\odot$ + 6.8 $M_\odot$,
10.25 $M_\odot$ + 6.97 $M_\odot$,
for these two respective papers)
¶ as is to be expected from
the failure of Aa,Ab to be tidally
locked, the system is young
(with 2016MNRAS.458.1964T
assigning as age 12.5 ± 1 My)
¶ the Aa,Ab binary is a polarimetric
variable (ISM material
entrained?), and a strong X-ray
source (colliding winds?)
¶ α Vir Ab is one of the few stars known
to exhibit Struve-Sahade variation
(en.wikipedia.org/wiki/Struve%E2%80%93Sahade_effect)
in its spectral line strengths

												¶ 1972JBAA...82..431K describes the 18.6-year 1940-through-2050 cycle of lunar occultation possibilities ¶ E(B-V)=+0.03	
ζ	Vir A	13 35.8 −0 43	3.38	0.11	A2 IV†	44	1.6	74	0.285	280	−13	a good marker of celestial equator (precession placed ζ Vir exactly onto equator in 1883 Feb.) ¶ rapid rotator (< 0.5 d; this renders puzzling the possible evidence for chemical anomalies, which would presuppose a quiet atmosphere) ¶ Astron. Alm. (epoch 2021.5) assigns MK type A2 IV$^-$ ¶ elusive red-dwarf companion ζ Vir B is mag. 10.0, with MK classification suggested as M4 V – M7 V; WDS documents just 9 ζ Vir AB measurements, 1.8″→1.8″, PA 145°→154°, 2004→2010; the discovery paper is 2010ApJ...712..421H, establishing inter alia shared proper motion, and proceeding from stellar coronography on adaptive-optics platforms, both at Palomar and at Hawai'i-Haleakala (rather than, as is more usual with a difficult binary, from interferometry (with ζ Vir B directly imaged in Figs. 1 and 4)); 2010ApJ...712..421H, while remarking that orbital coverage is as yet too brief for an orbital solution to be attempted, nevertheless (assuming 2.04 $M_\odot$ for ζ Vir A, 0.168 $M_\odot$ for ζ Vir B) computes approximate lower bounds for semimajor axis, for *e*, and for period as, respectively, 24.9 AU, 0.16, and 124 y; 2010ApJ...712..421H is a contribution to the important, and until recently unstudied, topic of low-mass companions for A stars (another contribution to this topic is, however, the circa-2010 discovery of an elusive M-type companion for the A5 Vn star which is Alcor); the topic in its turn is a building block in the overall theory of star and exoplanet genesis (could massive (e.g. A-type) stars acquire low-mass companion stars via condensation not directly from the parent molecular cloud, but rather from condensation in an unstable circumstellar disk?), the 2010ApJ...712..421H detection of red-dwarf ζ Vir B additionally affords an explanation for the puzzling ζ Vir X-ray emission observed by *ROSAT* (as a star of a spectral type lacking strong winds and lacking convection up to the photosphere, ζ Vir A would not itself be expected to emit X-rays: similar puzzles of X-ray emission from the putatively X-ray-dark A stars arise elsewhere also, and perhaps are similarly to be solved in terms of X-ray emission from (elusive) red-dwarf companions)	Heze
ε	Cen Aa	13 41.3−53 35	2.29†	−0.17	B1 III†	8	−3.3	400	0.019	233	+3	slight var., β Cep type, 2.29−2.31 in V multiperiodic, with primary period 0.169608 d ¶ rapid rotator (< 2.7 d) ¶ metals underabundant ¶ although we here assign MK luminosity class "III", Kaler at stars.astro.illinois.edu/sow/epscen.html discusses uncertainty ¶ we state mag. for combined light; WDS documents just one single measurement, from interferometry, for the elusive ε Cen Ab (mag 4.90, 0.2″ from Aa)	
η	UMa	13 48.4+49 12	1.85	−0.10†	B3 V†	31	−0.7	104	0.122	263	−11 SB?	resembles α UMa, at the other extreme of the Big Dipper, in not belonging to UMa Moving Group; 1921LicOB..10..110T asserts membership in what was at that time called the "Pleiades Group" ¶ rapid rotator (< 21 h), with some line variability (circumstellar ejecta disk?) ¶ X-ray source ¶ colour and mean temperature are anomalous for the MK type ¶ unusually young in our Sample S (< 15 My) ¶ E(B–V)=+0.02	**Alkaid**
ν	Cen	13 50.9−41 48	3.41†	−0.22	B2 IV†	~7.5	−2.2	440	0.034	233	+9 SB†	SB period is 2.622 d; system is rotating ellipsoidal slight variable (not eclipsing, but varying in light as the presented surface area changes); additionally, the primary is a pulsator in the β Cep class (observed mag. range from these two causes is 3.40–3.42 in V) ¶ MK luminosity class "IV" notwithstanding, primary is still a stable fuser of core hydrogen ¶ possible weak instance of	

μ	Cen Aa	13 51.0−42 35	3.47v†	−0.17	B2 IV–V pne†	~6.4	−2.5	510	0.031	232	+9 SB	“Be phenomenon” (with the outbursts possibly temporary) variable of γ Cas type,: 2.92–3.47 in V rapid rotator, and (consistently with the γ Cas behaviour) an instance of the “Be phenomenon”; additionally said to be a multiperiodic non-radial pulsator; BSC5: “line profiles of MgII 4481 change in period 0.505 d, about five times the period of weaker absorption”; variable Hα; “variable line profiles”; short-term photometric and polarimetric variability has also been reported (cf p. 46 of 2013A&ARv..21…69R, which notes a rapid rise, over just a few days, in photometric brightness or line-emission intensity, with a subsequent slower decline) ¶ we give mag. for combined light: WDS documents just one single measurement for μ Cen Aa,Ab, via 2010 Sydney interferometry (as Aa mag. 3.50, Ab mag. 6.70, with angular separation 0.1″) ¶ Astron. Alm. (epoch 2021.5) assigns MK type “B2 IV–Vpne (shell)”
η	Boo A	13 55.8+18 17	2.68	0.58	G0 IV†	88	2.4	37	0.361	190	0 SB	**Muphrid** unusually metal-rich ¶ an X-ray source (hot corona) ¶ 2007ApJ…657.1058V discusses recent work (MOST helioseismology, PTI interferometry) ¶ η Boo B is mag. 9.99; discrepant proper motions for the AB pair (126″→115″, PA 119°→85°, 1822→2020) establish that their pairing is a mere line-of-sight coincidence
ζ	Cen	13 57.0−47 24	2.55	−0.18	B2.5 IV†	8.5	−2.8	380	0.073	232	+7 SB2	SB period 8.02 d; SB as yet unresolved, even by interferometry (so WDS not yet able to write “Cen A,” “Cen B”) ¶ primary is a rapid rotator (<1.5 d) (BSC5: “expanding circumstellar disk”, and yet not (as viewed 2022 March 3) catalogued as an instance of the “Be phenomenon” in Paris-Meudon BeSS database) ¶ MK luminosity class “IV” notwithstanding, primary is possibly only halfway through its core hydrogen fusing ¶ E(B–V) =–0.02
β	Cen Aa,Ab	14 05.4−60 29	0.58†	−0.23	B1 III + B1 III	9.0†	−4.8	360†	0.041	235	+6 SB2†	**Hadar** B:3.94, A1mA7 IV–V, 0.3″ (2019) (more fully: 1.1″→0.3″, PA 257°→168°, 1935→2019); AB orbit is already constrained by the existing observations, with period 125−220 y, and 2016A&A...588A..55P indicates that it should be possible to compute the orbit by ~2025 or ~2030 or so); the β Cen A system (reported in 1999MNRAS.302..245R as resolved at AAT through spectrally dispersed aperture-masking interferometry) is β Cen Aa, β Cen Ab, comprising a pair of fast rotators of nearly equal mass, and with oddly disparate (high) rotation speeds (the slower rotator is known to be magnetic, so magnetic braking is possible) and with orbit so eccentric (how can the molecular-cloud ISM condensation have allowed this to happen?) as to make the periastron tight (at < 10 R*; so could there be tidal interaction between Aa, Ab at periastron, perturbing the variability that we discuss below?); Aa,Ab is additionally reported in 1999MNRAS.302..245R, on the strength of ESO spectra, to be not just SB, but SB2; since the masses are nearly equal, it becomes a delicate question which to take as the primary, i.e. to which to apply the label “Aa” and with it the IAU-official name “Hadar”; this question is answered by WDS in its usual terms, with “Aa” deemed to be the (very slightly) more luminous, more massive star (Aa mag. 1.29, Ab mag. 1.44), and yet the contrary decision has also been taken in the literature, since it is the less massive star that has the clearer, because the less severely (rotationally) broadened, spectrum; the observational challenges notwithstanding, the observational record for Aa,Ab is favourable, with WDS now documenting 53 measurements for 1995→2018

												(not only aperture-masking interferometry, but also speckle interferometry has been done); the Aa,Ab orbit is 357 d, with *e*=0.8; 2002A&A...384..209A finds both β Cen Aa and β Cen Ab to be β Cet variables; the Aa,Ab system is of astrophysical significance, as one of the rare cases of β Cep variability amenable to good-precision mass studies (and indeed 2016A&A...588A..55P, reporting precision two-filter photometry with the BRITE constellation, discusses prospects for future asteroseismology, noting also that in addition to β Cep ("pressure-wave") pulsation, there is SPB-type ("gravity-wave", i.e. buoyancy-driven) pulsation (the question which of Aa,Ab is presenting which of the 17 detected pulsation modes is, however, difficult, and the magnetic field of Ab is a further complication (with the overall topic of asteroseismology for magnetic β Cep variables only sparsely explored, at any rate as of ~2016)) ¶ we take π and D not from *HIPPARCOS* but from 2016A&A...588A..55P ¶ E(B–V) =+0.02 ¶ the traditional (Latin-derived?) name "Agena" is not IAU-official
π	Hya	14 07.7−26 47	3.25	1.09	K2 IIIb[†]	~32.3[†]	0.8	~101[†]	0.148[†]	163	+27[†] V	negative cyanide ion lines are anomalously weak relative to metal lines, consistent with this star's anomalously high velocity relative to Sun (suggesting interloper in our own galactic region; however, π Hya is more metal-rich than the celebrated interloper α Boo (Arcturus)) ¶ Astron. Alm. (epoch 2021.5) assigns MK type K2^{-} III Fe–0.5 ¶ in evolutionary terms, in "Red Clump" of core-He fusers (but uncertain whether recent arrival in clump or longtime denizen)
θ	Cen A	14 08.0−36 29	2.06	1.01	K0 IIIb[†]	55	0.8	59	0.734[†]	225	+1[†]	**Menkent** high velocity with respect to Sun suggests interloper status (and yet metallicity is approximately solar) ¶ Astron. Alm. (epoch 2021.5) assigns MK type K0^{-} IIIb
α	Boo A	14 16.7+19 04	−0.05	1.24	K1.5 III Fe–0.5[†]	89	−0.3	37	2.279[†]	209	−5[†] V	**Arcturus** high space velocity a metal-poor interloper (from galactic thick disk? but galaxy-merger scenario has also been suggested), and member of Arcturus Moving Group (2009IAUS..254..139W) ¶ a magnetic cycle (< 14 y?) has been detected ¶ still ascending RGB, with He flash impending? (but a later evolutionary stage has also been suggested) ¶ publication of α Boo A line atlas 1968pmas.book.....G (R.Griffin) was a major event in postwar spectroscopy ¶ α Boo A has been studied in recent asteroseismology ¶ a single 1991 observation of a putative companion has constrained WDS to write "α Boo A", "α Boo B"; however, 1990s assertion of multiplicity was retracted in 1998; independently of this pair of developments, there have been suggestions of a sub-stellar-mass companion at the margin of *HIPPARCOS* detectability
ι	Lup	14 20.9−46 10	3.55	−0.18	B2.5 IVn[†]	~9.6	−1.5	340	0.013	249	+22	rapid rotator (possibly ~0.9 d), and yet no evidence of circumstellar disk, and in particular no Be-phenomenon spectral features ¶ the MK luminosity class "IV" notwithstanding, still performing stable core-hydrogen fusion
γ	Boo Aa[†]	14 33.0+38 13	3.04[†]	0.19	A7 IV^{+}	37.6	0.9	87	0.190	323	−37 V	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]** slight var. of δ Sct type, 3.02−3.07 in V, 0.2903 dSeginus ¶ IR excess (from circumstellar debris, so far unexplained) ¶ Aa,Ab resolved in speckle interferometry, angular separation 70 mas

η	Cen	14 36.9−42 15	2.33v†	−0.16	B1.5 IV pne†	11	−2.5	310	0.048	227	0 SB	var. of γ Cas and λ Eri types, possibly shell, 2.29−2.47 in V multiperiodic, with primary period 0.6425 d; BSC5: Hα variable, Hβ “sometimes bright, sometimes dark and double or multiple”; consistently with γ Cas variability, a rapid rotator (< 1 d) and an instance of “Be phenomenon”; consistently with γ Cas variability, Astron. Alm. (epoch 2021.5) assigns MK type “B1.5 IVpne (shell)”	
α	Cen B†	14 41.1−60 56	1.35	0.9	K1 V†	750	5.7	4.3	~3.703	~283	−21 V?	AB 5.2″ (2019) orbit 79.9y min = 2″ (1955); max = 22″; PA (2017) 325°; separation 11.2 AU min, 35.6 AU max; Kaler at stars.astro.illinois.edu/sow/rigil-kent.html has map of apparent AB orbit (note further here that Kaler’s green, violet, and blue denote micrometry, photography, and interferometry, respectively: as the error bars suggest, the α Cen AB orbit is one of the most precisely known wide binary-system orbits in visual-binary astrometry); since plane of orbit is inclined at 79° to plane of sky, the apparent orbit is more severely elliptical than the true orbit (for which *e*=0.5); Kaler’s map can accordingly be usefully supplemented with the apparent-versus-true-orbit diagram in en.wikipedia.org/wiki/Alpha_Centauri ¶ whereas magnetic activity of α Cen A is in steep decline since 2005 (analogue of Maunder Minimum? or, rather, mere regular cycle?), α Cen B shows more magnetic activity than α Cen A does, and its cycle is brief (8.2 y in spot numbers, 16.4 y in magnetics; this is not unlike the Sun, for which the corresponding pair of periods is ~11y, ~22 y); and 2005A&A...442…315R reports a flare on α Cen B ¶ although 2012 α Cen B exoplanet claim is now discounted, an exoplanet is possible (2019 transit has been suggested; cf also 2015MNRAS.450.2043D) ¶ Einstein-ring event expected with 45% probability in 2028, early in May	Toliman
α	Cen A+1P†	14 41.1−60 56	−0.01	0.71	G2 V†	750	4.4	4.3	~3.710	~277	−22 SB	Ca (Proxima), 12.4, M5e, 2.2° SW of A still the closest known object in the population of stars and brown dwarfs, despite intense surveying of entire population over the past 20 or 30 years ¶ gravitational binding of AB+C was finally established with high probability in 2017A&A…598L…7K, and an orbit is considered to be known (~550,000 y: min > 4300 AU, max 13,000 AU) ¶ Ca is elusive, as one faint object in a sea of faint objects (detection was not achieved until the 1915 work of Innes, with blink comparator); nevertheless, violent flaring has been known to take Ca, briefly, up to the threshold of naked-eye visibility (2018ApJ...860L..30H reports peak V mag. 6.8 for a superflare of 2016 March 18, of duration ~1 h; the situation in UV and millimetre waves is also extreme, as reported in 2021ApJ...911L..25M) ¶ 2016Natur.536..437A announces an approx. Earth-mass exoplanet, α Cen Cb, in the habitable zone of its host α Cen Ca; 2018ApJ...860L..30H analyzes the germicidal implications of flaring, finding that in a habitable-zone Earth-like exoplanet atmosphere the ozone UV shield would be destroyed; this paper, however, like some others, leaves open the possibility of life on Cb; breakthroughinitiatives.org/initiative/3 advocates nanocraft exploration, and en.wikipedia.org/wiki/2069_Alpha_Centauri_mission has footnote links to reports of small-scale discussions	**Rigil Kentaurus**

												at NASA and in the USA Congress, envisaging the launch of some (nanocraft?) mission to celebrate the Apollo 11 centenary ¶ in 2020, a more distant exoplanet, either a super-Earth or a mini-Neptune, was suggested, as α Cen Ac (orbit 1930 d, whereas Ab has orbit 11 d; an unexpectedly bright detection with the VLT SPHERE instrument has been interpreted as the possible signature of rings around the putative Ac); an exoplanet Cd was announced in 2022 Feb.	
α	Lup A	14 43.4−47 29	2.30†	−0.15	B1.5 III	7	−3.5	460	0.032	221	+5 SB	sight var. of β Cep type, 2.29–2.34 in V, 0.26 d actually multiperiodic, with primary period (unusually long) 0.2598466 d given by AAVSO(VSX) as viewed 2021 Jan. 16, and again 2022 March 3	
α	Cir A†	14 44.4−65 04	3.18†	0.26	A7 Vp (Sr)	60.4	2.1	54.1	0.303	220	+7 SB?	B: 8.5, K5 V, 15.7″, PA 263°→224°, 1826→2016 AB probably true binary, with orbit ≥ 2600 y ¶ Astron. Alm. (epoch 2021.5) assigns MK temperature type A7p Sr Eu and does not assign an MK luminosity class ¶ the brightest (slight) variable of the AAVSO "rapidly oscillating Ap" type (features of the type include rapid non-radial pulsation with a stellar-rotation signal), with V mag. range 3.17–3.19; magnetically an oblique rotator (4.4790 d, with field strength ~500× solar); 2009MNRAS.396.1189B discusses the rotation, two notably stable putative equatorial chemical-anomaly regions, and asteroseismology, with history and fresh WIRE+SAAO observations	
ε	Boo A	14 46.0+26 59	2.58	1.34	K0 II–III†	16†	−1.6	200†	0.044	288	−17 V	B:4.81, 2.8″, PA 318°→347°, 1822→2018 orbit well over 1000 y ¶ ε Boo B is of MK type A0 V, and is SB, with at least one component a rapid rotator ¶ stars.astro.illinois.edu/sow/izar.html discusses difficulties in determination of the individual magnitudes and of the binary system's distance ¶ F.G.W. von Struve: "pulcherrima" ("the loveliest")	Izar
β	UMi A+1P†	14 50.7+74 04	2.07	1.46	K4 III†	24.9	−0.9	131	0.035	289	+17 V	useful for aligning small equatorial mount (since NCP, although not quite coincident with α UMi, does lie near the great-circle arc linking β UMi with α UMi: arksky.org/Kochab.htm) ¶ Fe underabundant, Ba possibly slightly overabundant ¶ 2008A&A...483L..43T suggests (via CORIOLIS-SMEI) two short-lived radial-pulsation mods ¶ 2014A&A...566A..67L announces exoplanet	**Kochab**
α	Lib Aa†	14 52.1−16 08	2.75	0.15	A3 III–IV†	43	0.9	76	0.126	237	−10 SB†	B. 5.2, 231″ (2012) angular distance from α Lib B, which shares the proper motion of α Lib A, entails Aa,Ab–B distance ≥ 5500 AU; if B and the A system are gravitationally bound, then their period is ≥ 200,000 y; alternative names for the α Lib Aa,Ab pairing and the single star α Lib B are α^2 Lib and α^1 Lib, respectively, with "1" signalling the fact that α^1 Lib, lying to W of α^2 Lib, although fainter than "2", is the earlier of the two in its crossing of the local meridian ¶ α Lib Aa, Ab are resp. mags. 3.30, 3.70; for this tight pairing (an SB as well as a resolved system), WDS documents just 4 measurements (from 2017), for angular separations perhaps < 0.1″ (distance between α Lib Aa and α Lib Ab may be a few tenths of 1 AU); one of α Lib Aa,Ab is overabundant in some metals, perhaps due to the influence of its SB companion; α Lib B is likewise now resolved as α Lib Ba, α Lib Bb, at resp. mags. 4.30 and 7.70, with just 2 astrometry measurements (1999 and 2018, angular separations 0.4″, 0.2″ respectively); celestial-sphere neighbours α Lib C and α Lib E are faint; gravitationally bound neighbour D (at very small angular separation) is mag. 7.31 (only 3 astrometry measurements, 1991→1999)	**Zubenelgenubi**

												¶ lunar occultations are possible, planetary occultations possible yet rare
β	Lup	15 00.0−43 13	2.68†	−0.18	B2 IV†	9	−2.7	380	0.054	222	0 SB	has been claimed to be slight (β Cep) var.: dominant period 0.232 d ¶ fast rotator (< 3.4 d) ¶ low metallicity
κ	Cen Aa†	15 00.6−42 12	3.13†	−0.21	B2 V	9	−2.2	400	0.029	218	+8 SB	strictly a triple system, Aa+Ab+B; B mag. 11.5; AB 4″, PA 84°→83°, 1926→2000, separation ≥ 470 AU; ≥ 3000 y; Aa-to-Ab distance possibly ~10 AU, period possibly ~10 y (stars.astro.illinois.edu/sow/kappacen.html discusses various physical uncertainties), with mags. resp. 3.34, 4.71, and with angular separation 0.1″→0.1″, 1991→2020 ¶ line profiles vary; although the Aa+Ab binarity has made variability classification difficult, κ Cen Aa is classified by AAVSO(VSX), viewed 2021 Jan. 16 and 2022 March 3, as a (slight) variable of the β Cep type (3.13−3.14 in V, 0.095325 d)
β	Boo	15 02.8+40 18	3.49	0.96	G8 IIIa†	14.5	−0.7	230	0.049	234	−20 V?	Nekkar Ba 0.4, Fe −0.5 1995A&A...296..509H discusses the puzzling flare seen by *ROSAT* 1993 Aug. 08 (unusual for a lone M giant; it is possible, but seems unlikely, that flare came instead from an undetected M-dwarf companion; the mild Ba enhancement is, admittedly, consistent with presence of such a companion); slow rotator (~200 d)
σ	Lib	15 05.4−25 22	3.25v†	1.67	M2.5 III	11	−1.5	290	0.083	239	−4	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]** semireg. var.: 3.20–3.46 in V, mean period 20 d Brachium there is also rapid microvariability ¶ highly evolved (on AGB, with dead carbon-oxygen core)
ζ	Lup A†	15 13.9−52 11	3.50	0.92	G8 III	~27.8	0.6	117	0.133	238	−10	B: 6.74; 71.7″ (2020), PA 249°→249°, 1826→2020 A-to-B distance ≥ 2600 AU; shared proper motion suggests true binarity (period possibly ≥ 68,000 y) ¶ ζ Lup A is in evolutionary terms on “Red Clump” (Sun-like when still on MS, but helium flash now finished, core-helium fusion now underway)
δ	Boo A†	15 16.4+33 14	3.46	0.96	G8 III Fe–1†	~26.8	0.6	122	0.140	143	−12 SB	a very wide true binary: B is mag. 7.89, 105″ (2017) PA 84°→78°, 1780→2017, A-to-B distance ≥ 3800 AU, period 120,000 y (with shared proper motion indicating true binarity); the SB which is δ Boo A is not as yet resolved, even in interferometry (so WDS cannot as yet write “δ Boo Aa”, “δ Boo Ab”) ¶ δ Boo A is CN weak; δ Boo B could be a subdwarf, consistently with the observed low metallicity of δ Boo A ¶ δ Boo A is in evolutionary terms a “Red Clump” star (core-helium fusion now underway)
β	Lib	15 18.2 −9 28	2.61†	−0.07	B8 IIIn	~17.6	−1.2	190	0.100	259	−35 SB	Zubeneschamali flagged by AAVSO(VSX), viewed 2021 Jan. 16 and 2022 March 3, as “suspected [slight] variable lacking deeper studies,” with V mag. 2.60–2.62 (and yet Eratosthenes, resp. Ptolemy, asserted β Lib to be brighter than, resp. equal to, α Sco) ¶ rapid rotator ¶ E(B–V) =–0.02
γ	UMi	15 20.7+71 45	3.00†	0.06	A3 III†	6.7	−2.9	490	0.025	315	−4 V	Pherkad a rapid rotator, and (despite being in MK type A, not B) said to be a variable shell star (cf 2000A&A...354..157H; BSC5: “shell possibly variable,” H and CaII variable); AAVSO(VSX), however, viewed 2021 Jan. 16 and 2022 March 3, classifies this as a low-amplitude variable in the δ Sct group
γ	TrA	15 21.0−68 46	2.87	0.01	A1 IIIn†	17.7	−0.9	184	0.074	244	−3 V	has been asserted to be chemically anomalous (Eu overabundance), and also, not quite consistently, has been classed as a rapid rotator (< 1.2 d)

												¶ although we here give MK luminosity class III, class V has also been asserted; Astron. Alm. (epoch 2021.5) assigns MK type A1 III ¶ IR excess has been asserted (circumstellar disk?)	
δ	Lup	15 22.9−40 44	3.22†	−0.23	B1.5 IVn	4	−3.9	900	0.032	218	0 V?	rapid rotator (< 2.4 d) ¶ a (low-amplitude) variable in the β Cep Group, 3.2–3.24 in V (AAVSO(VSX), viewed 2021 Jan. 16 and 2022 March 3), with a single period known, 0.16547 d (cf 2007MNRAS.377..645S)	
ε	Lup Aa†	15 24.2−44 46	3.37†	−0.19	B2 IV–V	6	−2.6	500	~0.030	~230	+8 SB2	A: 3.56; B: 5.04, 0.1″, PA 285°→53°, 1883→2019 orbit 737 y: in more detail, a (probable) hierarchical quadruple; although B experiences A as essentially a point mass, in fact A is SB, interferometrically resolved as ε Lup Aa, ε Lup Ab with a single 2010 measurement (mags. 3.60, 5.10, angular separation 0.1″), for which 2005A&A...440..249U gives SB period 4.55970 d (classifying the primary as a suspect β Cep variable and the secondary as a new β Cep variable); experiencing AB, on the other hand, as essentially a point mass is the (probably) gravitationally bound C (mag. 9.10; 19″→26″, PA 174°→168°, 1826→2020; AB-to-C distance ≥ 4100 AU; if gravitationally bound, then period ≥ 60,000 y); in its stable kinematics, this putative hierarchical quadruple may be contrasted with the unstable, nonhierarchical θ Ori system, and in its detailed organization with the stable, hierarchical, but mere "double-double" ε Lyr system; AAVSO(VSX) as viewed 2021 Aug. 11 and 2022 March 3, 2022 indicates just slight variability, 3.36−3.38 in V	
ι	Dra A+1P†	15 25.4+58 53	3.29	1.17	K2 III†	32.2	0.8	101	0.019	334	−11	2002ApJ...576..478F announces substellar-mass companion and discusses possibility of transits; this is the first discovery of a planet or brown dwarf (IAU name: Hypatia) orbiting a star that has finished stable core-hydrogen fusion; exoplanet.eu/catalog/HIP%2075458_b/ may from time to time have updates; its substellar companion notwithstanding, ι Dra has metallicity only slightly greater than solar ¶ since the ι Dra A substellar-mass companion is known through radial-velocity work, and has not yet been resolved (not even interferometrically), WDS is not as yet able to write "ι Dra Aa", "ι Dra Ab"; the widely separated celestial-sphere neighbour catalogued in WDS as ι Dra B (mag. 8.87, 255″→253″, PA 50°→51°, 1879→2020) is not gravitationally bound to the ι Dra A system)	Edasich
α	CrB	15 35.6+26 38	2.22v†	0.03	A0 IV (composite) †	43	0.4	75	0.150	127	+2 SB†	ecl.: 2.21–2.32 in B band, 17 d (more precisely, from AAVSO(VSX) as viewed 2021 Jan. 16 and 2022 March 3, 17.359907 d): this SB is a detached binary, with neither component filling its Roche lobe; distance between components 0.13 AU min; as with β Per, so also with α CrB, instrumental photometry reveals both the primary and the secondary eclipse; components have not been interferometrically resolved (so WDS-conformant designation is still "α CrB", not "α CrB A" and "α CrB B") ¶ individual MK types are difficult: primary possibly A0 V, secondary possibly G5 ¶ primary has IR excess (debris disk?) ¶ secondary is X-ray visible and is a rather rapid rotator (~9 d or ~7 d or less, so not tidally locked) ¶ non-IAU name Gemma denotes α CrB as "gem of the Northern Crown"	**Alphecca**
γ	Lup A†	15 36.6−41 14	2.80	−0.22	B2 IVn†	8	−2.8	400	~0.030	~212	+2 V	A: 3.0; B: 4.4; similar spectra 0.8″ (2019) PA 94°→275°, 1835→2019;	

													max angular distance 1980, min ang. dist. 2075; orbit 190 y: γ Lup AB orbit is seen nearly edge-on; separation 41 AU min, 128 AU max, 84.5 AU average; stars.astro.illinois.edu/sow/gammalup.html has an orbit map, showing that observational coverage is imperfect (green for micrometry (with large error bars), violet for photography, blue for interferometry); γ Lup A is itself SB (2.801 d, unresolved), making this a hierarchical triple system, with the primary in the γ Lup A pairing a fast rotator (< 1 d, so not tidally locked) ¶ BSC5 asserts expanding circumstellar shell, and (citing 1987 Vainu Bappu spectra) notes emission peaks in Hα profiles, says possibly in transition from B to Be	
α	Ser A	15 45.4	+6 21	2.63?†	1.17	K2 IIIb CN1†	44	0.9	74	0.141	71	+3 V?	slight semiregular variable? (range ~0.2 in V?) ¶ a "strong-lined giant" (although [Fe/H] metallicity is not very much above solar) ¶ a modest X-ray source ¶ has borne also the (not IAU-official) name Cor Serpentis ("Heart of the Serpent"), despite being the principal luminary of Serpens Caput ("Serpent Head")	Unukalhai
μ	Ser A	15 50.8	−3 30	3.54	−0.04	A0 III	19	0.0	170	0.104	255	−9 SB†	binary resolved with speckle interferometry, and subsequently (2010NewA...15..324G) analyzed with astrometry: B is mag. 5.39, 0.2″→0.4″, 1991→2018; 2010NewA...15..324G offers an orbital solution, with period 36±2 y, e=0.4±0.3; these authors remark that the low precision of their orbit-based mass determinations leaves various possibilities open regarding the nature of μ Ser B ("A or F dwarf, subgiant, giant or even a pair of late-type dwarfs")	
β	TrA A	15 57.1	−63 30	2.83	0.32	F0 IV†	~80.8	2.4	40.4	0.444	205	0	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]** *Spitzer Space Telescope* finds IR excess (debris disk?) ¶ rapid rotator (slightly < 1 d), with detectable magnetic field ¶ metals vary widely (some overabundant, some underabundant) ¶ duplicity analyzed through speckle interferometry: A is mag. 3.75, B is mag. 5.39, 0.2″→0.4″, 1991→2018	
π	Sco Aa	16 00.2	−26 11	2.89†	−0.18	B1 V†	6	−3.4	600	0.029	203	−3 SB2	Aa,Ab ecl.(?) SB; mags. 3.4, 4.5; 1.57d; 2.88−2.91 in V (more precisely, from AAVSO(VSX) viewed 2021 Dec. 3, 1.570103 d (but a published orbital solution gives instead 1.5700925 d), WDS asserts, and AAVSO(VSX) denies, that orbit is so close to edge-on as to be eclipsing (if not eclipsing, then the observed slight variability is the effect of ellipsoidally distorted stars presenting different surface areas to the photometer at different stages in their mutual orbit); orbit is circular or nearly circular (two published orbital solutions disagree slightly, asserting e=0, e=0.15), possibly with tidal locking, Aa-to-Ab distance possibly ~0.07 AU; the binarity has been detected also via occultation; WDS documents just one Aa,Ab measurement (in the year 2000, with angular separation 2″); although system has been said to be of β Lyr type, the AAVSO(VSX) classification is, rather, "rotating ellipsoidal variable" (the stars so close as to be gravitationally distorted into ellipsoids, but neither star deformed into the teardrop shape possible in one β Lyr scenario (the β-Lyr variability-type scenario, namely, in which a component becomes so grossly distended as to fill its Roche lobe; in any β Lyr variable, the shape distortion is by definition so severe as to leave no constant-light segments in the light curve));	Fang

												inspection of AAVSO archive, as viewed 2021 Dec. 3, indicates a longstanding shortage of photometry (and Kaler at stars.astro.illinois.edu/sow/pisco.html additionally discusses some difficulties in astrophysical modelling)) ¶ E(B–V) =+0.08
T	CrB A	16 00.4+25 51	9.8v†	1.34	M3 III†	—	0.6	2500?	0.011	329	−29 SB	recurrent nova 1866&1946 mags 3&2; V~9.6 2022 Feb. 15 only ten galactic recurrent novae are currently known (2010ApJS..187..275S; these are by definition novae known to recur, and yet lacking the short periods of dwarf novae) ¶ T CrB A partner in the recurrent-nova activity, T CrB B, is WD with hot circumstellar accretion (dominating the aggregate T Cr AB signal in UV) of MK type Bep, orbit 227.5 d or 227.6 d, A-to-B distance ~0.5 AU; angular separation has been measured only twice (in 1946 (considered doubtful by WDS) and 2010, as 0.3″ and 0.7″ respectively) ¶ long documented in Handbook as mag. 10.08, T CrB AB (combined light) brightened from February 2015, attaining ~9.2 in April 2015 (with V mag. 9.555, on the other hand, is reported in general AAVSO database (not AAVSO(VSX) database) for 2022 Feb. 15); Bob King in *Sky&Telescope* 2016 Apr. 20 gives recent history, and AAVSO has a backgrounder at www.aavso.org/t-crb; next eruption 2026, or earlier?
η	Lup A†	16 01.6−38 28	3.37	−0.21	B2.5 IVn†	7	−2.2	440	0.033	211	+8 V	B: 7.50, 14.8″, PA 22°→19°, 1834→2020 orbit ≥ 26,000 y: a hierarchical system, with remote outlier D at angular distance 135″ (separation ≥ 18,000 AU, period ≥ 750,000 y), with D experiencing the AB pair as essentially a point mass; η Lup C is not part of this (triple) system, C's angular proximity to AB being a mere line-of-sight coincidence (mag. 9.39; 60″→115″, PA 245°→248°, 1825→2015) ¶ although η Lup A is a rapid rotator (< 1.1 d), there is no evidence of a circumstellar disk, and in particular there seems to be no documentation of "Be phenomenon" spectral behaviour ¶ η Lup B is chemically peculiar
δ	Sco A†	16 01.7−22 41	2.29v†	−0.12	B0.5 Ve†	7	−3.6	440	~0.037	~196	−7	pro-am photm., spectr. data needed esp. 2022-May Dschubba (when this binary passes through periastron): the δ Sco AB interferometrically resolvable SB has been measured since 1973, although binarity was reported from lunar occultation as early as 1901; period is 10.8 y; previous recent periastra were in 1990, 2000, and 2011; orbit is discussed in. e.g., 2012ApJ...757...29C (slightly refining the orbital solution of 2011ApJ...729L...5T); orbit is remarkable for its extreme elongation (*e* = 0.94; for most Be binaries with a non-degenerate secondary component, *e* ≈ 0); the suggestion that δ Sco A is itself a tight binary, with period ~20 d, is not now generally favoured (2013ApJ...766..119M argues against the suggestion); connected with the unexpectedly high-eccentricity AB orbit, however, might be some as-yet-undetected distant orbiter, with period ~200 y, perhaps participating in a Lidov-Kozai interaction (2013ApJ...766..119M, and additionally cf. en.wikipedia.org/wiki/Kozai_mechanism); AB angular separation can become as great as ~200 mas, but at periastron diminishes to 5.9 mas, with AB physical distance diminishing to within 0.8 AU (a distance ~25x the radius of δ Sco A, so tidal interaction, perhaps even the generation of a tidal trail of ejecta, is to be expected at periastron); orbital plane is inclined only rather gently to the plane of the sky (2020ApJ...890...86S gives the angle as 38°), making this

binary far from eclipsing
¶ important instance of "Be phenomenon", offering opportunity to examine a recent disk-building event: δ Sco system seemed unremarkable in much of the 20th century, and was even taken as a B0 IV MK standard; 1993A&A...274..870C, however, reported Be phenomenon, from spectroscopy at or near the 1990 periastron; at or near the periastron of 2000, Be-phenomenon behaviour in δ Sco A became for the first time strongly evident, with pronounced Hα emission in spectroscopy, and also with brightening in photometry; the system faded somewhat in 2005, both in V band and in IR, while Hα equivalent width (a signature of material toward the outer reaches of the Be-phenomenon disk around δ Sco A) increased; the system again brightened in V band in 2010 (a signature of material being added to the inner reaches of the disk), and stayed bright, with minor V fluctuations, through the 2011 periastron (and for at least some of this period, notably for 2009 through 2012, cyclic photometric variability was observed on timescales of ~60 d to ~100 d (similar behaviour had been found also for 2000 through 2002; but the orbital periods of the inner and outer portions of the disk are on the order of 0.5 d and 1.5 d, and so are on a different timescale), consistently with a variable rate of mass transfer upward to the disk out of the δ Sco A photosphere); the speed of disk growth seems unusually high in the general population of Be stars, where an episode of growth can take decades, and yet rapid rotator δ Sco A does not appear to be rotating so rapidly as to diminish effective photosphere gravity at equator down to ~0; on the modelling of 2020ApJ...890...86S, the Be-phenomenon Hα-emitting portion of the disk of equatorially ejected gas around δ Sco A was of radius 10 R* in 2000, 14 R* in 2002, 11 R* in 2007 (there was a temporary partial dissipation of the disk in 2005), 46 R* by 2018 (the temporary partial dissipation was followed on this modelling by a period of variability from 2005 to 2009, and by a disk-growth process from 2010 to 2011, with a rather steady state attained from 2011 to 2018 or beyond); 2020ApJ...890...86S remains agnostic on the question whether gravitational perturbations, especially at the outer reaches of the disk, have affected dissipation and growth (for instance, through tidal effects, including the tidal locking, at periastron, of a local density enhancement, such as a spot or a spiral wave? with disk possibly even overflowing δ Sco A Roche lobe at periastron, yielding mass transfer to δ Sco B (although mass transfer is rejected by a CHARA disk-imaging team, at 2012ApJ...757...29C)?)); it is to address this question that photometry and spectroscopy are sought, especially around the 2022 May periastron, both from professionals and from amateurs (cf pro-am 2022-campaign request at www.aavso.org/delta-sco-campaign; AAVSO additionally has a circa-2011 background briefing, with emphasis on the photometry, at www.aavso.org/vsots_delsco; 2013ApJ...766..119M describes pro-am spectroscopy contributions at the 2011 periastron, involving on the amateur side nearly 20 observers from Australia, France, Germany, Portugal, Spain, and the USA)
¶ our apparent and absolute magnitudes do not reflect the post-2000 combined light of δ Sco AB: www.aavso.org/delta-scorpi has recent forum discussion, notably on choice of comparison stars for visual photometric estimates; 4 typical AAVSO V-filter photometry reports, from one and the same observer, are 2020 June 20 mag. 1.86, 2020 June 26 mag. 1.81, 2021 June 26 mag. 1.82, and 2021 July 5 mag. 1.78 (suggesting, therefore, just slight variability in V at present, below the naked-eye threshold); δ Sco B (not known

												to be variable) is reported by WDS as of mag. 4.62 ¶ we follow 2020ApJ...890...86S in assigning MK type B0.5 V, which we take to be appropriate for the combined binary-system light (but elsewhere in the literature, a slightly different MK type is assigned); 2020ApJ...890...86S suggests B2 V as an MK basis for modelling δ Sco B ¶ we follow 2020ApJ...890...86S in assigning, consistently with our policy for using rounding-off to reflect uncertainties, D = 440 ly, and on this basis asserting π to be, with reasonable rounding-off, 7 mas ¶ δ Sco AB (and its remoter gravitationally bound third star, if there is such a companion) may possibly be a low-velocity runaway system with ISM bow shock ¶ E(B–V)=+0.16	
β	Sco Aa†	16 06.7−19 52	2.56	−0.06	B0.5 V	8	−2.9	400	0.025	192	−1 SB	Aa: 2.9; B: 10.6, 0.3″ (2019); C: 4.52, 13″ (AC astrometry in more detail: 14″→13″, PA 25°→20°, 1779→2019); in gross terms a visual binary (as AC), with A-to-C distance ≥ 2200 AU, period > 16,000 y, but in fact putatively a sextuplet; en.wikipedia.org/wiki/Beta_Scorpii summarizes the sextuplet hierarchy in a diagram (Aa with Ab (6.82 d), and B experiencing Aa+Ab as essentially a point mass (610 y); E as an as a not-yet-resolved SB (although Wikipedia in an informal spirit writes "Ea" and "Eb", WDS, whose terminology we in this Handbook article take as normative, cannot as yet do so; period of this unresolved SB is 10.7 d), and C experiencing the binary system which is as yet just (in WDS-formal terms) "β Sco E" as essentially a point mass (39 y, β Sco E combined light mag. 6.60); the B+(Aa,Ab) triple is in a wide, > 16,000-y orbit with the C+E triple, around the centre of mass shared by this pair of triples, thereby delivering the gross visual-binary phenomenology studied as β Sco AC since 1779; the Aa,Ab angular separation is too small to yield a measure for WDS; the CE separation is tight, measured as 0.1″ in 2019; the entire sextuplet has an outlying celestial-sphere neighbour, physically unrelated to the sextuplet, WDS-catalogued as β Sco D (mag. 7.5; 520″→518″, PA 31°→30°, 1860→1998) ¶ lunar occultations possible, planetary occultations possible yet rare (1971 May 14 occultation by Jovian satellite Io) ¶ the name Graffias is not IAU-official	Acrab
δ	Oph A	16 15.5 −3 45	2.73†	1.58	M1 III†	~19.1	−0.9	171	0.150	198	−20 V	slight var.? (2.72–2.75 in V?) ¶ slow rotator ¶ high metallicity ¶ although δ Oph has finished core hydrogen fusion, its exact evolutionary state is uncertain (cf stars.astro.illinois.edu/sow/yedprior.html) ¶ Astron. Alm. (epoch 2021.5) assigns MK type M0.5 III ¶ naked-eye neighbour Yed Posterior is a mere optical companion, too greatly separated in space for true binarity; the "prior" and "posterior" in the traditional, and as of 2016 IAU-official, names denote the order in which these two (physically unrelated) stars cross the local meridian ¶ listed in NSV as a suspected variable, and in AAVSO(VSX), viewed 2021 Jan.16, as unobserved; 1992IBVS.3792....1P finds no variability, but says that since NSV V-amplitude is just 0.03 mag., variability cannot be excluded	Yed Prior†
ε	Oph A	16 19.5 −4 45	3.23	0.97	G9.5 IIIb†	31	0.7	106	0.093	64	−10 V	cyanogen and carbon notably underabundant, suggesting that ε Oph is an interloper from outside the galactic thin disk; Astron. Alm. (epoch 2021.5) assigns MK type G9.5 IIIb Fe–0.5	Yed Posterior
σ	Sco Aa1†	16 22.6−25 39	2.91†	0.13	B1 III†	5	−3.7	700	0.019	213	+3 SB†	slight var.: 2.86–2.94, 0.25 d; B: 8.4, B9 V, 20.5″ (2019)	Alniyat

												(AB astrometry in more detail: 22″→20″, PA 270°→269°, 1783→2019); recent studies, including lunar occultation measures, show σ Sco to be a quadruple system, with σ Sco Aa1,Aa2 in fact SB (33.0 d; considered by WDS to be successfully resolved, although only one PA measurement, from 2010, is documented in WDS), and with the entire 3-star σ Sco A configuration (where the (Aa1+Aa2), Ab period is > 100 y) in some slow, wide orbit with σ Sco B; orbital solution has been published for Aa1,Aa2, but not for Aa,Ab; 2007MNRAS.380.1276N announces interferometric solution for the SB orbit, proposing for primary and secondary the respective MK types B1 III, B1 V; in the SB pair, the primary is a variable of the β Cep type (AAVSO(VSX) as viewed 2022 Feb. 18 gives V-mag. range 2.86–2.94, period 0.246839 d; 1992A&A...261..203P discusses period changes) ¶ photography shows σ Sco to be embedded in diffuse nebula ¶ E(B–V) =+0.4 (pronounced reddening)	
η	Dra A	16 24.3 +61 28	2.73†	0.91	G8 IIIab	35.4	0.5	92	0.059	343	−14 SB?	B: 8.2, 4.4″, PA 150°→143°, 1843→2015 orbit ≥ 1000 y, separation ≥ 140 AU ¶ a “Red Clump” resident (evolved, presently stable, performing core-helium fusion) ¶ believed to be a slow rotator (~400 d) ¶ a modest X-ray source ¶ both the SB which is η Dra A and its close celestial-sphere neighbour η Dra B are suspected of variability (cf NSV (Kukarkin et al)) ¶ the η Dra AB pair has a reasonably bright outlying celestial-sphere neighbour, WDS- catalogued as η Dra C: mag. 8.10, angular separation from η Dra A ~565″, PA ~240°, angular separation and PA nearly constant 1866→2003 ¶ near the radiant of the η Draconids meteor shower	Athebyne
α	Sco A†	16 30.8 -26 29	1.06v†	1.86	M1.5 Iab†	6	−5.1	600	0.026	207	−26 SB	semireg. var.: 0.75–1.21, 5.97 y; B: 5.40, 2.8″ (2019) PA 273°→276°, 1847→2016; orbit 2500 y? the large magnitude difference makes this a difficult binary (the binarity was discovered only in 1819, through controverted lunar-occultation work (1972JBAA...82..431K describes the 18.6-year 1940-through-2050 cycle of lunar occultation possibilities); the binarity was established beyond controversy still later, in the 1840s) ¶ AAVSO(VSX): semireg. (with some discussion of period; cf also 2013AJ....145...38P, where a true period is found for radial-velocity variations, and the detected variation is judged to be more likely of pulsational than of orbital origin), V mag. 0.75–1.21 (but the variability has also been called irregular); 2018AuJAn..29...89H reports that variability was observed by, and incorporated into the oral tradition of, aboriginals in southern Australia; asserted by Eratosthenes to be fainter than β Lib, and by Ptolemy to equal β Lib ¶ Astron. Alm. (epoch 2021.5) assigns MK type M1.5 Iab–Ib ¶ radius has been studied interferometrically and via lunar occultations (up to 3.4 AU; however, even apart from the problem of pulsation, radius determination of highly evolved red stars is wavelength-dependent); one of the two first-mag. supergiants (the other being α Ori Aa (Betelgeuse)) ¶ significant stellar wind, with mass loss	**Antares**†

												almost 1e-6 $M_{\odot}$/y, within which α Sco B has created a locally ionized region ¶ the most massive member of the Sco-Cen Association (the nearest OB association) ¶ B shares in the proper motion of A, indicating true binarity: AB separation is ≥ 530 AU, period possibly ~1200 y ¶ location (within zodiac) makes the classical Greek name for “rival of Mars” appropriate not only as regards naked-eye colour but also as regards sky geometry ¶ 1972JBAA...82..431K describes the 18.6-year 1940-through-2050 cycle of lunar occultation possibilities
β	Her Aa	16 31.2 +21 27	2.78†	0.95	G7 IIIa†	23	−0.4	140	0.100	261	−26 SB†	Kornephoros† SB period computed 1908, and again 2008, in both cases ~410 d; 1977ApJ...214L..79B announces speckle-interferometry resolution of the β Her Aa,Ab SB, with angular separation 43 mas; WDS documents 4 astrometric measurements of the Aa,Ab SB as a visual binary, 1975→1984; Aa,Ab possesses, and AB lacks, a published orbit ¶ suspected slight variable (NSV (Kukarkin, et al., online) suggests V-mag. range 2.76–2.81, and AAVSO(VSX) as viewed 2022 Feb. 18 concurs) ¶ X-ray emission from the SB primary indicates magnetic activity ¶ Astron. Alm. (epoch 2021.5) assigns MK type G7 IIIa Fe–0.5 ¶ Kaler, noting that primary has N enhanced relative to C, says in his overall summation “a very normal star for its state of age” ¶ “Kornephoros” = Gk “club-bearer,” in reference to the weapon of Hercules (compare α Her, which in the pictorial-atlas tradition, marks the hero’s head)
τ	Sco	16 37.3 −28 16	2.82	−0.21	B0 V†	7	−3.0	500	0.025	203	+2 V	Paikauhale† intrinsically more luminous than σ Sco, but more heavily obscured by ISM ¶ anomalous in its UV lines (P Cyg profile) ¶ O and Fe are underabundant ¶ 2006MNRAS.370..629D discusses τ Sco magnetic topology (poloidal, with also a warped toroidal component of modest strength), including both its origin (more likely a fossil field from the star’s (recent) birth than a dynamo effect) and its connection with winds and with the observed hard X-ray emission; the authors note that the topology is stable over the 1.5-y period of their observations (in contrast with a strongly differential-rotation star, such as Sun); in additionally announcing a (refined) rotation period of 41.033 d, the authors comment, “the second-slowest rotator so far known among high-mass stars” ¶ Kaler: “among the most-observed stars in the sky” ¶ E(B-V)=+0.06 ¶ the τ Sco name Paikauhale was IAU-approved in 2018 Aug. 10; the not-IAU-official “Al Niyat,” or “the arteries of the Heart,” on the other hand, denotes σ Sco and τ Sco jointly, as flanking α Sco
ζ	Oph	16 38.4 −10 37	2.54†	0.04	O9.5 Vne†	9	−2.7	370	0.029†	32	−15 V†	the nearest O-type star (and consistently with this extreme temperature, resident in an H II region) ¶ unusual in being an “Oe”, i.e. an O-star instance of the “Be phenomenon” ¶ “runaway star” (consistently with this extreme speed-relative-to-LSR, forming bow shock in ISM), perhaps formerly the secondary in a binary pair whose primary perished in a supernova;

													2011AN....332..147H confirms magnetic field, discusses X-ray properties, suggests PSR B1919+10 as remnant of the hypothesized defunct companion ¶ line of sight to ζ Oph is one of the most used in spectroscopic studies of ISM ¶ 2014MNRAS.440.1674H is a recent discussion of variability, from radial and non-radial pulsation modes; AAVSO(VSX), assigning magnitude range 2.56–2.58 in V, follows GCVS in treating ζ Oph as a variable with Be-phenomenon behaviour, and yet lacking the history of outbursts founds in the γ Cas class; ζ Oph is, on the other hand, classified as γ Cas-variable (and is termed a shell star) in BSC5; still elsewhere, ζ Oph has been treated as a prototype for the "ζ Oph variables" ¶ E(B-V)=+0.32 (pronounced reddening; if ISM were not present, ζ Oph would reach nearly first mag.) ¶ recapitulations of recent ζ Oph studies include 2012MNRAS.427L..50G (MK classification problem, also mass-loss rate in context of "weak-wind problem"), 2014MNRAS.440.1674H (rotation, pulsation, Hα emission episodes, inferred circumstellar decretion disk, satellite-based photometry), 2015ApJ...800..132C (distance, age, mass, effective temperature, bow shock in ISM, ...); additionally, 2012A&A...543A..56D is among the papers describing not only the specific interaction of ζ Oph with ISM, but also the quite general ISM bow-shock topic (noting inter alia that not all runaway stars produce bow shocks)
ζ	Her A†	16 42.1+31 34	2.62	0.65	G1 IV†	93	2.7	35	~0.575	~307	−70†	SB	B: 5.40, G7 V, 1.6″, PA 110° (2019), orbit 34.45 y orbit well studied since F.G.W. von Struve 1826 micrometry (however, it was Herschel, not von Struve, who discovered the binarity); separation 8 AU min, 21 AU max, 15 AU average, 34.45 y; considered one of the few binaries in which ratio of B mass to sum of A and B masses can be studied both via traditional (non-interferometric) astrometry and via spectroscopy; WDS indicates, however, that an inner binary with orbit ~12 y has been suspected repeatedly, and that the inner-binary component has been detected in IR speckle interferometry (but WDS, at any rate as viewed 2021 Dec. 13, continues to write simply "ζ Her A", without as yet distinguishing between ζ Her Aa and ζ Her Ab) ¶ Astron. Alm. (epoch 2021.5) assigns MK type G0 IV ¶ ζ Her A is unusual in its evolutionary phase, being in the Hertzsprung Gap (and so in rapid evolutionary transition) ¶ 2001A&A...379..245M summarizes previous work, presents detailed physical modelling for A and B, and discusses asteroseismology, remarking in conclusion that "among the binaries to be calibrated with some confidence, ζ Herculis is one of

												the most interesting owing to the difference of evolutionary state of components” ¶ high velocity relative to Sun
η	Her A	16 43.7+38 53	3.48	0.92	G7.5 IIIb Fe–1†	30.0	0.9	109	0.092	157	+8 V?	the outlying (116″, 84″) and faint (mag. 11.7, 13.9) celestial-sphere neighbours η Her B and η Her C aside, a close (0.3″) celestial-sphere neighbour of η Her A was suspected in 1842, without subsequent detection ¶ in evolutionary terms a resident of the “Red Clump” (fusing helium in stable core) ¶ Astron. Alm. (epoch 2021.5) assigns MK type G7 III Fe–1 ¶ Fe is notably underabundant
α	TrA A	16 51.1−69 04	1.91	1.45	K2 IIb–IIIa†	~8.4	−3.5	390	0.036	150	−3	**Atria** anomalous for its MK type, with flares and X-ray emission, perhaps from as-yet-undetected magnetically active companion (a companion would indeed be indicated by the claimed “barium star” status of α TrA; stars.astro.illinois.edu/sow/atria.html, in discussing the possibility of a companion, also remarks, however, “the classic ‘hybrid star,’ a giant that shows evidence for blowing a cool wind from its surface, yet having a hot surrounding magnetic corona at the same time”; ntrs.nasa.gov/archive/nasa/casi.ntrs.nasa.gov/20040086627.pdf further discusses both α TrA and β Dra, as (solitary) stars, which are in this particular posited sense “hybrid”; the faint (mag. 11.4) and outlying (angular separation 92″) celestial-sphere neighbour α TrA B cannot be the postulated flaring and X-ray-bright companion
ε	Sco	16 51.6−34 20	2.29	1.14	K2 III	51	0.8	64	0.666†	247	−3†	Larawag slow rotator (possibly even 1.3 y) ¶ evolved, and yet not a clump star; stars.astro.illinois.edu/sow/epssco.html discusses the uncertainty in evolutionary stage (brightening, with He core as yet awaiting ignition? dimming, with He core fusion in progress? or brightening, with dead C-and-O core, He-core fusion now over?) ¶ high velocity relative to Sun indicates origin outside the galactic thin disk (and metal underabundances are consistent with such an origin)
μ^1	Sco A	16 53.4−38 05	3.00v†	−0.20	B1.5 IVn	7	−2.9	500	0.024†	206	−25† SB2†	Xamidimura ecl.: 2.94–3.22, 1.4463 d (more precisely, in AAVSO(VSX) as viewed 2021 Dec. 14, 1.44626907 d); 2 published solutions for the orbit give, respectively, *e*=0.019, *e*=0.0; semidetached, partially eclipsing binary system, with mass transfer, resembling β Lyr in its never-constant light and in exhibiting both primary and secondary minima; 1948MNRAS.108..398S gives the light curve, and also discusses early observational history (this is the third eclipsing SB discovery in astronomy (made by Bailey, 1896)); distance between components is ~0.07 AU; since the SB is not as yet resolved, even interferometrically, WDS is not as yet able to write “μ^1 Sco Aa”, “μ^1 Sco Ab” ¶ μ^1 Sco has celestial-sphere neighbor μ^2 Sco, at mag. 3.6 (so just barely fainter than our chosen “Brightest Stars” magnitude cutoff), at angular separation ~5.8′ (Mizar-Alcor angular separation is just under 12′; normal naked-eye resolution is taken in ophthalmology, under at any rate some reasonable selection of consulting-room eye chart, to be ~1′); μ^1 Sco and μ^2 Sco are the “Little Cat’s Eyes”, as distinct from the “Cat’s Eyes” which are λ Sco and υ Sco; the IAU-official name

												“Xamidimura” applies just to the primary in the SB which is the right, unresolved, two-star μ^1 Sco system; for the Khoekhoem nomadic pastoralists of SW Africa, on the other hand, “Xami di mura” is “eyes of the lion”, as a designation for the naked-eye challenge pair μ^1 Sco, μ^2 Sco; μ^1 Sco and μ^2 Sco are not gravitationally bound, although both belong to the (gravitationally unbound) “Upper Sco” subgroup of the Sco-Cen Association ¶ μ^1 Sco and μ^2 Sco are not gravitationally bound, although both belong to the (gravitationally unbound) “Upper Sco” subgroup of the Sco-Cen Association ¶ a little confusingly, μ^2 Sco is also formally “μ^1 Sco H” (μ^1 Sco AH astrometry: 333″→347″, PA 71°→72°, 1752→2015); additionally, the unresolved μ^1 Sco A SB has two other WDS-documented celestial-sphere neighbours, both reasonably bright, μ^1 Sco B (mag. 8.9, 8.9″→9.2″, PA 211°→210°, 1999→2000) and μ^1 Sco G (mag. 9.4, 81″→81″, PA 257°→257°, 1935→2016)	
κ	Oph	16 58.7 +9 20	3.19†	1.16	K2 III	36	1.0	91	0.292†	268	−56 V†	slow rotator (possibly as slow as 1.6 y) ¶ historical assertion of variability may be due to a confusion between κ Oph and χ Oph; completely apart from this historical problem, however, 2001BaltA..10..593A discusses the possible variability both of κ Oph and of other Red Clump stars ¶ high velocity relative to Sun suggests origin outside the galactic thin disk	
ζ	Ara	17 00.5−56 01	3.12	1.55	K4 III	7	−2.7	490	0.041	206	−6	one of the rather rare instances of a giant excessively bright in far IR (1997A&A...323..513P suggests that such giants are more likely to be radiating their IR excess from circumstellar debris disks than from winds, and so are to be considered evolved-star analogues of the unevolved (and IR-bright) α Lyr)	
ζ	Dra A	17 08.9+65 41	3.17	−0.12	B6 III†	10	−1.8	330	0.028	314	−17 V	ζ Dra A and ζ Dra B are mags. 3.2, 4.2 respectively; a difficult binary, resolved interferometrically, but as of at any rate 2021 Dec. 14 with just 10 astrometry measurements (AB 0.0″→0.1″, 1981→1994); on the preliminary orbital solution offered in 1998A&AS..133..149M, the orbital plane coincides with the plane of the sky, and the orbit is circular, with radius 67 mas; 1998A&AS..133..149M suggests, tentatively, that ζ Dra A and ζ Dra B are “a pair of giants” ¶ given the recent formation of the ζ Dra system, Fe is anomalously underabundant ¶ E(B–V) =+0.03	Aldhibah
η	Oph AB†	17 11.7−15 45	2.43	0.06	A1 IV + A1 IV?†	37	0.3	90	~0.107	~22	−1 SB	A: 3.0; B: 3.3, A3 V, 0.5″ (2019), orbit 87.6 y highly eccentric orbit: separation 2 AU min, 65 AU max ¶ under IAU rules, “Sabik” designates η Oph A, not η Oph B ¶ our present assignment of MK types (confident for η Oph A, tentative for η Oph B) is from the literature; our Handbook predecessor R.F. Garrison, however, himself favoured “A2.5 Va,” perhaps for the AB composite; Astron. Alm. (epoch 2021.5), perhaps again for the AB composite, assigns MK type “A2 Va$^+$ (Sr)” ¶ it is possible that A, or B, or both A and B, are superabundant in metals	**Sabik**†
η	Sco A	17 13.8−43 16	3.32	0.44	F5 IV†	~44.4	1.6	73	0.290	175	−28	we now take MK type (slightly evolved beyond stable core-hydrogen fusion, because in luminosity class “IV”) from NASA NStars work summarized at 2006AJ....132..161G (with Garrison the third author); Garrison had himself previously, in this Handbook table, proposed the dwarf MK type “F2 V:p(Cr)"; the intricacy of	

												his previous type hints at difficulties in classification, and indeed even “dwarf barium star” has been asserted elsewhere; Astron. Alm. (epoch 2021.5) assigns the same MK type as legacy-Garrison ¶ rapid rotator (< 1 d); the observed X-ray emission is consistent with magnetic effects (including coronal heating?) stemming from rapid rotation
α	Her Aa†	17 15.7+14 22	2.78†	1.16	M5 Ib–II	9	−2.4	400	0.032	347	−33 V	semireg. var.: 2.73–3.60 in V; B: 5.4, 5″ (2019) Rasalgethi† the second-nearest AGB star (the nearest being the more dramatic visual variable o Ceta Aa (Mira) ¶ WDS catalogues α Her Aa,Ab, but with the caveat that the asserted duplicity may not be real (only 3 measurements are available (0.2″→0.2″, 1986→1991), and attempts to resolve the asserted binary failed over the period 1985→1997, even with speckle interferometry at BTA-6; a period of ~10 y has been suspected if the Aa,Ab pairing is real); the IAU-official name Rasalgethi applies to α Her Aa if the Aa,Ab pairing is real, and to α Her A otherwise; although we use WDS nomenclature (for this as for all stars in this table), the literature also (e.g. 1993A&A...274..838T, 2013AJ....146..148M) uses an alternative terminology, on which Rasalgethi is designated “α^1 Her A”, and the two as-yet-unresolved SB components are designated “α^2 Her A” and “α^2 Her B”; the as-yet-unresolved SB (MK types G8 III and A9 IV–V, period 51.578 d, distance between the unresolved components ~0.4 AU) which both for WDS and for us is designated α Her B is mag. 5.4 (9″→5″, PA 117°→103°, 1777→2020); AB orbital solution asserts period 3600 y; this makes the α Her system at least a (kinematically stable, hierarchically organized) triple; the faint celestial-sphere neighbours α Her C (mag. 15.5) and α Her D (mag. 11.1), on the other hand, are not part of the system ¶ the mass of Rasalgethi has been controverted (mass as high as ~15 $M_{\odot}$, putting Rasalgethi into the same mass league as (admittedly less evolved) Betelgeuse and Aldebaran; or, rather, as low as ~2 $M_{\odot}$, putting Rasalgethi into the same mass league as its AGB colleague Mira?); 2013AJ....146..148M, proceeding from the surely safe assumption that the three known stars in the α Her AB system are of the same age (being surely born in a single ISM cloud condensation event), and taking the system age from the MS star which is the secondary in the α Her unresolved SB (age is more safely determined from an MS star than from an evolved star), and using photometry sensitive to TiO to obtain a fluctuating effective temperature for Rasalgethi (and thereby, via the Stefan-Boltzmann law, fluctuating luminosity and fluctuating radius) via a 55-track grid of evolutionary models (with each track tracing evolution stepwise until the depletion of core helium which marks a star’s arrival on AGB) deduces that the Rasalgethi mass is in the range [2.175 $M_{\odot}$, 3.250 $M_{\odot}$]; the authors remarks that “very few AGB stars have reliable ages and masses”; the luminosity and effective temperature deduced for Rasalgethi agree to within uncertainties with the Rasalgethi interferometry results of 2004A&A...418..675P ¶ in the classification of AAVSO(VSX), Rasalgethi is a “semi-regular late-type giant”, with V-mag. range 2.73−3.60; 2010Ap&SS.328..113M reports “up to seven” pulsation modes, with one period of ~1343 d (a radial pulsation), other periods on the order of ~125 d (and cf. also light curve in Fig. 5 of 2001PASP..113..983P); two typical measurements in V band, from AAVSO database, from one and the same observer, are mag. 3.37 (2021 June 27) and mag. 3.07 (2021 Sep. 29) ¶ Rasalgethi radius variation, associated with the photometric variation, is assessed in 2013AJ....146..148M as 264 $R_{\odot}$ min, 303 $R_{\odot}$ max; en.wikipedia.org/wiki/List_of_largest_stars shows ranking of Rasalgethi in the overall known cosmic population of highly evolved stars

												¶ 1956ApJ...123..210D initiated the successful analysis of mass loss from Rasalgethi (mass loss is to be expected, in an AGB star, and in the case the consequent circumstellar material is so copious as to extend into our line of sight to α Her B); 1956ApJ...123..210D proceeds by examining the radial velocity of the Rasalgethi ejecta (as seen in absorption, as very cold gas, in the spectrum of the α Her B brighter component, and (crucially) not sharing in the ~52 d fast SB orbital motion of that component); 2007ApJ...658L.103T finds interferometric evidence for an exceptionally violent episode of mass loss, comprising about 10e−6 $M_\odot$ (cp. Mira, for which we take as usual mass-loss rate ~2.5e−7 $M_\odot$ / y), with the ejected material (condensing as dust) flowing outward at exceptionally high speeds, possibly ≥ 72 km/s, and with no similarly drastic mass loss until at least the conclusion of the study (~2003); this suggests to us, as the Handbook team, the advisability of ongoing photometric surveillance ¶ in the pictorial-atlas tradition, α Her marks the head of hero Hercules (with β Her marking his club; for summer-evening observers in the northern hemisphere, the hero is to be visualized inverted, with feet high in the sky, club and head lower; indeed the IAU-official Arabic name "Rasalgethi" derives from the Arabic for "the kneeler's head"
π	Her	17 15.8+36 47	3.16†	1.44	K3 IIab†	8.7	−2.2	380	0.027	276	−26 V?	¶ Astron. Alm. (epoch 2021.5) assigns MK type K3 II ¶ low-amplitude photometric variations with low-amplitude radial-velocity variations, 613 d, perhaps favour the hypothesis of non-radial pulsation over the competing hypotheses of an undetected low-mass companion and of rotation with starspots
δ	Her Aa†	17 16.0+24 49	3.12	0.08	A1 Vann	43.4	1.3	75	~0.158	~188	−40 SB†	B: mag. 8.3, 14″ (2020) is mere optical companion Sarin δ Her A, being SB (and also resolved as a binary in interferometry, with angular separation 60 mas; inter-component distance ≥ 1.45 AU, period ≥ 335 d; WDS documents just 5 astrometry measurements (0.1″→0.1″, 1978→1989)), is strictly δ Her Aa,Ab ¶ δ Her Aa is a fast rotator (< 9 h) ¶ as with δ Her B, so also δ Her C and δ Her D, at respective angular separations 174″ (2013) and 192″ (2009), are most likely mere optical companions
θ	Oph A†	17 23.4−25 01	3.27†	−0.19	B2 IV	~7.5	−2.4	440	0.025	197	−2 SB†	slight var., β Cep type, 3.26–3.29 in V, 0.14 d θ Oph A is classified as an unresolved SB, with period variously stated in secondary literature as 11.44 d (Kaler) and 56.71 d (Wikipedia), and with inter-component distance proposed by Kaler as ~0.25 AU; θ Oph B is mag. 6.2, 0.2″→0.1″, 1992→2020, with 10 astrometry measurements documented in WDS; lunar occultations occur, and indeed a lunar occultation event might possibly (Kaler, stars.astro.illinois.edu/sow/thetaoph.html) have split the SB which is θ Oph A ¶ according to AAVSO(VSX) as viewed 2022 March 3, the unresolved θ Oph A pairing presents slight variability of both the β Cep type and "slowly pulsating B" type, with period 0.1405278d
β	Ara	17 27.2−55 33	2.84	1.48	K3 Ib–IIa†	5	−3.6	600	0.027	199	0	slow rotator (possibly as much as 2.33 y) ¶ high metallicity ¶ not gravitationally bound to γ Ara AB
γ	Ara A	17 27.3−56 24	3.31	−0.15	B1 Ib	~2.9	−4.4	1100	0.016	182	−3 V	broad lines for Ib ¶ γ Ara A is rapid rotator (both "~4.8 d" and "< 2.5 d" have been asserted, and yet rapid rotation is unusual for the (evolved) γ Ara A luminosity class) ¶ 1997A&A...318..157P finds via *IUE* spectroscopy that, consistently with this rapid rotation, the stellar wind of γ Ara A

												may be equatorially enhanced (and more generally, that the wind is variable, and is structured with two components, its structure being not typical of stars in this portion of the HR diagram) ¶ γ Ara AB is not gravitationally bound to β Ara; γ Ara B is faint (mag. 10.21: 18″→18″, PA 324°→326°, 1835→2016) ¶ E(B–V) =+0.08	
β	Dra A	17 30.9+52 17	2.79†	0.95†	G2 Ib–IIa	8.6	−2.5	380	0.020	308	−20 V	in evolutionary terms, β Dra A is somewhat unusual, as a yellow more-than-giant (having been a stable core-hydrogen fuser just 0.5 My ago, the star is in transition to being redder, and of still larger radius) ¶ it is also odd that β Dra A, while lying in the HR diagram IS, has not been observed to pulsate	Rastaban
υ	Sco	17 32.3−37 19	2.70	−0.18	B2 IV†	6	−3.5	600	0.030	185	+8 SB	although we here give spectral type B, type Be has also been asserted ¶ υ Sco and λ Sco are not gravitationally bound (although both belong to the (gravitationally unbound) Sco-Cen OB association, and have as an optical double been called the "Cat's Eyes") ¶ E(B–V) =+0.02	Lesath
α	Ara A	17 33.6−49 53	2.84	−0.14	B2 Vne†	12†	−1.7	300†	0.075	206	0 SB	an instance of the "Be phenomenon," with (since the star, with its equatorial ejecta, is seen nearly equator-on) "shell" spectrum: 2007A&A...464...59M says, "For the first time, we obtain the clear evidence that the [equatorial ejecta] disk is in Keplerian rotation, closing a debate that has continued since the discovery of the first Be star γ Cas by Father Secchi"; on the authors' modelling, α Ara is rotating near breakup speed (and consequently is oblate), with an enhanced wind from its poles; the authors note the possibility that equatorial ejecta disk is truncated by an unseen companion at 32 stellar radii ¶ the SB which is α Ara is not as yet resolved, even in interferometry (so WDS cannot as yet write "α Ara Aa", "α Ara Ab") ¶ IR excess is unusually high for a Be star ¶ for problem of distance (the *HIPPARCOS* distance given here may be too high) cf 2005A&A...435..275C and 2007A&A...464...59M	
λ	Sco Aa,Ab†	17 35.1−37 07	1.62†	−0.23	B1.5 IV + n.a.	>6†	−4.6	400†	0.032	195	−3 SB2	slight ecl. var., 1.59−1.65 in V, 0.2137 d λ Sco Aa,Ab are respectively of mags. 2.1, 2.7, and are an instance of an interferometrically resolved SB2 (38 measurements, 1999→2019; cf 2006MNRAS.370..884T); period of the SB2 is 5.9525 d; strictly a hierarchical triple system, however, in which the wide λ Sco AB pairing has period ~1000 d; B is elusive, at mag. 14.9; there is additionally a less elusive celestial-sphere neighbour, λ Sco C, at mag. 9.2, at a rather wide angular separation from A (AC: 95″→94″, PA 331°→330°, 1897→2016) ¶ the IAU-official name "Shaula" applies not to the λ Sco Aa,Ab system, but only to λ Sco Aa ¶ we take here the *D* and π values implied by comparing the Aa,Ab orbit semi-major axis angular measure against the computed Aa,Ab orbit physical semi-major axis length rather than the π value measured by *HIPPARCOS* and its corresponding *D*; generally speaking, *HIPPARCOS*, like other fine-grained trigonometric-parallax determinations of distance, risks degradation through astrometric wobble if a star has a stellar-mass gravitationally bound companion ¶ although λ Sco Aa,Ab has a published orbit, λ Sco AB does not (and indeed WDS takes no position on the	**Shaula**

												question whether this wide double is a true binary) ¶ λ Sco Aa is a β Cep variable; since full orbital coverage is available in this case (as also with β Cep itself; in most or all other β Cep-class cases, full orbital coverage is presently unavailable), mass determination becomes feasible, making the λ Sco Aa,Ab binary important in β Cep-variable research; λ Sco Ab is itself of interest, as a possible pre-main-sequence star (this would be consistent with the observed X-ray emission) ¶ 1975MNRAS.173..709L gives some photometry ¶ a flare was observed in vicinity of λ Sco on 1975 Jun. 01 ¶ 2004A&A...427..581U summarizes previous work on λ Sco, considers masses, and discusses tidal effect on β Cep pulsation ¶ λ Sco and υ Sco are not gravitationally bound (although both belong to the (gravitationally unbound) Sco-Cen OB association, and have as an optical double been jointly called the "Cat's Eyes") ¶ E(B–V) =+0.03	
α	Oph <u>A</u>†	17 36.0+12 33	2.08†	0.16	A5 Vnn	67	1.2	49	0.247	154	+13 SB	slight var., δ Sct type & γ Dor type, ~0.006 in V AAVSO(VSX), as viewed 2022 Feb. 21, gives period 0.05357 d ¶ α Oph A is a fast rotator (oblateness has been imaged interferometrically), seen nearly equator-on ¶ B is mag. 5.0 ¶ the SB system α Oph Ab, although tight, has been observed interferometrically, with 17 measurements 1982→2018; the binary system has become better understood with the recent, 2011ApJ...726..104H, determination of masses and orbit geometry, through coronagraph and adoptive optics (period 3148.4 d, angular distance at periastron passage ~50 mas; the now-achieved determination of masses in this particular system has implications for astrophysics generally, since it potentially facilitates the refining of numerical models for rapidly rotating hot stars) ¶ asteroseismology mission MOST has identified ~50 pulsational modes in α Oph A	**Rasalhague**
ξ	Ser Aa†	17 38.9−15 25	3.54†	0.26	F0 IIIb†	31	1.0	105	0.073	215	−43 SB	hierarchically organized triple system, comprising the resolved (with just one measurement, in 1987: 0.3″) single-lined SB which is ξ Ser Aa and ξ Ser Ab, experienced as essentially a point mass by the outlying ξ Ser B; period of the single-lined SB Aa,Ab is 2.29 d; ξ Ser B is faint, at mag. 13.0, with AB period possibly ~15,000 y; there is additionally a faint celestial-sphere neighbour ξ Ser C, at mag. 13.8 ¶ ξ Ser Aa has been asserted to be very slightly hotter than Garrison's "F0 IIIb", and moreover to be chemically peculiar, being on this (more recent?) determination of MK type A9 IIIp Sr; additionally, somewhere in the literature, δ Sct variability has been asserted or conjectured (Kaler comments at stars.astro.illinois.edu/sow/xiser.html: "the star /…/ remains cryptic") **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	
θ	Sco A†	17 38.9−43 01	2.0	0.41	F1 III	~11	−3.0	300	0.006	119	+1	rapid rotator, in the sense that *v* sin *i* is (according to 2005yCat.3244....0G) 125.0 km/s; since, however, θ Sco is a (rapidly evolving) giant, its high *v* sin *i* may correspond to a not-spectacularly short rotation period, of up to 10 d; if, as asserted in literature, θ Sco A truly is a rapid rotator, it will resemble	Sargas†

												β Cas A (a rapid rotator that, strikingly, has evolved beyond the MS) ¶ θ Sco B is catalogued by WDS as mag. 5.4, with θ Sco AB 6″→6″, PA 322°→315°, 1896→1991, but with the caveat (in WDS Notes) that θ Sco B information may be erroneous (grid-step error in *HIPPARCOS* data reduction? erroneous work in 1896?) ¶ although it is the Sumerian name Sargas that is IAU-official as of 2016 Aug. 21, θ Sco, like κ Sco, has also been known under the different, not IAU-official, name Girtab (originally applied by the Sumerians to an entire asterism)	
κ	Sco†	17 44.0−39 02	2.41†	−0.17	B1.5 III	7	−3.5	480	0.026	193	−14 SB2	slight var., β Cep type, 2.41–2.42 in V, ~0.2 d ¶ since the SB has not yet been resolved, even interferometrically, WDS cannot yet write “κ Sco A” and “κ Sco B”; the SB has orbital period 195.65 d, with inter-component distance 1.7 AU ¶ the SB primary is a rapid rotator (1.9 d), in addition to being a variable of β Cep type (0.1998303 d in AAVSO(VSX) as viewed 2021 Jan. 16 and 2022 March 3; 1975MNRAS.173..709L gives some photometry, confirming a beat period) ¶ κ Sco, as a single naked-eye object, has (like θ Sco) been known under the different, not IAU-official, name Girtab (originally applied by the Sumerians to an entire asterism)	
β	Oph	17 44.6 +4 34	2.76†	1.17	K2 III†	~39.8	0.8	82	0.165	345	−12 V	¶ Astron. Alm. (epoch 2021.5) assigns MK type K2 III CN 0.5 ¶ 1996ApJ...468..391H finds multiple pulsation periods, in behaviour paralleling α Boo (“it may well be that these [two] stars represent a new class of radially and unradially pulsating stars”), and also a possible long period of 142.3 d; the authors suggest that if the latter is real, then although the more likely explanation is a 142.3-d rotation, nevertheless gravitational pull from an unknown exoplanet is conceivable	Cebalrai
μ	Her Aa†	17 47.3+27 43	3.42	0.75	G5 IV	~120.3	3.8	27.1	0.804	201	−16 V	BC: 9.78comb., A−BC 36″, PA 240°→249°, 1781→2015 this stable “double double” is the third-closest quadruple star system to the Sun, and is one of the best-studied double doubles (2016AJ....151..169R consequently writes that μ Her “serves as an archetype for understanding stellar system formation”): μ Her Aa,Ab is in tight orbit, and μ Her BC is in tight orbit, with each of these two pairs experiencing the other as essentially a point mass; 2016AJ....151..169R gives Aa,Ab a period of ~100 y, with wide uncertainties, concluding also that Ab is an M-dwarf (and thus more massive than a mere substellar object; Aa is mag. 3.5, Ab mag. 12.7, with this magnitude difference making the astrometry difficult; the Aa,Ab binary has been measured 24 times, 1998→2015 (angular separation 1.8″ in 2015); BC has period 43.127 y, B-to-C distance 1.5 AU min, 3.6 AU max, 2.2 AU average (angular separation 0.6″ in 2016), abundantly observed (360 measurements, 1854→2016), with B, C of respective mags. 10.2, 10.7; the distance between the Aa,Ab centre of mass and the BC centre of mass is ≥ 300 AU, with orbital period ≥ 3700 y; in contrast with the two pairings Aa,Ab and BC, there is no published orbital solution for the wide and slow pairing AB ¶ despite having finished core-hydrogen fusion, μ Her Aa is a fast rotator, and is consequently magnetically active and an X-ray source	

ι^{1}	Sco A	17 49.2−40 08	2.99	0.51†	F2 Ia	2†	−5.9	2000†	0.006	180	−28 SB	since the SB which is ι^{1} Sco A is not as yet resolved (even by interferometry), WDS is not as yet able to write “ι^{1} Sco Aa”, ι^{1} Sco Ab” ¶ a rare instance of a yellow supergiant (dead helium core; the star is now cooling, and is now in transition to the less exotic status of red supergiant) ¶ radius estimates vary; CADARS (2001A&A...367..521P) value is ~1.9 AU ¶ mass loss ~1e–7 $M_{\odot}$/y ¶ slow rotator (≥ 0.5 y) ¶ distance and mass are rather uncertain ¶ the modest angular distance of ι^{1} Sco from ι^{2} Sco is the result of a mere line-of-sight coincidence (with ι^{2} ~2 times as distant as ι^{1}; again by coincidence, not ι^{1} alone, but also ι^{2}, is a supergiant)
G	Sco A	17 51.4−37 03	3.19	1.19	K2 III	25.9	0.3	126	0.049	56	+25	HR6630, Fuyue although masses of K giants are in general uncertain, in this particular case the mass is known via WIRE salvage-mission asteroseismology (being determined in 2008ApJ...674L..53S as 1.44 $M_{\odot}$, with just a 15% uncertainty)
γ	Dra A	17 57.1+51 29	2.24	1.52	K5 III†	21.1	−1.1	154	0.024	200	−28	**Eltanin** in 1728, James Bradley used γ Dra to demonstrate aberration of light (“velocity aberration”); his demonstration strongly confirmed the heliocentric (and thus non-Ptolemaic) kinematics of the Solar System ¶ Fe is slightly underabundant
ν	Oph +2P†	18 00.3 −9 46	3.32	0.99	G9.5 IIIa†	22	0.0	150	0.117	185	+13	brown-dwarf companions, not optically resolved (so WDS cannot write “ν Oph A”, “ν Oph B”, “ν Oph C”) with masses ≤ 24× Jupiter and ≤ 27× Jupiter (deuterium fusion begins at a lower mass, 13× Jupiter), periods 530.3 d and 3190 d (Quirrenbach et al. 2011, and additionally 2012PASJ...64..135S; the latter paper suggests formation in circumstellar disk, with subsequent migration, in a scenario reminiscent of planet and exoplanet formation): this is the third star found to be hosting two brown dwarfs ¶ slow rotator (≤ 234 d) ¶ far-IR variability has been suspected ¶ CN underabundant, Fe overabundant ¶ Astron. Alm. (epoch 2021.5) assigns MK type G9 IIIa
γ^{2}	Sgr	18 07.3−30 25	2.98	0.98	K0 III†	34	0.6	97	0.189	197	+22 SB	Alnasl metals underabundant ¶ Astron. Alm. (epoch 2021.5) assigns MK type K0^{+} III ¶ ε Sgr and the γ^{2}–γ^{1} Sgr pair serve as pointers to Baade’s Window ¶ angular proximity of γ^{1} Sgr (= W Sgr; variable, mag. range 4.28−5.10 in V; ~50′, to ~N of γ^{2} Sgr) is a mere line-of-sight coincidence
η	Sgr A†	18 19.2−36 45	3.10†	1.5	M3.5 IIIab	22	−0.2	~146	0.211	218	+1 V?	slight irreg. var.: 3.05–3.12; B: 8.00, G8: IV:, 3.5″ (2016) PA 100°→110°, 1879→2016; orbit ≥ 1270 y, A-to-B distance ≥ 165 AU ¶ η Sgr A is variously asserted to be on the (very highly evolved) HR diagram AGB or at the tip of the RGB ¶ η Sgr A is in the AAVSO(VSX) classification an “LB,” i.e. a slow irregular variable ¶ temperature of η Sgr A not yet well determined?
δ	Sgr A	18 22.4−29 49	2.72	1.38	K2.5 IIIa†	9	−2.4	350	0.041	128	−20 V?	Kaus Media† ¶ Astron. Alm. (epoch 2021.5) assigns MK type K2.5 IIIa CN 0.5 ¶ possibly a weak barium (Ba) star, δ Sgr A possesses (as expected for a Ba star) a WD companion ¶ temperature of δ Sgr A not yet well determined? ¶ “Kaus” is Arabic “bow,” with Kaus Borealis (λ Sgr), Kaus Media (δ Sgr), and Kaus Australis (ε Sgr) the three delineating stars of

												the archer's bow; by coincidence, the archer turns out to be aiming rather close both to Baade's Window and (prolonging the line of firing) to the Sgr A* black hole at the galaxy's centre
η	Ser A	18 22.5 −2 53	3.23	0.94	K0 III–IV†	54	1.9	~60.5	0.890	218	+9 V?	slow rotator (but ≤ 1.9 y) ¶ high velocity relative to Sun suggests that η Ser is an interloper (born outside the galactic thin disk? consistently with this conjecture, Fe is underabundant)
ε	Sgr A	18 25.7−34 22	1.79	−0.03	A0 IIn (shell?) †	23	−1.4	~143	0.130	198	−15	**Kaus Australis**† fast rotator (consistent with shell-star classification); as might be predicted for a fast rotator, a magnetic field, and also X-ray emission, have been detected ¶ Astron. Alm. (epoch 2021.5) assigns MK type "A0 II⁻n (shell)" ¶ has been classified as a λ Boo star, apparently in error ¶ IR excess indicates debris disk (possibly also detected in polarimetry), at average separation 155 AU; and yet a companion (other than the WDS-documented celestial-sphere neighbours ε Sgr B, ε Sgr C, ε Sgr D) is also asserted, surprisingly present within this debris-disk radius ¶ ε Sgr B,C,D are at mags. 14.3, 8.4, 9.0 respectively; AC astrometry is 2.2″→2.4″, PA 146°→142°, 1992→1999; AD astrometry (as a pair at wide angular separation) is 858″→858″, PA 36°→36°, 1980→1999
α	Tel	18 28.6−45 57	3.49†	−0.18	B3 IV†	12	−1.2	280	0.056	198	0 V?	stars.astro.illinois.edu/sow/alphatel.html remarks that MK luminosity class IV notwithstanding, α Tel is still on the astrophysical (as opposed to the MK-phenomenological) main sequence (in other words, is still fusing core hydrogen) ¶ said in 2005ApJS..158..193S to be among the (rare) He-rich stars; these authors list α Tel as a candidate-and- unconfirmed β Cep variable, and say they suspect it is a variable in the slowly pulsating B-star class; although α Tel has *HIPPARCOS* microvariability (0.909 d), it is absent from the AAVSO(VSX) database as viewed 2022 March 3 **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**
λ	Sgr A	18 29.4−25 24	2.82	1.02	K1 IIIb	~41.7	0.9	78	0.191	194	−43 V?	Kaus Borealis† modest X-ray emission indicates some magnetic activity (not usual in a duly evolved, stable core-He-fusing, HR diagram "clump star") ¶ λ Sgr B is mag. 9.9; AB 101″→82″, PA 188°→184°, 1911→2019 ¶ lunar occultations are possible, planetary occultations possible yet rare; most recent planetary occultation was by Venus, on 1984 Nov. 19 ¶ unusual in occupying fully three roles in the Western pictorial traditions: as northernmost star of the Archer's Bow, as westernmost (handle-tip) star of the Little Milk Dipper, and as uppermost (lid-knob) star in the Teapot
α	Lyr A	18 37.7+38 48	0.03†	0.00	A0 Va†	130	0.6	25.0	0.350	35	−14 V	pole-on rapid rotator with circumstellar disk **Vega** ¶ pole-on rotators are useful for asteroseismology, since all but the axisymmetric modes (whether radially symmetric or radially not symmetric) are helpfully rendered invisible to photometry (and in a rather analogous way, equator-on rotators are also useful through suppression of some modes); α Lyr pole-on orientation represents an extreme on a continuum whose other extreme is represented by the equator-on rapid rotator α Leo A (Regulus); the α Lyr A

rapid rotation (2015A&A...577A..64B
now confidently asserts 0.68 d)
yields oblate spheroid shape
(here as with α Leo A);
it is this, with
consequent latitude-varying
photosphere (severe temperature
and luminosity gradients along
the arcs of photospheric longitude,
with equator coolest and darkest),
rather than any evolution
beyond core-hydrogen-fusion stage, that
explains the anomalously
high luminosity (α Lyr A
is in MK luminosity class Va,
rather than in the slightly dimmer V class
that would be observed if
its orientation was equator-on)
¶ α Lyr A is now known to harbour
all three of the classical
circumstellar-dust regimes (~1500 K, near-IR;
~120-170 K, mid-IR, as an
analogue of our own zodiacal dust; and
~50 K, far-IR, as an analogue
of our own Kuiper Belt:
for regimes overview without
specific reference to α Lyr,
cf 2013ApJ...763..118S, section 1):
2013ApJ...763..118S is the paper
announcing discovery of the second
of these around α Lyr
(with sections 5.1 and 5.2,
respectively, summarizing
previous α Lyr A work on the
first and third of the three regimes):
a question of recent interest
is the origin of the α Lyr A
exozodiacal (warm-regime, mid-IR)
dust (episode analogous to our own
planetary system's Late Heavy
Bombardment? or, rather,
some steady-state replenishment mechanism?);
efforts at detecting exoplanet(s)
to account for the complex
inferred, and indeed in some
wavelengths also now directly imaged,
disk structure have not yet succeeded
¶ 2007ASPC..364...305G,
reviewing the history of α Lyr A
photometry, considers modest variability likely,
the historical use of α Lyr A as a photometric
standard notwithstanding
(and indeed α Lyr A
is described at AAVSO(VSX) as
a low-amplitude δ Sct variable,
in the now-obsolete AAVSO(VSX) "DSCTC"
classification bin, fluctuating between
−0.02 and +0.07 in V, with period 0.19 d)
¶ 2010A&A...523A..41P,
with 2014A&A...568C...2P corrigendum,
is a recent discussion of α Lyr A magnetism
(the authors note that α Lyr A
"may well be the first confirmed
member of a much larger, as yet
unexplored, class of weakly-magnetic
stars now investigatable with
the current generation of stellar
spectropolarimeters"; for origin, they
somewhat favour dynamo over fossil,
and radiative dynamo over core dynamo):
consistently with magnetism,
2015A&A...577A..64B finds,
via line-profile variations,

												multiple (bright, not dark) star spots, in some undetermined complex pattern (authors comment that this is “first strong evidence that standard A-type stars can show surface structure”); 2015A&A...577A..64B is additionally one of several papers summarizing recent work on an interrelated complex of α Lyr A themes, comprising (in addition to magnetism) rotation, spots, photovariability, and pulsation ¶ in 2010ApJ...725.2401F, Fig. 8 with its accompanying discussion summarizes studies on elemental abundances (important because α Lyr A, as a rather “normal” star for MK temperature type A, might serve as a benchmark for appraising chemically peculiar A stars) ¶ WDS documents an exceptionally large number of celestial-sphere neighbours for α Lyr A (B,C,D,E,F,G,H, Ia,Ib,J,Ka,Kb,L,M,N,O,P,Q,R,S,T); all of these are faint apart from B and E (each at mag. 9.5, and each known to be not gravitationally bound to A): AB 37″→84″, PA 117°→184°, 1781→2017; AE 150″→87″, PA 40°→39°, 1831→2017 ¶ E(B–V) =0.00	
φ	Sgr†	18 47.1−26 58	3.17	−0.11	B8 III	14	−1.2	240	0.051	89	+22	apparent duplicity now discounted (erroneous lunar-occultation observation) ¶ stars.astro.illinois.edu/sow/phisgr.html discusses some difficulties in physical modelling (if pole-on rotator, then there will be troublesome temperature and luminosity gradients along the arcs of photospheric longitude)	
β	Lyr Aa1†	18 50.9+33 23	3.52v†	0.00	B7 Vpe (shell) †	~3.4†	−3.8	~960†	0.004	152	−19 SB	eclipsing, 3.30–4.35 in V, 13 d period is increasing at constant rate of ~19 s/y; orbital solution has been published for β Lyr Aa1, β Lyr Aa2 (although not for the wider binary β Lyr Aa, β Lyr Ab, and not for the still wider binary β Lyr A, β Lyr B); this orbit is seen nearly edge-on; prototype of the β Lyr class of eclipsing systems (but has also been assigned to the new class of “W Ser stars”: 1980IAUS...88..251P); AAVSO supplies information both via VSX database (showing, e.g. the high-precision recent determinations of period) and via www.aavso.org/vsots_betalyr (a detailed astrophysics discussion, with bibliography): alternating deep and shallow visible-light minima, with the object eclipsed in the deep minima (the “donor”) a Roche-lobe-filling giant, currently ~3 $M_{\odot}$ and diminishing, and the object eclipsed in the shallow minima (the “gainer”) embedded in a thick accretion disk, currently ~13 $M_{\odot}$ and increasing; mass transfer is copious (~2e–5 $M_{\odot}$/y); this disk renders the gainer dim, and its eclipses consequently shallow, even though the (presently dim) gainer is (now, at this rather late stage in mass transfer) already ~4 times more massive than the (bright) donor (cf 1963ApJ...138..342H); further, instabilities in the accretion disk, from which ~20% of the light comes, make the light curve liable to vary slightly from cycle to cycle; the presently dim gainer is destined to be first (1) brightening, and spun up by conservation of angular momentum, as its obscuring accretion disk disappears by being dumped down into photosphere, and then (2) to become a slower rotator,	Sheliak

												tidally locked with the secondary; at stage “(1)”, the system will be a so-called “Rapidly Rotating Algol,” at stage “(2)”, on the other hand, the system will be simply a “classical Algol” ¶ 2008ApJ...684L..95Z presents the first (CHARA-interferometric) binary-resolving imaging, achieving resolution ~0.5 mas or ~0.7 mas (and for the first time in astrophysics deduces a β Lyr Aa1,Aa2 astrometric orbit); the bright low-mass donor, and the presently dim high-mass gainer, are evident, corroborating the overall conception of 1963ApJ...138..342H ; 2008ApJ...684L..95Z discusses also polar outflow jets on the gainer (these do not alter the essential situation: for the gainer, equatorial gain exceeds polar loss), and deduces a distance to ±15% (a distance consistent-to-within-uncertainties with the *HIPPARCOS* distance) ¶ 2012ApJ...750...59L discusses possible hot spot at edge of accretion disk, on the basis of spectropolarimetry (and ` has modelling that provides for hot spot, and additionally for a bright spot, on the accretion disk) ¶ some observations have been made in radio and (a regime especially relevant to hot-spot studies) X-ray ¶ strictly speaking, this is a hierarchical system, Ab experiencing the binary that is Aa1+Aa2 as essentially a point mass; for the Aa+Ab pairing, and for possibility of further pairings (AB, AC, … , Be, …), cf WDS and (a source that reports inter alia *Ga*ia) en.wikipedia.org/wiki/Beta_Lyrae ¶ although we here, following Garrison, assign a rather straightforward spectral type, this should be taken only as a starting point: cf, eg., 2000A&A...353.1009B, which lists six systems of spectral lines, while repeating an old O. Struve warning that spectrum involves circumstellar matter ¶ Kaler comments in stars.astro.illinois.edu/sow/sheliak.html “one of the most confusing, heavily studied, and important stars of the nighttime sky” ¶ the rather long period, with the large magnitude swing, and the readily discoverable difference in depths of the alternating minima, make this object a suitable binoculars-or-naked-eye photometry project (using γ Lyr A as a comparison) even from locations suffering rather frequent cloud **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**
σ	Sgr Aa	18 56.7−26 16	2.05	−0.13	B3 IV	14	−2.2	230	0.056	164	−11V	**Nunki** fast rotator ¶ the σ Sgr Aa,Ab duplicity was discovered, with suggestion that mags. are roughly equal, but with the system not resolved, through the Narrabri intensity interferometer (1974MNRAS.167..121H); WDS documents a single measurement of Aa,Ab (in 1991, with separation found to be ~12 mas), through aperture-masking interferometry at AAT (1994A&A...290..340B); σ Sgr B, at mag. 9.95, is known to be not gravitationally bound to σ Sgr Aa,Ab; AB 309″→349″, PA 244°→239°, 1837→1999 ¶ lunar occultations are possible, and planetary occultation possible-yet-rare (most recently Venus, 1981 Nov. 17) ¶ E(B–V) =+0.02
ξ^2	Sgr	18 59.1−21 05	3.52	1.15	K1 III	9	−1.7	400	0.034	113	−20	occultations (at any rate lunar) are possible ¶ the angular proximity of ξ^1 Sgr is a mere line-of-sight coincidence **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**
γ	Lyr A	18 59.8+32 43	3.25	−0.05	B9 II†	5	−3.1	600	0.003	290	−21 V†	Sulafat has been both asserted and denied to be SB ¶ 2001A&A...371.1078A reports many metals underabundant
ζ	Sgr AB†	19 04.0−29 51	2.60	0.06	A2 IV–V + A4:V:	37	0.4	90	n.a.	n.a.	+22 SB	A: 3.3; B: 3.5, 0.3″ (2020), 21.1 y, compos. spectrum Ascella

												PA is 223° in 2020; the system is well observed, with 210 astrometry measurements 1867→2020; A-to-B distance is 10.6 AU min, 16.1 AU max, average 13.4 AU ¶ under IAU rules, "Ascella" designates ζ Sgr A, not ζ Sgr B ¶ stars.astro.illinois.edu/sow/ascella.html discusses uncertainty in masses, remarks that temperatures are not yet directly measured ¶ Sgr C (17.6″ in 2013) is probably a mere optical companion	
ζ	Aql A†	19 06.4+13 54	2.99	0.01	A0 Vann	~39.3	1.0	83	0.096	184	−25 SB	among the most rapidly rotating stars known (period 16 h) ¶ since the SB is not as yet resolved (not even interferometrically), WDS is as yet unable to write "ζ Aql Aa", "ζ Aql Ab"; in the faint angular-proximity grouping ζ Aql B,C, D, E, ζ Aql B, at mag. 12.0, has been asserted to be a gravitationally bound companion of the ζ Aql A SB pair (5″→7″, PA 60°→46°, 1874→2016; A-to-B distance is ≥ 125 AU, period ≥ 800 y); also, the exceedingly faint (mag. 16.2) ζ Aql E may be gravitationally bound ¶ 2008A&A...487.1041A reports near-IR excess around ζ Aql A, and suggests that an unseen close companion is a more likely source than a close-in hot debris disk	Okab
λ	Aql	19 07.4 −4 51	3.43	−0.10	B9 Vnp (kB7HeA0) †	26	0.5	120	0.093	192	−12 V†	possibly SB ¶ rapid rotator (< 21h) ¶ suspected chemically anomalous (metals-weak, in λ Boo class); Astron. Alm. (epoch 2021.5) assigns MK type "A0 IVp (wk 4481)"	
τ	Sgr	19 08.3−27 38	3.32	1.17	K1.5 IIIb†	27	0.5	120	0.255†	191	+45†	possibly SB ¶ high velocity relative to Sun suggests origin outside galactic thin disk; underabundance of metals is consistent with this conjecture ¶ slow rotator (≤ 270 d)	
π	Sgr AB†	19 11.1−20 59	2.88†	0.38	F2 II–III + n.a.	6	−3.1	500	0.036	182	−10	tight triple system, seldom successfully resolved π Sgr A is mag. 3.6, π Sgr B is mag. 3.6, and π Sgr C is mag. 6.0; AB 0.1″→0.1″, PA 152°→179°, 1936→1989 (only 5 astrometry measurements), with A-to-B distance ≥ 13 AU, AB orbit ≥ 15 y; AB,C 0.4″→0.3″, PA 122°→136°, 1936→1939 (only 3 astrometry measurements), with AB-to-C distance ≥ 40 AU, orbit ≥ 100 y; under IAU rules, "Albaldah" designates π Sgr A, not π Sgr B or π Sgr C ¶ in HR-diagram terms, π Sgr A lies on blue edge of IS, without being presently observed to pulsate ¶ lunar occultations of ABC are possible, planetary occultations possible yet rare (next by Venus, 2035 Feb. 17)	Albaldah
δ	Dra A	19 12.6+67 42	3.07	0.99	G9 III	33.5	0.7	97	0.133	46	+25 V		Altais
δ	Aql Aa	19 26.6 +3 10	3.36†	0.32	F2 IV†	64	2.4	51	0.268	72	−30 SB†	fast rotator (> 0.9 d) ¶ Astron. Alm. (epoch 2021.5) assigns MK type F2 IV–V ¶ the SB which is δ Aql Aa,Ab has been observed just once in successful astrometry (speckle interferometry, 1979, with 1.9 m telescope at OHP, with angular separation found to be 132±10 mas); the Aa,Ab duplicity was discovered astrometrically, through periodic perturbation in proper motion of δ Aql Aa, in 1929 (Alden, as reported in 1936AJ.....45..193A); an astrometric-spectroscopic orbit is offered in 1989AJ.....98..686K, with period 3.426±0.006 y, *e*=0.36±0.07 (rotational broadening makes the spectroscopy difficult; also, this paper deduces a large magnitude difference between δ Aql Aa and δ Aql Ab, diminishing the prospects for easy interferometry)	

												¶ stars.astro.illinois.edu/sow/deltaaql.html discusses points of uncertainty (incl. the just-mentioned binarity, and possible δ Sct variability; although in 2018 δ Aql was not in the AAVSO(VSX) database, it is now (as viewed 2022 Feb. 22) listed as a slight variable, of the γ Dor type, 3.36–3.37 in V, 1.04524 d)	
β	Cyg Aa†	19 31.6+28 00	3.36	1.09	K3 II†	10†	−2.3	330†	0.009	229	−24 V	B: 4.68, 35″; Aa, Ab, Ac ≤ 0.3″ (2020) B is of MK type B9.5 Ve ¶ if AB is true binary, orbit is possibly ≥ 100 000 y; the competing mere-optical-companions thesis is argued by Bob King in *Sky & Telescope* 2016 Sep. 21; same conclusion is reached in 2018 by P. Plait at www.syfy.com/syfywire/long-standing-as tronomical-mystery-solved-albireo-is-not-a-binary-star, on strength of fresh *Gaia* data (which yield for β Cyg B **π** = 8.4 mas ±2%, implying distance for β Cyg B, to two significant figures, 390 ly; however, further analysis is needed, since astrometry of β Cyg A is potentially perturbed by the multiplicity of A (en.wikipedia.org/wiki/Albireo recaps literature, with some reference to recent interferometry)); WDS takes the firm view that AB is not a true binary, citing discrepancy in its accepted parallaxes, while also noting the binarity defence arguments of R.Griffin in 1999JRASC..93..208G (Griffin's assumed parallaxes are for their part not discrepant; in arguing for binarity, Griffin also *(a)* cites the statistical improbability of two such bright objects appearing on the celestial sphere at so tight an angular separation if they are not gravitationally bound, and *(b)* notes the possibility of shared proper motion (with, he stresses, proper motion observations rendered difficult by possible astrometric wobble in the photocentre of the Aa,Ac binary)) ¶ Aa,Ab,Ac are at mags. 3.4, 5.0, 5.2 respectively ¶ the Aa,Ac binary has been measured in astrometry 62 times, 1976→2021 (typical angular separation ~0.3″); 2008AN....329...54S offers a preliminary Aa,Ac orbit (with *e*~0.3; since period is ~210 y, the orbit will remain only imperfectly known over coming decades); the Aa,Ab binary has been measured in astrometry just twice (1978,1995, with the 1978 angular separation 0.1″) ¶ our values, for β Cyg A, of **π** = 10 mas (strictly, 9.5 mas ± 6.0%), with D consequently computed to two significant figures as 330 ly, are taken uncritically from *Gaia* ~2018, rather than (as in our previous Handbook editions) from *HIPPARCOS*; we do not here attempt a critical investigation of uncertainties ¶ β Cyg B is a fast rotator (< 0.6 d), and consistently with this is in emission (as "Be", rather than plain "B": being very evolved, this star is not, however, an instance of the "Be phenomenon" as discussed in the final subsection of our accompanying essay) ¶ β Cyg B is not now thought to be a binary (WDS note, and WDS "X" flag) ¶ the name "Albireo," colloquially associated with the AB pairing as visible in a small telescope, applies under IAU rules only to β Cyg Aa	Albireo
δ	Cyg A†	19 45.7+45 11	2.89†	0.00	B9.5 III	20	−0.7	160	~0.066	~42	−20 SB	B 6.3, F1 V; 1.9″→2.7″, PA 41°→213°, 1826→2020 orbit 780 y; separation 84 AU min, 230 AU max, 157 AU average, period 780 y ¶ δ Cyg A is a rapid rotator ¶ δ Cyg C is gravitationally bound to the AB pair: mag. 12, angular distance (2017) 62.5″, PA (for AC): 66°→67°, 1913→2017 ¶ variability has been suspected both in A and in B ¶ E(B–V)=+0.05	Fawaris

γ	Aql A	19 47.3+10 40	2.72†	1.51	K3 II†	~8.3	−2.7	390	0.017	100	−2 V	radius ~0.5 AU ¶ variability has been asserted ¶ a rare instance of a “hybrid” star (possessing a (hot) corona, like our Sun’s, and yet also emitting the cool high-mass wind typical in an evolved star)	Tarazed
α	Aql A	19 51.9 +8 56	0.76†	0.22	A7 Vnn	195	2.2	16.7	0.660	54	−26	rapid rotator (~7 h or ~8 h, latitude-dependent) the first MS star, other than the Sun, to yield a measurement of photospheric oblateness (2001ApJ…559.1155V); 2007Sci…317..352M announces CHARA imaging with angular resolution ~0.65 mas (the first direct imaging of any MS star other than the Sun; news.bbc.co.uk/2/hi/science/nature/6709345.stm is a news writeup; 2007Sci…317..342M shows oblate rotation-flattened photosphere, brighter at poles than at equator) ¶ found in 2005ApJ…619.1072B, via WIRE salvage mission, to be a slight variable of the δ Sct type, a classification now followed by AAVSO(VSX) (which indicates a fluctuation of 0.004 mag in V; WIRE makes this the brightest known δ Sct variable; second-brightest is β Cas); the 2005ApJ…619.1072B authors suggest that many δ Sct variables, as residents of the IS, may be oscillating at such low amplitudes as to evade detection except by such sensitive facilities as WIRE (their suggestion helps relieve a longstanding astrophysical puzzlement over IS residents that appear, inexplicably, not to be pulsating) ¶ drawing on interferometry, spectroscopy, and the 2005ApJ…619.1072B δ Sct asteroseismology, 2020A&A…633A..78B, while conceding a failure of uniqueness, and consequently conceding the need for further spectroscopy, offers a physical model that takes account of the rapid rotation (by assuming mere cylindrical symmetry, and not the outright spherical symmetry that would be appropriate in the modelling of a slow rotator); Table 5 of the paper summarizes its results, comparing them against earlier modelling; the paper finds a typical core rotation period ~0.6 of the rotation period of the photosphere, and with only modest latitude variation (shearing) in the rotation period in the photosphere (with middle altitudes ~7.7 h, equator ~7.8 h, immediate vicinity of poles ~8.1 h); the paper deduces a value for core metallicity that makes α Aql A young, aged only ~100 My (but some other recent literature proposes instead ~1.2 Gy; both suggested ages are consistent with the failure of α Aql A to have progressed significantly off the MS); the paper ascribes to α Aql A a remarkable variation in envelope temperature, with the envelope convective (because cooler) at low latitudes and radiative (because hotter) at high latitudes (a similar latitude-governed bifurcation in envelope characteristics is believed present in the rapid rotator α Cep A (Alderamin)); consistently with this latitude-dependent temperature variation, 2009A&A…497..511R finds modest coronal X-ray emission, attributed to modest dynamo activity at the low or intermediate latitudes (the authors note that of the stars not in a tight binary system, α Aql A is among the hottest known to have coronal X-ray emission) ¶ 2017A&A…608A.113N reports time-varying IR (K-band) excess, suggestive of tenuous circumstellar material (possibly debris disk: the “Be phenomenon”, present in many hot, young rapid rotators, is believed to involve a gas disk rather than a debris disk) ¶ since proper motions of α Aql A and α Aql B are discrepant, AB is not a true binary	**Altair**

													(AB astrometry: 143″→196″, PA 335°→286°, 1781→2015); under IAU rules, the name “Altair” designates just α Aql A	
η	Aql A†	19 53.6	+1 04	3.87v†	0.63	F6–G1 Ib	2	−4.3	1000	0.011	140	−15 SB	Cepheid variable, 3.49–4.30 in V, 7.2 d more precisely, AAVSO(VSX) as viewed 2022 March 3 gives 7.17679 d (same value as given 2021 Jan. 28 and in 2019 January); BSC5 asserts 7.176641 d with period changes; 2002ApJS..140..465B (in centre panel of the author’s Fig 1) gives (1990s?) photometry (to tighter than ±10 millimag), colour, and radial-velocity curves ¶ η Aql AB has been split with HST WFC3 (cf 2020ApJ...905...81E; A-to-B distance is ~200 AU, AB period ~900 y); the AB astrometry so far comprises just 2 observations (0.7″→0.7″, 2011→2012); A is for its part an SB not as yet resolved, for which 2020ApJ...905...81E gives period 4 y, suggesting that the orbit is eccentric and is seen nearly face-on (i.e. is seen in the orientation least favourable for the radial-velocity investigation required in orbit spectroscopy) ¶ in the case of novice Northern Hemisphere observers troubled by frequent cloud, its rather long period makes η Aql A a better high-amplitude Cepheid demonstration than the more celebrated δ Cep A **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	
γ	Sge	19 59.8	+19 33	3.51†	1.57	M0 III†	13	−1.0	260	0.070	71	−33 V?	radius 0.26 AU (from interferometry; the disk subtends an angle of 6.18 mas) ¶ Astron. Alm. (epoch 2021.5) assigns MK type M0$^-$ III ¶ slightly variable; already has a dead carbon core, is not yet a Mira **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	
θ	Aql Aa	20 12.5	−0 45	3.24	−0.07	B9.5 III†	11	−1.5	290	0.036	81	−27 SB2†	a good marker of celestial equator ¶ there may be some unclarity in the literature regarding “Aa”, “Ab” nomenclature (cf WDS Note): it is at any rate clear that the SB2 which is θ Aql A has now been interferometrically resolved, and that two quite similar orbital solutions have been published (period 17.124 d, *e*≈0.6); recent discussions include 1995AJ....110..376H, 2000A&AS..145..215P ¶ θ Aql A pairing is metal-rich ¶ Astron. Alm. (epoch 2021.5) assigns MK type B9.5 III$^+$	
β	Cap Aa†	20 22.3	−14 43	3.05	0.79	K0: II:†	10	−2.0	300	0.046	81	−19 SB	hierarchical quintuplet (or greater) en.wikipedia.org/wiki/Beta_Capricorni has a diagram summarizing the known gravitationally bound hierarchy: Aa, Ab1 (seen), Ab2 (unseen), Ba, Bb, where Aa is mag. 3.1, Ab1Ab2 is mag. 4.9, Ba is mag. 6.2, Bb is mag. 9.1 (but Wikipedia needs a caveat: since Ab is not yet resolved, even in interferometry, the designations “Ab1”, “Ab2” are not yet WDS-conformant); WDS also lists, as nearby in angular distance, C (mag. 8.8, 226″), D (mag. 13.0, 116″), and E (mag. 14.4, 3.9 ″ from D): Ab1, Ab2 period is 8.7 d; Aa experiences Ab1,Ab2 as essentially a point mass, recently at angular separation 50 mas (period 3.77 y, inter-component distance ~4 AU); Ba, Bb 0.5″, according to WDS (and yet en.wikipedia.org/wiki/Beta_Capricorni states 3″), PA 106°→54°, 1884→2019; AB 205″, PA 268°→267°, 1800→2012; each of (Aa,Ab), (Ba,Bb) experiences the other as essentially a point mass, at separation ≥ 0.34 ly, with the (Aa,Ab)+(Ba,Bb) orbit ≥ 700,000 y; orbital solutions have been published for Aa,Ab and for Ba,Bb, but not for the wide (Aa,Ab)+(Ba,Bb) pairing ¶ spectral type of β Cap A is controverted;	Dabih

												entire system appears in spectrograph as K0: II: + A5: V:n ¶ β Cap A is overabundant in Hg, Mn, and several other heavy elements ¶ lunar occultations are possible, planetary occultations possible-yet-rare	
γ	Cyg A†	20 23.0+40 20	2.23†	0.67	F8 Ib†	2	−6.5	2000	0.003	111	−8 V	unusual supergiant, because of MK type F (among supergiants, it is the hotter and the cooler types that are more usually encountered; γ Cyg A resides near the HR diagram IS: 2010AJ....140.1329G first surveys the observational literature, then discusses spectral variations (possibly pulsation-style oscillation, or alternatively large convection cells are possible; and indeed convection cells can be a driver of oscillation)) ¶ radius ~1 AU (stars.astro.illinois.edu/sow/sadr.html discusses uncertainty) ¶ BSC5: "no demonstrable connection" between γ Cyg and the so-called γ Cyg supernova remnant	Sadr
α	Pav A	20 27.4−56 40	1.94	−0.12	B2.5 V	18	−1.8	180	0.086	175	+2 SB†	SB, not as yet resolved (so WDS is not as yet able to write "α Pav Aa", "α Pav Ab"), 11.753 d, inter-component distance 0.21 AU ¶ the SB which is α Pav A has celestial-sphere neighbours α Pav B (mag. 9.1), α Pav C (mag. 9.7), α Pav D (mag. 9.7), at rather wide angular separations from α Pav A, with rather scant astrometry coverage (AB 245″→249″, PA 85°→80°, 1879→2008, with just 7 measurements; AC 226″→244″, PA 80°→77°, 1879→2010, with just 3 measurements; AD 59″→62″, PA 249°→254°, 1904→2010, with just 3 measurements; BC 18″→17″, PA 332°→332°, 1835→2010, with just 10 measurements) ¶ 1988A&A...201..273V discusses galactic-astronomy implications of puzzling deuterium paucity in α Pav A ¶ E(B–V)=+0.02 ¶ the name, although anomalously English, is nevertheless IAU-official: its origins lie in 1930s RAF Air Almanac project, which directed HM Nautical Almanac Office that no air-navigation star was to be left nameless	**Peacock†**
α	Ind A	20 39.1−47 13	3.11	1.00	K0 III CN–1†	33	0.7	98	0.083	37	−1	Fe overabundant (α Ind AB born in metal-rich ISM cloud?)	
α	Cyg A	20 42.2+45 22	1.25†	0.09†	A2 Ia	2†	−6.9	~1400†	0.003	47	−5	blue supergiant, of radius ~0.5 AU or ~1 AU for context pertaining to this particular BSG in the general population of hypergiants and supergiants, cf en.wikipedia.org/wiki/List_of_largest_stars (which adopts "~1 AU"); for current state of theoretical investigations into BSG populations (crossing Hertsprung-Russell diagram for the first time, redward? or, rather, after episode of mass loss, crossing for the second time, blueward?) cf, e.g. 2014MNRAS.439L...6G ¶ slightly variable, and the prototype of the α Cyg variables: AAVSO(VSX) gives V range 1.21–1.29; seemingly irregular (in the α Cyg variables, many short-period oscillations are superimposed); 2011AJ....141...17R discusses α Cyg A, reporting a 1977-through-2001 campaign in both photometric and spectroscopic variability ¶ α Cyg A core hydrogen-fusion career started in MK spectral type B, or possibly even in the rare MK spectral type O ¶ present mass-loss rate is ~8e-7 $M_{\odot}$/y ¶ slow rotator (period possibly as long as 0.5 y,	**Deneb**

													consistently with its large radius and its ongoing mass loss) ¶ public-outreach astro audiences enjoy comparing and contrasting distance, and therefore intrinsic luminosity, of α Cyg A with distance, and therefore intrinsic luminosity, of the other two Summer Triangle stars (nearby α Lyr A, nearby α Aql A; all three are similar not only in their apparent magnitudes, but also in falling within MK type A, and consequently in lacking tint, even through binoculars); it is perhaps worth stressing in such lectures that the α Cyg A distance, although large (1500 ly? more?), is nevertheless not yet well known; Kaler in stars.astro.illinois.edu/sow/deneb.html, accepting ~1500 ly, writes that if placed at distance of α Lyr A, α Cyg A "would /…/ be as bright as a well-developed crescent Moon, cast shadows on the ground, and easily be visible in broad daylight"
η	Cep A	20 45.7+61 56	3.41	0.91	K0 IV†		70.1	2.6	46.5	0.823†	6	−87†	high velocity relative to Sun indicates interloper status in galactic thin disk (and observed underabundance of Fe is consistent with interloper status)
β	Pav	20 47.0−66 07	3.42	0.16	A6 IV†		~24.1	0.3	135	0.044	283	+10	still a fast rotator (≤ 2.3 d), although core hydrogen fusion is ended or is close to ending ¶ Astron. Alm. (epoch 2021.5) assigns MK type A6 IV⁻
ε	Cyg Aa†	20 47.1+34 03	2.48	1.02	K0 III		44.9	0.7	73	0.486†	47	−11† SB†	Aljanah the SB which is ε Cyg Aa,Ab (period ≥ 15 y, only one set of lines visible) has been interferometrically measured-and-resolved just twice, in the period 1983→1991 ¶ velocity of Aa,Ab (and of faint, outlying, gravitationally bound red dwarf C, mag. 13.4) relative to Sun is high
ζ	Cyg Aa†	21 13.9+30 19	3.21	0.99	G8 IIIa Ba†	0.5	23	0.0	140	0.069	175	+17 SB	in evolutionary terms, possibly a Red Clump resident (stable helium fusion in core); but it might also be the case that core-helium fusion has yet to begin ¶ SB is resolved, as ζ Cyg Aa, ζ Cyg Ab: 1920Obs...112..168G discusses spectroscopy, reviewing the history at a level of detail so instructive as to make this a case study for spectroscopy technique generally, even outside the particular domain of the ζ Cyg system; an orbital solution for Aa,Ab has been published, asserting period 6489 d (~18 y), with *e*=0.22; on this solution inter-component distance is 8 AU min, 13 AU max, 11 AU average, and angular length of semimajor axis of the rectified orbit is ~190 mas; Ab is a WD, of mag. 13.2; the Aa,Ab binary was first split with far-UV imaging from *IUE* (in general, UV is a desirable regime for observing binaries with an elusive WD secondary, since WDs, although faint in the V band, are UV-bright); additionally, 2001MNRAS.322..891B announces direct imaging with HST WFPC2 (elongated smear, WD partly resolved, possibly 36 mas; but this observation was made under unfavourable conditions, near periastron); as viewed 2022 Jan. 14, WDS documents just one successful astrometric data point for Aa,Ab, from the year 2000; the ζ Cyg Aa,Ab SB binary, long observed spectroscopically, and now open to space-based UV astrometry, is of interest for WD studies, since determination of orbit yields determination of WD mass (admittedly, a very small number of WD orbits, notably including the respective WD companions of two stars in this Handbook table, Sirius (in α CMa) and Procyon (in α CMi), have been determined even from terrestrial astrometry; with the advent of space-based UV astrometry, the overall WD mass-determination

												situation, historically something of a bottleneck, may now be expected to improve) ¶ ζ Cyg Aa is chemically a mild barium star (Astron. Alm. (epoch 2021.5) assigns MK type "$G8^{+}$ III–IIIa Ba 0.5"); before becoming a WD, ζ Cyg Ab, as a mass-shedding AGB star, deposited barium onto ζ Cyg Aa	
α	Cep A	21 19.1+62 41	2.45†	0.26	A7 Van†	66.5	1.6	49.1	0.158	72	−10 V	fast rotator (< 12h); the rotational shape distortion, into an oblate spheroid, gives α Cep A a remarkable variation in envelope temperature, with the envelope convective (photosphere ~6600 K) at equator and radiative (photosphere ~8600K) at poles (the transition temperature is ~8300 K): a similar latitude-governed bifurcation in envelope characteristics is present (cf 2011ApJ…732…68C Fig. 9) in the rapid rotator α Aql A (Altair) ¶ Astron. Alm. (epoch 2021.5) assigns MK type A7 $V^{+}n$ ¶ listed by AAVSO(VSX) as δ Sct variable, with V mag. range 2.41–2.47 ¶ several factors, including X-ray emission (consistent with corona, as might be expected for convection-harbouring latitudes of the envelope) indicate magnetic activity	Alderamin
β	Cep Aa†	21 28.9+70 40	3.23v†	−0.20	B1 III†	5	−3.4	700	0.015	56	−8 SB	variable: 3.16–3.27 in V, 0.19 d; B: 8.6; 13.5″ (2016) PA 255°→251°, 1779 →2016 ¶ the archetype of the β Cep variables (although stars of this same type are sometimes called the "β CMa variables"), and (as is typical for the type) known to be multiperiodic: AAVSO supplies a 2010 Apr. 13 backgrounder at www.aavso.org/vsots_betacep; AAVSO(VSX) as viewed 2021 Jan. 16 asserts period 0.1904881 d; AAVSO archives a notice for an August 2009 β Cep campaign (coordinated photometry, spectroscopy, CHARA) at www.aavso.org/aavso-special-notice-162 ¶ β Cep Aa is a magnetic star ¶ system comprises at least (the much-studied variable) Aa and Ab (mag. 6.6, probably a Be-phenomenon star, and the origin of the Be-phenomenon behaviour observed in AaAb); Aa,Ab period is variously suggested as 50 y, 85 y; astrometry is now quite good, with 62 measurements over the period 1971→2007 (angular separation 0.3″→0.2″); if β Cep B is gravitationally bound to the Aa,Ab SB (no AB orbital solution has been published), then period is ≥ 40,000 y, with (Aa,Ab)-to-B distance 3,000 AU ¶ MK luminosity class III ("giant") notwithstanding, β Cep Aa is still fusing hydrogen in its core	Alfirk
β	Aqr A	21 32.7 −5 28	2.9†	0.83†	G0 Ib†	6	−3.2	500	0.020	114	+7 V?	a rare instance of a yellow supergiant β Aqr A is possibly now evolving blueward in a second crossing of the HR diagram ¶ spectroscopically a "hybrid" star, combining signature of hot corona with signature of cool massive wind; 2005ApJ...627L..53A, in a study jointly covering β Aqr A and the astrophysically similar hypergiant (likewise a hybrid) α Aqr A, reports *Chandra* observation of coronal X-rays (first X-ray detection from a hybrid G supergiant; such supergiants are X-ray deficient, their coronae notwithstanding) ¶ β Aqr lies in the IS on the HR diagram, and yet is not known to be a pulsator ¶ their ~10° separation on the celestial sphere notwithstanding, β Aqr A and α Aqr A have shared proper motion and similar parallaxes (and WDS β Aqr A is the same object as WDS α Aqr C; this pairing of β Aqr A a.k.a.	Sadalsuud

												α Aqr C with α Aqr A is further discussed in bestdoubles.wordpress.com/2014/12/15/the-alpha-beta-gamma-of-aquarius-%CE%B1-%CE%B2-and-%CE%B3-aquarii/)	
ε	Peg A	21 45.3 +9 59	2.38†	1.52†	K2 Ib†	5	−4.2	700	0.027	89	+5 V	slight irregular var.: 2.37–2.45 in V (flare in 1972) 1972IAUC.2392....1W reports extreme flare-like brightening, ~10 minutes, to V mag. 0.7 ¶ orange-class supergiant ¶ Astron. Alm. (epoch 2021.5) assigns MK type K2 Ib–II ¶ 1987MNRAS.226..563S discusses abundances, finding that, earlier literature notwithstanding, ε Peg A is unremarkable in its Ba (and unremarkable in its Sr), and therefore discounting an earlier suggestion that ε Peg A outer layers have hosted nucleosynthesis in slow-neutron capture ¶ BSC5 suggests “cooler shell surrounding” ¶ WDS documents as celestial-sphere neighbours the faint ε Peg B (mag. 12.8, known to be not gravitationally bound to ε Peg A (angular separation 83″ in 2013), and additionally the less faint ε Peg C (mag. 8.7; AC astrometry is 138″→144″, PA 323°→318°, 1825→2018)	**Enif**
δ	Cap A	21 48.3−16 01	2.85v†	0.18	A3mF2 IV:†	84	2.5	38.7	0.396	139	−6 SB†	eclipsing binary: V 2.81–3.05, 1.0 d, 3.2 + 5.2 since SB has not been measured as a visual binary (not even interferometrically; the binarity has, admittedly, been demonstrated in at least one occultation: lunar occultations are possible, planetary occultations possible-yet-rare), WDS is not as yet able to write “δ Cap Aa”, “δ Cap Ab”; the δ Cap A pair is classified at AAVSO(VSX) as an Algol-type eclipsing binary, 1.0227688 d in AAVSO(VSX) as viewed 2021 Jan. 16, but 1.0227672 d as viewed 2022 Feb. 23; AAVSO(VSX) also yields O-C, i.e. period-monitoring, plotting from 2016); secondary is ~3 mag. fainter than primary and is judged in 1992MNRAS.259..251W to be mildly active, possibly tidally locked, with large spot; A is known to be SB since 1906 (Slipher), and yet is known to be eclipsing only as of 1956PASP...68..541E ¶ 1994MNRAS.266L..13L rebuts earlier assertion of δ Sct variability, and remarks that “given the brightness of the system, δ Cap is poorly observed,” with period awkward for any one solitary observatory (an implication of this remark is that coordinated intercontinental photometry would now be helpful) ¶ Astron. Alm. (epoch 2021.5) assigns MK temperature type F2m and does not assign an MK luminosity class	Deneb Algedi
γ	Gru	21 55.3−37 15	3.00	−0.08	B8 IV–Vs	15	−1.1	210	0.099	98	−2 V?		Aldhanab
α	Aqr A	22 06.9 −0 13	2.95†	0.97†	G2 Ib†	6	−3.1	~520	0.021	117	+8 V?	a good marker of celestial equator a rare instance of a yellow supergiant; possibly now evolving blueward in a second crossing of the HR plane; resides in the IS (under at least one definition of IS) and yet is nonpulsating (cf further 2017AstL...43..265U) ¶ spectroscopically a “hybrid star,” combining signature of hot corona with signature of cool, massive wind; 2005ApJ...627L..53A, in a study jointly covering α Aqr A and the astrophysically similar supergiant (likewise a hybrid star) β Aqr A reports *Chandra* observation of coronal X-rays (first X-ray detection from a hybrid G supergiant; such supergiants are X-ray deficient, their coronae notwithstanding) ¶ despite their ~10° separation on the celestial sphere, α Aqr A and β Aqr A have shared proper motion and similar parallaxes (and WDS β Aqr A is the same object as WDS α Aqr C; this pairing of β Aqr A a.k.a. α Aqr C with α Aqr A is further discussed in bestdoubles.wordpress.com/2014/12/15/the-al	Sadalmelik

												pha-beta-gamma-of-aquarius-%CE%B1-%CE%B2-and-%CE%B3-aquarii/)
α	Gru A	22 09.6−46 51	1.73	−0.07	B7 Vn	32	−0.7	101	0.194	139	+12	**Alnair** rapid rotator (< 1d) ¶ E(B–V)=–0.02
θ	Peg	22 11.3 +6 19	3.52†	0.09	A2mA1 IV–V†	35	1.3	90	0.284	84	−6 SB2	Biham rapid rotator (< 20 h); consistently with rapid rotation, and therefore with a stirred atmosphere, elemental abundances are unremarkable ¶ earlier assertion of δ Sct variability is now discounted
ζ	Cep	22 11.6+58 19	3.39†	1.56†	K1.5 Ib†	3.9	−3.7	800	0.014	69	−18 SB	**[THIS STAR ONLY IN ONLINE VERSION OF TABLE]** orange supergiant either approaching core-helium fusion or already in core-helium fusion ¶ an eclipsing companion has been suggested, with suggestion later questioned ¶ metals somewhat overabundant
α	Tuc	22 20.0−60 09	2.87	1.39	K3 III†	16	−1.1	200	0.081	241	+42 SB†	SB 11.5 y, separation possibly 11.5 AU ¶ primary in the SB is a giant, with carbon underabundant, nitrogen overabundant ¶ stars.astro.illinois.edu/sow/alphatuc.html discusses uncertainties in the evolutionary stage of this giant, offering three scenarios
δ	Cep A†	22 30.0+58 32	4.07v†	0.78	F5–G2 Ib	4†	−3.0	900†	0.016	77	−15 SB†	the prototype Cepheid variable: 3.49–4.36 in V, 5.4 d second-nearest Cepheid (α UMi is still nearer) ¶ AAVSO offers a tutorial at www.eso.org/public/outreach/eduoff/aol/market/collaboration/varstar/pg2.html and an initial backgrounder at www.aavso.org/vsots_delcep; the first three sections of a paper directed inter alia to AAVSO observers, 2016JAVSO..44..179N, constitute a deeper backgrounder on the Cepheids ¶ AAVSO(VSX) has, as viewed 2021 Jan. 28 and again at 2022 Feb. 24, period 5.366266 d (is this value possibly now stale?); although Cepheids experience both period jitter and (monotonic) period slide, with a slide of even 200 s/y possible, 2014ApJ...794...80E finds δ Cep period sliding slowly, at just −0.1 s/y (period decrease-increase is a signature of evolution, specifically of density increase-decrease, as a Cepheid passes across the HR diagram (δ Cep is now making its second such passage, moving blueward)) ¶ 2015ApJ...804..144A announces that δ Cep A is SB, with period 2201 d ¶ accurate distances to Cepheids are foundational in cosmology, which needs independently known (galactic) Cepheid distances before embarking on its external-galaxy distance deductions through applications of the Cepheid Period-Luminosity (PL) Law; it is reassuring that the 2007 *HIPPARCOS* distance and the distance implied by the usual PL calculation agree to within uncertainties; although we have here stated the 2007 *HIPPARCOS* parallax, on which distance of δ Cep depends, as 4 mas, our cited 2007 *HIPPARCOS* determination is more formally, with decimal fractions and the uncertainty made explicit, 3.77±0.16 mas; 2015ApJ...804..144A proposes instead 4.09±0.16 mas, with the remark that impending *Gaia* may be expected, in part in the light of these authors' SB announcement, to secure an authoritative parallax; an already reassuring state of affairs may thus be expected to improve further ¶ mass loss ~1e–6 $M_{\odot}$/y; bow shock in ISM has now been detected ¶ δ Cep C, at mag. 6.1, is in slow and wide orbit

												with δ Cep A (no orbital solution published; period 345,000 y; AC astrometry is 42″→41″, PA 195°→191°, 1800→2018) **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	
ζ	Peg A	22 42.6+10 57	3.41†	−0.09	B8.5 III†	16	−0.6	210	0.078	98	+7 V?	our (Garrison) MK type notwithstanding, B8 V has been suggested ¶ fast rotator (< 1.4 d) ¶ microvariable (2007PASP..119..483G discusses satellite detection of amplitude ~0.5 millimag); assigned by AAVSO(VSX) to the class of “slowly pulsating B stars”	Homam
β	Gru	22 44.0−46 46	2.07v†	1.61†	M5 III†	18	−1.6	180	0.135	92	+2	semiregular variable, 1.90–2.3 in V, 37d among the rather uncommon cool red giants, with radius slightly > 0.8 AU ¶ Astron. Alm. (epoch 2021.5) assigns MK type M4.5 III ¶ classified at AAVSO(VSX) as semiregular late-type giant, perhaps on the basis of 2006JAVSO..34..156O (this paper might serve as a case study for effective amateur-budget intercontinental photometry collaboration)	Tiaki
η	Peg Aa†	22 44.1+30 20	2.93†	0.85	G8 II−III + F0 IV	15	−1.2	210	0.029	153	+4 SB	compos. spectrum, SB period 813 d the η Peg Aa,Ab SB is resolved in speckle interferometry, with orbital solution published (66 measurements: 0.1″→0.1″, 1975→2005); BC is itself a tight pairing of equally bright stars, probably a true binary (0.3″→0.2″, 1889→2011), probably in orbit with the Aa,Ab binary, making this a hierarchically organized quadruple system; A,BC astrometry is 90″→92″, PA 339°→338°, 1824→2012 ¶ we take the MK type from WDS Note ¶ we give mag. of η Peg as the combined light of η Peg Aa, η Peg Ab; it is known that individually, in the R band as distinct from the V band, η Peg Aa and η Peg Ab are respectively mag. 4.1, mag. 6.9 ¶ η Peg Aa is a slow rotator (818 d?)	Matar
ε	Gru	22 49.9−51 12	3.49	0.08	A2 Va	25	0.5	130	0.126	121	0 V	rapid rotator (< 0.65 d) **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	
ι	Cep	22 50.5 +66 19	3.50	1.05	K0 III†	28.3	0.8	115	0.141	208	−12	Astron. Alm. (epoch 2021.5) assigns MK type K0^{-} III **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	
μ	Peg	22 51.1 +24 43	3.51	0.93	G8 III†	31	0.9	106	0.151	106	+1	Astron. Alm. (epoch 2021.5) assigns MK type G8^{+} III **[THIS STAR ONLY IN ONLINE VERSION OF TABLE]**	Sadalbari
δ	Aqr	22 55.8 −15 42	3.27	0.07	A3 IV–V	20	−0.2	160	0.051	237	+18 V	weak λ4481 ¶ rapid rotator (< 3.0 d)	Skat
α	PsA Aa†	22 58.9−29 30†	1.17	0.14	A3 Va†	130	1.7	25.1	0.368	1 17	+7	2008 (HST) image was debris cloud, not exoplnt HST putative 2008 “exoplanet” α PsA Ab was IAU-named Dagon, after a Semitic deity; at ~125 AU, in the outermost of the debris rings; Dagon was in always-wide (albeit eccentric) orbit, making direct imaging, as opposed both to spectroscopy (for star Doppler wobble) and astrometry (for star transverse wobble) the tool of choice: 32 AU min, 320 AU max; period ~1700 y; in more recent years, it was suggested that Dagon could be a mere dust cloud, or an aggregation of rubble, or a single rocky body; an explanation was needed for the fact that Dagon proved so readily HST-visible (e.g. visibility enhanced by circumplanetary dust sphere, or by circumplanetary ring system?); Dagon mass was uncertain (< 2× Jupiter, perhaps even ~Earth); but with Dagon now no longer HST-visible, it would appear that the 2008 HST image was of an expanding debris cloud, now become too tenuous for detection ¶ the nested circumstellar dust rings extend as far as radius ~150 AU (a distance recalling the Solar System Kuiper Belt); 2017ApJ...842....8M reports complete outer debris-ring mapping,	**Fomalhaut**

												via ALMA (223 GHz radio), finding ring mass of 0.015 Earths, eccentric, with α PsA Aa offset from the ring centroid ¶ α PsA Aa is a fast rotator (< 1d) ¶ in evolutionary terms, α PsA Aa is sufficiently young to be undergoing an analogue of the Solar System's Late Heavy Bombardment (and consistently with this, 2017ApJ...842....9M writes that exocometary gas is detected, in ALMA 230 GHz radio) ¶ 2017ApJ...842....8M comments that "given its unique characteristics and architecture, the Fomalhaut system is a Rosetta stone for understanding the interaction between planetary systems and debris disks" ¶ α PsA Aa has low metallicity ¶ 2013AJ....146..154M, working both from proper motion (across the celestial sphere) and from velocities along the line of sight, concludes that α PsA Aa, α PsA B, and α PsA C belong to the same system: B (a flare star) is V mag. 7, at angular separation almost 2° from Aa (period ≥ 7.6 My), while C is V mag. 13, at enormous angular separation 5.7° from Aa (and yet at a sufficiently low distance from Aa-with-B to have the Aa-with-B gravitational field dominate the general external gravitational field at its location; period is ≥ 35 My) ¶ β Peg, α Peg serve as pointers: since α PsA Aa lies a couple of arcminutes N of DEC=−30°, α PsA Aa rises (if briefly) above the horizon even for such Canadian subarctic communities as Churchill, and for such Scandinavian communities as Stavanger	
β	Peg A	23 04.9+28 12	2.44v†	1.66	M2 II–III†	16.6	−1.5	~196	0.232	54	+9 V	semiregular variable, mag. 2.31–2.74 in V, 43.3 d with period 43.3 d ¶ Astron. Alm. (epoch 2021.5) assigns MK type M2.5 II–III ¶ an intermediary between straightforward red giant and red bright giant (radius ~0.5 AU); mass-loss rate is notably low for such a star (≤ 1e-8 $M_\odot$/y; i.e. ~100× lower than mass loss rate of α Ori Aa (Betelgeuse); *IRAS* detected no IR excess)	Scheat
α	Peg	23 05.9+15 20	2.49	0.00	A0 III–IV	24	−0.6	133	0.073	124	−4 SB	rapid rotator (1.5 d)	**Markab**
γ	Cep A†+1P†	23 40.3+77 45	3.21	1.03	K1 III–IV†	71	2.5	46	0.135	339	−42 V?	the binary γ Cep A, γ Cep B has been split with adaptive optics at Subaru, with direct imaging reported in 2007A&A...462..777N (AB astrometry: 3 measurements, 0.9″→0. 9″, PA 257°→256°, 2006→2006); the IAU-official name "Errai" applies to γ Cep A rather than to the two-star system γ Cep AB; Errai hosts an exoplanet, IAU-named Tadmor, which is not as yet astrometrically observed by any optical technique (so WDS is not as yet able to write "γ Cep Aa", "γ Cep Ab"); Tadmor, circumstellar without being circumbinary, is among the few exoplanets discovered in a two-star system; Tadmor orbital period is 2.47 y, with average distance from Errai 2.05 AU, and with mass between 3× Jupiter and 16× Jupiter; A-to-B distance is 12 AU min, 25 AU max, AB orbital period (orbital solution has been published) is 66 y or 67 y ¶ Errai rotation period is possibly 781 d (making this star a slow rotator) ¶ Astron. Alm. (epoch 2021.5) assigns MK type K1 III−IV CN 1	Errai